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The structure and origin of magnetic clouds in the solar wind

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Abstract. Plasma and magnetic field data from the Helios 1/2 spacecraft have been used to investigate the structure of magnetic clouds (MCs) in the inner heliosphere. 46 MCs were identified in the Helios data for the period 1974–1981 between 0.3 and 1 AU. 85% of the MCs were associated with fast-forward interplanetary shock waves, supporting the close association between MCs and SMEs (solar mass ejections). Seven MCs were identified as direct consequences of Helios-directed SMEs, and the passage of MCs agreed with that of interplanetary plasma clouds (IPCs) identified as white-light brightness enhancements in the Helios photometer data. The total (plasma and magnetic field) pressure in MCs was higher and the plasma- β lower than in the surrounding solar wind. Minimum variance analysis (MVA) showed that MCs can best be described as large-scale quasi-cylindrical magnetic flux tubes. The axes of the flux tubes usually had a small inclination to the ecliptic plane, with their azimuthal direction close to the east-west direction. The large-scale flux tube model for MCs was validated by the analysis of multi-spacecraft observations. MCs were observed over a range of up to $\sim 60^\circ$ in solar longitude in the ecliptic having the same magnetic configuration. The Helios observations further showed that over-expansion is a common feature of MCs. From a combined study of Helios, Voyager and IMP data we found that the radial diameter of MCs increases between 0.3 and 4.2 AU proportional to the distance, R , from the Sun as $R^{0.8}$ (R in AU). The density decrease inside MCs was found to be proportional to $R^{-2.4}$, thus being stronger compared to the average solar wind. Four different magnetic configurations, as expected from the flux-tube concept, for MCs have been observed in situ by the Helios probes. MCs with left- and right-handed magnetic helicity occurred with about equal frequencies during 1974–1981, but surprisingly, the majority (74%) of the MCs had a south to north (SN) rotation of the magnetic field vector relative to the

ecliptic. In contrast, an investigation of solar wind data obtained near Earth's orbit during 1984–1991 showed a preference for NS-clouds. A direct correlation was found between MCs and large quiescent filament disappearances (disparition brusques, DBs). The magnetic configurations of the filaments, as inferred from the orientation of the prominence axis, the polarity of the overlying field lines and the hemispheric helicity pattern observed for filaments, agreed well with the in situ observed magnetic structure of the associated MCs. The results support the model of MCs as large-scale expanding quasi-cylindrical magnetic flux tubes in the solar wind, most likely caused by SMEs associated with eruptions of large quiescent filaments. We suggest that the hemispheric dependence of the magnetic helicity structure observed for solar filaments can explain the preferred orientation of MCs in interplanetary space as well as their solar cycle behavior. However, the white-light features of SMEs and the measured volumes of their interplanetary counterparts suggest that MCs may not simply be just H α -prominences, but that SMEs likely convect large-scale coronal loops overlying the prominence axis out of the solar atmosphere.

1 Introduction

The term magnetic cloud (MC) was introduced by Burlaga *et al.* (1981) to characterize the magnetic field and plasma signatures of an interplanetary post-shock flow observed by five spacecraft separated over more than 30° in solar longitude between 0.9 and 2 AU. The outstanding feature of this solar wind transient was the smooth rotation of the magnetic field vector nearly parallel to a plane over a time interval of the order of one day at 1 AU. This peculiar transient flow showed extremely low plasma- β values, i.e., the plasma's magnetic pressure ($B^2/8\pi$) dominated the thermal pressure (nkT) due to its higher than average solar wind magnetic

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field strength (B) and its unusually low plasma temperature (T) and density (N). The MC maintained its characteristic magnetic structure during its outward propagation up to at least 2 AU, although it had considerably expanded in its radial direction.

Subsequent analyses of solar wind data from various spacecraft near Earth and in the outer heliosphere revealed that MCs are relatively common phenomena in the solar wind (Klein and Burlaga, 1982; Burlaga and Behannon, 1982; Burlaga, 1991). MCs are often, but not always associated with interplanetary shocks (Klein and Burlaga, 1982; Zhang and Burlaga, 1988). MCs that strike the Earth may cause intense geomagnetic storms (Wilson, 1987; Zhang and Burlaga, 1988; Tsurutani *et al.*, 1992; Bothmer, 1993; Bothmer and Schwenn, 1995).

The search for the solar origin of MCs brought up a close association of MCs with solar mass ejections (SMEs) and erupting prominences (disappearing filaments) (Burlaga *et al.*, 1982; Wilson and Hildner, 1984, 1986; Marubashi, 1986; Bothmer and Schwenn, 1994). Note, that in this study we use the more general term ‘solar mass ejection’ (SME), rather than ‘coronal mass ejection’ (CME), because the origin and evolution of transient events in the solar atmosphere is not yet known from current white-light coronagraph observations and in situ solar wind measurements. Further, evidence has been found that at times not only coronal but also chromospheric solar material is ejected into interplanetary space (Schwenn *et al.*, 1980). Here we prefer to reserve the term SME for the solar event and we call the interplanetary counterparts of SMEs ejecta, interplanetary plasma clouds (IPCs), or magnetic clouds (MCs) if they exhibit a specific magnetic topology. According to previous studies MCs represent just a subset of all IPCs found in the solar wind (Gosling, 1990; Bothmer and Schwenn, 1996). Among all IPCs identified by bidirectional electron (BDE) flows, $\sim 1/3$ had a large coherent internal magnetic field rotation characteristic of a MC (Gosling, 1990, 1993). Bothmer and Schwenn (1996) found from correlated white-light and in situ solar wind measurements that 41% of the driver gases (ejecta, IPCs), caused by SMEs directed towards the Helios spacecraft, were MCs.

Goldstein (1983) first suggested that the specific magnetic properties of MCs might be explained through force-free, large-scale, cylindrical magnetic flux tubes with $\nabla \times B = \alpha B$, i.e. the electric currents are flowing everywhere parallel or antiparallel to the magnetic field lines. Maximum magnetic field strength should be observed at the center, i.e., on the axis of the flux tube under ideal force-free conditions. Burlaga (1988), Marubashi (1986) and Lepping *et al.* (1990) found a good agreement between the in situ observed magnetic field characteristics of MCs and self-consistent magnetohydrodynamic (MHD)- calculations based on the force-free flux tube concept. In contrast to the model of cylindrical shaped clouds, Vandas and Fisher (1991) and Vandas *et al.* (1993) suggested that MCs might rather be explained in terms of spheroidal topologies, and Farrugia *et al.* (1995) found that the flux tube model better

describes in situ field structure of MCs than the spheromak model.

The bidirectional flows of supra-thermal electrons along magnetic field lines, as observed inside IPCs, support the assumption of closed field lines, in agreement with observations that these plasma and magnetic field entities in the solar wind stem from regions in the solar atmosphere where the magnetic field lines form closed loops (Gosling *et al.*, 1973; Montgomery *et al.*, 1974; Bame *et al.*, 1981; Gosling *et al.*, 1987). It has been proposed that the ejecta from SMEs evolve into large-scale magnetic flux ropes that remain magnetically connected to the solar atmosphere (e.g. Gosling, 1990; Gosling *et al.*, 1992). Observations of solar energetic particles (SEPs) associated with MCs and IPCs (Kahler and Reames, 1991; Bothmer *et al.*, 1996b), of flare particles inside a MC (Farrugia *et al.*, 1993) and of BDE-anisotropies (Phillips *et al.*, 1992) are also in favor of this picture. However, BDEs could be present on closed magnetic field lines disconnected from the Sun as in plasmoids, or on closed field lines that are rooted with its footpoints in the solar corona as in magnetic flux ropes. Unfortunately, current spacecraft measurements do not allow us to infer the overall topology of MCs unambiguously.

MCs have hitherto been studied extensively only at near Earth-orbit and beyond. The Helios mission has provided us with solar wind data in the inner heliosphere between 0.3 and 1 AU over a time interval covering more than one solar activity cycle. The purpose of our study is to extend the investigations of MCs to this range of the inner heliosphere. A further scientific goal of the work is to clarify the origin of MCs at the Sun, especially to validate the probable association with erupting solar prominences (disappearing filaments).

2 Identification of magnetic clouds in the Helios data

For the study presented here we used hourly averaged plasma and magnetic field data of the instruments E1 & E2 onboard the Helios probes. A detailed description of the instruments can be found in the papers by Rosenbauer *et al.* (1981), Schwenn *et al.* (1975) and Musmann *et al.* (1975). The complete set of interplanetary magnetic field (IMF) and plasma data of Helios 1 covers the time interval December 1974 through July 1981. Data of Helios 2 span the time-interval January 1976 through March 1980.

Only a small number of magnetic clouds have been studied previously in the inner heliosphere at distances closer to the Sun than 1 AU (e.g., Burlaga *et al.*, 1982). Thus, we had to establish appropriate identification criteria applicable to identify MCs between 0.3 and 1 AU. According to Burlaga (1991), MCs can be identified at 1 AU through the following criteria: (1) the magnetic field direction rotates parallel to a plane through a large angle during a time-interval of the order of one day; (2) the magnetic field strength inside a MC is higher than in the average solar wind; and (3) the temperature inside a MC is lower than average. Note,

that the rotation of the magnetic field direction can occur in any direction relative to the ecliptic. Previous MC observations at 1 AU have shown that the criteria (2) and (3) imply low plasma- β values for MCs (Klein and Burlaga, 1982). Each individual criterion might occur independently of a MC, e.g. in association with planar magnetic structures (PMS) (Nakagawa *et al.*, 1989; Neugebauer *et al.*, 1993), co-rotating interaction regions (CIRs) or sector boundaries (SBs) (Bothmer and Schwenn, 1992a; 1996). For a unique identification of a MC all three criteria have to be fulfilled simultaneously.

The enormous radial size of MCs (~ 0.25 AU at 1 AU) was first evidence for an expansion of MCs in the heliosphere (Klein and Burlaga, 1982). The Voyager spacecraft indeed observed larger radial sizes for MCs in the outer heliosphere (Burlaga and Behannon, 1982). The sizes of MCs observed near Earth's orbit cannot be considered to be typical of those that might be observed closer to the Sun. Further, the spacecraft's trajectory through a MC influences the observed duration time and thus its apparent size. Due to these uncertainties we concentrated our search for MCs in the Helios data on solar wind events associated with the basic property of a MC, i.e., a smooth coherent directional change of the magnetic field vector on time scales of several hours to days. These changes can be identified in the magnetic field components B_x , B_y , B_z which are the Cartesian components of the IMF in solar ecliptic coordinates (SEC), and in the latitudinal and longitudinal excursions of the magnetic field vector, i.e., by variations in the angles θ and φ . B_x points from the spacecraft towards the Sun, B_y points in the ecliptic plane normal to B_x towards the East, and B_z points normal to the ecliptic, towards the ecliptic north pole. Accordingly, θ and φ are the magnetic field's polar ($\theta = 90^\circ \equiv N$, $\theta = -90^\circ \equiv S$) and azimuthal ($\varphi = 0^\circ \equiv$ sunward direction, $\varphi = 90^\circ \equiv E$, $\varphi = 270^\circ \equiv W$) angles.

Based on these criteria we investigated plots of the plasma and magnetic field data of the Helios spacecraft taking into account all events in which variations in the magnetic field direction occurred in θ or φ and in the cartesian components over time intervals of several hours. We included events with irregular or weak field variations and variations at sector boundaries. All events were then investigated with the minimum variance analysis (MVA) (Sonnerup and Cahill, 1967) described in the Appendix to identify smooth rotations of the magnetic field vector on time scales of the order of several hours. The MVA-method can be accurately applied for directional changes of the magnetic field vector exceeding $\sim 30^\circ$ (Burlaga and Behannon, 1982). Thus MCs were identified by a smooth rotation of the field direction in the plane of maximum variance, if the plasma temperature was lower and the magnetic field strength higher than in the surrounding solar wind (see criteria 2 and 3), and finally the error criteria of the minimum variance method were satisfied ($\lambda_2/\lambda_3 \geq 2$; λ_2 and λ_3 correspond to the eigenvalues of the directions of intermediate and minimum variance, see Appendix).

Figure 1 shows an example of a MC in the Helios data for a radial distance of 0.4 AU from the Sun. The

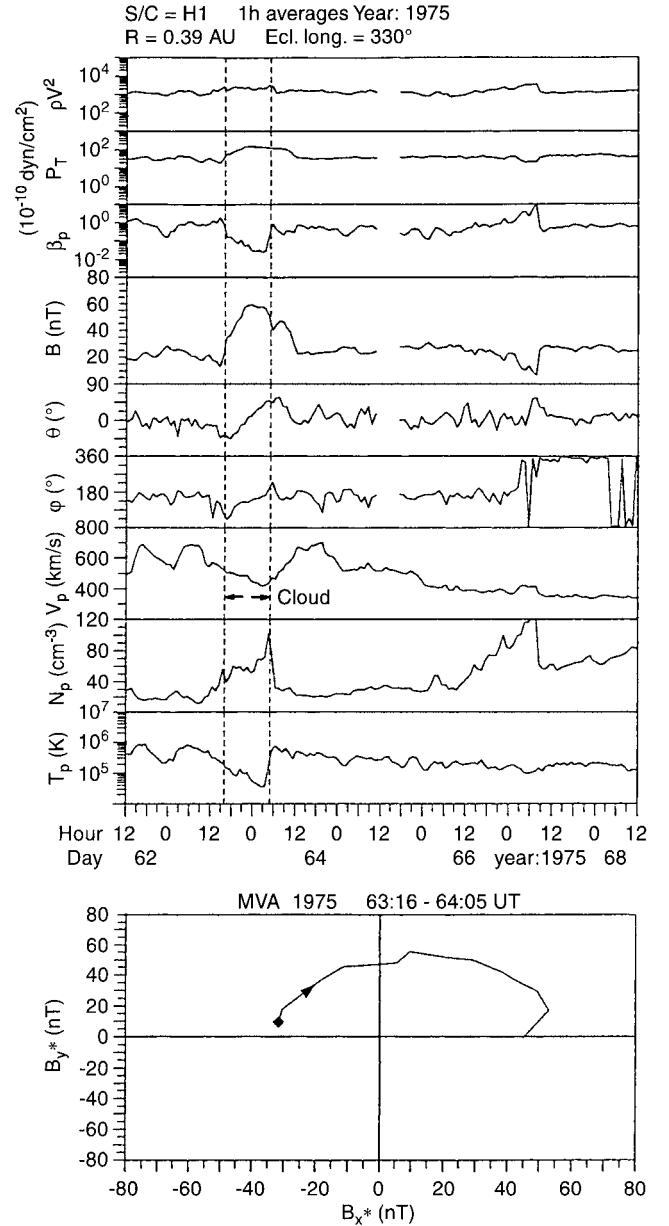


Fig. 1. Solar wind parameters for an MC (dashed lines) observed with Helios 1 at 0.4 AU in March 1975, day 63, 16 UT – day 64, 05 UT. *Top to bottom:* solar wind momentum flux $\rho_p V^2$, total plasma pressure P_T (10^{-10} dyn/cm 2), plasma- β_p calculated from the proton parameters, magnetic field magnitude B [nT], polar and azimuthal angles θ , φ of the magnetic field vector in solar ecliptic coordinates, proton velocity V_p , proton number density N_p , and proton temperature T_p . The *bottom part* of the figure shows the rotation of the magnetic field vector in the plane of maximum variance for the time interval of the MC. The *diamond* denotes the initial point of the plot and the *arrow* indicates the direction of the rotation of the magnetic field

MC was basically identified from the smooth south to north turning of the magnetic field vector relative to the ecliptic. Note the rotation of θ in Fig. 1 covering the time interval marked with dashed lines, and that of the magnetic field vector in the plane of maximum variance. Here the diamond denotes the initial point of the plot and the arrow indicates the direction of the rotation of

Table 1. MCs identified in the Helios 1/2 data from December 1974 until July 1981 between 0.3–1 AU. The columns from left to right denote: time of the MC, time of associated shock, s/c, distance to Sun, Earth-Sun-Helios-angle, heliographic latitude, direction of minimum variance in solar ecliptic coordinates (φ = azimuthal angle in the ecliptic plane, θ = polar angle), direction of MC-axis (direction of intermediate variance), error criteria for the minimum variance analysis (ratio of the eigenvalues of intermediate, λ_2 , to minimum variance, λ_3 , ratio of the component of minimum variance to the average field magnitude inside the MC), flux tube type of the MC. Further abbreviations indicate if MCs had also been observed at Earth's orbit (@ \equiv yes, $- \equiv$ no) and MCs for which an IPC had been detected with the Helios photometers (* \equiv yes; m \equiv maybe; S, N, E, W denote the directions with respect to the s/c if a white-light brightness enhancement occurred in a preferred direction of the sky). + symbols mark MCs for which a Helios directed SME has been observed (Sheeley *et al.*, 1985)

Magnetic cloud		$\langle V_{MC} \rangle$ [km/s]		Associated shock		S/C	R	Longitude Latitude		Min-Var		Axis		Error criteria		MCType
Year	Day	Hours		Day	Hour		[AU]	[°]	[°]	[°]	[°]	[°]	[°]	λ_2/λ_3	$\langle B_z \rangle / \langle B \rangle$	
1975	7@	00–10*	705	6	20		0.92	352	-3	154	+23	102	-56	16	0.5	SEN
	63,64	16–05m	467	-			0.39	30	-7	222	-36	124	-10	20	0.1	SEN
	92	06–16*	384	91	13		0.48	157	+7	234	-02	144	+13	7	0.1	SEN
	313	03–18mN	390	-			0.81	355	+4	220	+47	179	-35	3	0.1	NWS
	321@	06–18-	329	-			0.87	355	+3	145	-21	42	-31	14	0.1	SEN
1976	90@	09–21*E	465	90	4 (17)		0.47	009	-6	204	+6	91	+75	4	0.1	SWN
	187	03–21-	378	-			0.98	156	-1	181	+12	91	-1	16	0.2	SWN
1977	29,30	10–10-	456	29	1		0.95	323	-2	191	-15	103	+8	7	0.4	NES
	76	05–20*W	369	75	19		0.71	345	-7	201	+13	76	+68	2	0.3	SWN
	78,79	22–08*	435	78	7		0.57	330	-7	185	-43	24	-45	6	0.2	NWS
	159,160	18–13*	413	159	7		0.86	150	+3	207	-27	121	+6	4	0.2	SEN
	240,241-	14–10*	360	240	6		0.84	135	-4	205	+17	112	+12	120	0.4	SEN
	268,269	14–12-	658	268	2		0.57	145	-7	146	-1	55	-32	4	0.7	NES
	335	14–00*	413	335	5,8		0.75	324	+5	197	-41	119	+14	27	0.2	SEN
	3,4@	15–17-	733	3	8		0.95	320	+2	265	+40	124	+43	10	0.2	SWN?
	4,5@	08–10*	570	3	14		0.94	355	-3	212	+38	104	+21	16	0.2	SWN
	6	01–13*	567	5	3		0.95	354	-3	158	+4	74	-59	40	0.3	SWN
1978	17@	01–23-	325	-			0.98	351	-4	175	+10	82	+13	23	0.1	NES
	29,30@	12–01-	518	29	21		0.98	347	-5	139	+30	45	+7	7	0.3	SEN
	37,38	16–16-	403	37	7		0.98	344	-2	244	+27	154	-1	2	0.2	NWS
	46,47@	14–20-	484	46	1		0.95	307	-5	224	-23	71	-65	3	0.1	NWS?
	47,48@	03–09*	430	46	1		0.95	341	-6	164	+16	76	-6	12	0.3	SWN
	61,62-	01–01*	446	60	12		0.87	305	-4	189	-20	119	+44	17	0.1	SEN
	92@	02–07-	431	91	8		0.61	343	-7	206	-8	122	+33	126	0.3	SWN
	114	12–19*	395	-			0.32	039	-1	143	+23	46	+17	5	0.3	NWS
	189,190	22–22-	456	189	10		0.94	134	+2	167	-23	76	-23	8	0.1	SEN
	292,293-	01–14*	488	291	16		0.47	143	-7	143	+8	51	+18	7	0.7	NES
	358,359	15–15-	401	358	8		0.85	351	-1	208	+25	142	-40	2	0.4	SEN
	363,364	09–14-	671	362	23		0.85	309	-30	230	+47	177	-30	5	0.2	SWN

1979	58,59	15-15-	582	58	2	H1	0.96	293	-2	198	-23	28	-67	5	0.6	SWN
	62-	09-17-	648	62	2	H1	0.94	292	-2	146	-19	50	-16	11	0.5	SEN
	93@	02-18-	436	92	20	H2	0.68	333	-7	164	-19	79	+14	10	0.3	SWN
	129	06-12*	625	129	3,(8)	H2	0.30	84	+6	244	-28	173	+32	12	0.4	SEN
	148,149+	23-07*	515	148	18	H1	0.43	91	+7	223	+5	133	-3	6	0.3	SEN
	305	03-19*	325	-		H1	0.50	124	-7	116	-23	25	-4	8	0.6	SWN
	365	00-10-	397	364	15	H2	0.84	345	-1	212	+36	68	+48	3	0.2	NES
1980	82,83+	17-12-	356	82	14	H1	0.92	276	-3	213	+28	100	+37	17	0.1	NES
	90+	01-15-	553	89	11	H1	0.88	275	-3	174	-25	81	-6	4	0.5	SEN
	162,163	17-01-	467	161	18	H1	0.41	72	+7	213	+25	130	-15	2	0.2	SWN
	172+	02-20*	407	171	19	H1	0.53	92	+7	168	+26	87	-17	36	0.1	NWS
	175+	05-17m	350	174	20	H1	0.57	96	+7	141	-30	72	+31	14	0.1	NWS
	231	00-18-	370	230	18	H1	0.97	103	+1	184	-32	90	-6	22	0.1	SWN
1981	117,118	09-03-	478	117	3	H1	0.79	260	-5	263	+50	3	+8	4	0.8	SWN
	131,132+	15-03-	818	131	7	H1	0.66	264	-6	99	-12	25	+54	4	0.1	SEN
	146,147-	03-07*	477	145	20	H1	0.48	281	-7	197	-3	157	+86	4	0.4	SEN
	170+	03-09*	644	170	1	H1	0.34	30	+5	143	+53	56	-3	9	0.3	SWN

the magnetic field. The plasma- β_p (calculated from the proton parameters only) was much lower inside the cloud whereas the total plasma pressure $P_T(nkT+B^2/8\pi)$ was much higher compared with the surrounding flow. The solar wind momentum flux (ρV^2) remained fairly constant during the time interval shown in Fig. 1. Another interesting feature is the high plasma density compared with the ambient solar wind which may indicate that the low densities often found in MCs may be a result of considerable expansion of MCs with increasing distance from the Sun (see Klein and Burlaga, 1982; Burlaga and Behannon, 1982 and the results presented in Sec. 7).

Our search through the Helios data yielded a set of 46 certain MCs for the years 1974–1981 over the distance range 0.3–1 AU (Table 1). Note that many more cases hampered by data gaps or other disturbances were excluded. It is interesting to remark that generally solar wind structures with unique internal field rotations identified in the inner heliosphere were associated also with low plasma- β values ($\beta \lesssim 0.1$). This finding agrees with the characteristics of MCs at 1 AU (Burlaga, 1991).

3 Magnetic clouds, interplanetary shocks and SMEs

We examined the high resolution plasma and IMF data set of Helios 1/2 to investigate the association of the identified MCs with interplanetary shock waves. 39 (85%) of the 46 MCs observed between 0.3–1 AU were associated with fast-forward shocks (see Table 1). The arrival times of the MCs lagged the shocks by 5–25 h with an average delay of ~ 11 h. Solar wind momentum flux ρV^2 often peaked near the “nose” of the MCs several hours after arrival of the interplanetary shock (e.g., Figs. 5, 6, 7). In agreement with theoretical models for interplanetary shocks (see, e.g., Hundhausen, 1972), MCs seem to be shock drivers in the inner heliosphere.

Figure 2 shows a histogram in 100 km/sec bins for the average plasma speeds, V (derived from the protons only), measured inside the 46 MCs observed by Helios 1/2. Most MCs were found to have a propagation speed similar to that of typically slow (300–500 km/s) solar wind. Only a few MCs had speeds comparable to the typically high (~ 700 –800 km/s) speed solar wind flow from coronal holes. The peak of the speed distribution in Fig. 2 lies in the range 400–500 km/s with an average value of $\langle V \rangle = 478$ (rms = 118) km/s. As indicated in Fig. 2 by the open portion of the bars, the MCs that were not associated with shocks were amongst the slowest MCs with speeds typically less than 400 km/s. The finding that most of the MCs with speeds beyond 400 km/s were associated with shocks is in agreement with the result obtained from direct correlations of coronagraph observations of SMEs directed towards Helios 1 and in situ shock disturbances at the spacecraft: When Helios 1 was located inside the angular extent of a SME seen in the coronagraph’s field of view as viewed from Earth’s orbit, it was definitely hit by an interplanetary shock as consequence of the SME, if the SME’s speed was greater 400 km/s, independent of the space-

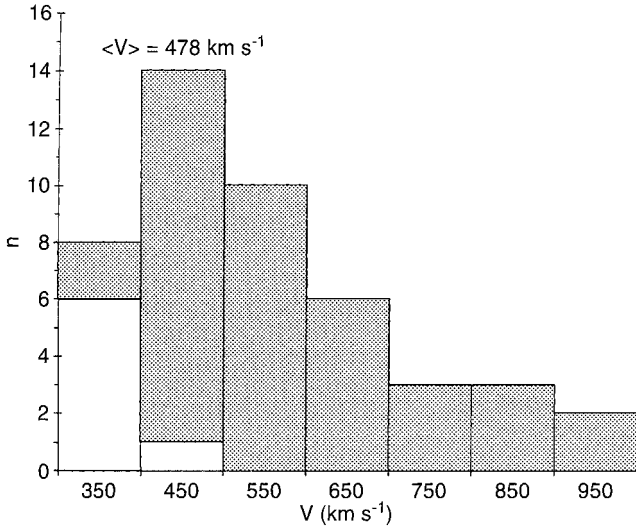


Fig. 2. Frequency distribution for the average proton speeds of MCs in bins of 100 km/s as observed by Helios 1/2 during 1974–1981 between 0.3–1 AU. Open portions of the bars denote the number of MCs not associated with a shock. The average speed for MCs was 478 km/s

craft’s distance to the Sun (Sheeley *et al.*, 1985). The average speed of MCs shown in Fig. 2 is comparable to the speed distribution derived for SMEs (see Hundhausen *et al.*, 1994) in the sense that the shape of the distribution drops off at speeds beyond 600 km/s, i.e., very fast MCs and SMEs are less frequent. However, the fact that the observed “low speed” MCs were so well associated with shocks may be indicative for deceleration of the MCs on their way out from the Sun depending on the ambient solar wind flow through which they had to plow. This should be especially important during solar maximum when the heliosphere consists of very inhomogeneous solar wind flow.

The high percentage of MCs associated with shocks found here is similar to that obtained for MCs observed near Earth orbit during the years 1978–1982 (Zhang and Burlaga, 1988) but higher than that observed during the years 1967–1978 (Klein and Burlaga, 1982). None of the 46 MCs observed between 0.3–1 AU was associated with a forward-reverse shock pair, which has been observed for interplanetary counterparts of SMEs detected by the Ulysses spacecraft at high heliographic latitudes in fast solar wind coronal hole flow (Gosling *et al.*, 1994). Since it has been uniquely shown that fast SMEs at the Sun are the prime sources of transient shocks in the interplanetary medium (Sheeley *et al.*, 1985; Schwenn, 1986; Cane *et al.*, 1987) our findings support a close association of MCs with SMEs.

We compared our list of shock-associated MCs with the SME shock event list derived from correlated Helios-Solwind observations during 1979–1982 published by Sheeley *et al.* (1985) and found that 7 MCs occurred as direct consequences of SMEs towards Helios 1 (see Table 1). Another study based on this event list showed that 17 out of 25 Helios-directed SMEs which caused an interplanetary shock wave at Helios 1 were associated with driver gases, but only seven of the driver gases were MCs (Bothmer and Schwenn, 1996).

Note that this study included only those events of the original list of Sheeley *et al.* (1985) in which the presence or absence of a piston had been uniquely identified and for which both plasma and magnetic field measurements had been obtained.

Both observations are plausible since an interplanetary shock should commonly have a broader spatial extent than the associated driver gas (Borrini *et al.*, 1982; Richardson and Cane, 1993) and IPCs might have different magnetic structures, as has been pointed out by Bothmer and Schwenn (1996).

The Helios photometers, designed to measure the zodiacal light, mounted at an angle of 16°, 31° and 90° with respect to the spacecraft’s spin-axis, have been used as wide-angle coronagraphs in order to identify SMEs by their white-light brightness increases close to the Sun and also as IPCs in the interplanetary medium (Webb and Jackson, 1990). Striking evidence was found that plasma clouds resulting from SMEs did pass the Helios spacecraft. We compared our list of MCs with the list of IPCs detected by the Helios photometers (Jackson *et al.*, 1994). Without considering the duty-cycle of the instruments we find 23 of our MCs (see Table 1) associated with an interplanetary white-light event identified as an IPC, i.e., MCs are intimately connected to IPCs from SMEs (see also Webb *et al.*, 1993). Brightness enhancements in the 16° and 31° photometers usually preceded the peak of the 90° photometer by 1–2 days. The peak in the 90° photometer data is thought to indicate the passage of a plasma cloud right at Helios. We find also that there were IPCs identified from the Helios photometer data that had not the structure of MCs, in agreement with the finding that MCs are only a subset of all IPCs (Gosling, 1990; Bothmer and Schwenn, 1996).

4 Magnetic clouds as magnetic flux tubes

Goldstein (1983) has proposed that the magnetic field characteristics of MCs could be explained through force-free large-scale cylindrical magnetic flux tubes. The idealized sketch in Fig. 3 depicts the magnetic field signatures expected for a cylindrical magnetic flux tube which moves radially away from the Sun, i.e., in the x-direction, and centrally passes a Helios spacecraft in the ecliptic. Here, the maximum change in the magnetic field direction would occur normal to the ecliptic in the B_x^* -component due to the south (S) to north (N) turning of the magnetic field vector. The component of minimum variance, B_z^* , corresponds to the radial component B_x , because B_x is zero in the whole MC. At the center of the flux tube, i.e., at the cloud’s axis, the magnetic field is directed in the azimuthal (+y-) direction [eastern (E)-direction]. This is also the direction of intermediate variance, because B_y^* is zero at the cloud’s outer boundaries (on the surface of the cylinder). In reality, the flux-tube axis may have any orientation relative to the ecliptic plane and to the radial direction i.e., the actual variance directions x^* , y^* , z^* usually differ from the solar ecliptic coordinate system in x, y, z. Whether the model is appropriate to describe the structure of

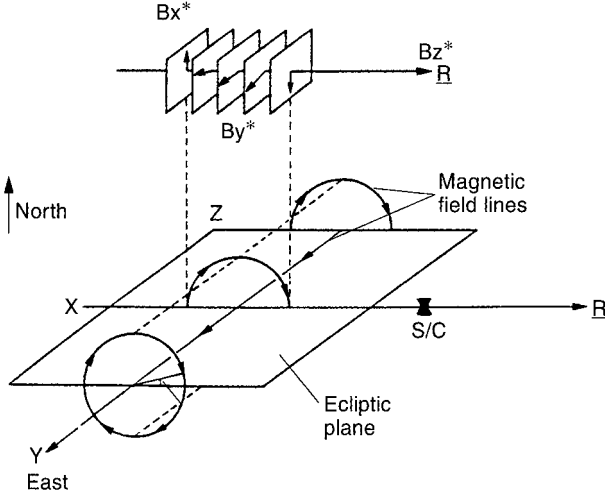


Fig. 3. Idealized sketch showing the expected magnetic field signatures during passage of a large-scale cylindrical magnetic flux tube over a s/c (adapted from Goldstein, 1993). B_x^* , B_y^* , B_z^* are the components of maximum, intermediate and minimum variance of the magnetic field vector. The south to north [$-B_x^*$ (S) to $+B_x^*$ (N)] turning of the magnetic field vector is associated with passage of the circular field lines at the cloud's outer boundaries. At the center of the flux tube (at the axis) the magnetic field is directed along the $+B_y^*$ (east)-direction. A semicircular rotation appears when B_y^* is plotted versus B_x^* (top diagram)

MCs can be inferred from the results of MVA (Klein and Burlaga, 1982) applied to measured data.

In Fig. 4 the in situ characteristics of a MC associated with an interplanetary shock observed by Helios 1 are displayed together with the results of the variance analysis of the magnetic field data. A smooth rotation of the magnetic field direction from south to north (note the changes of θ in Fig. 4) starts in coincidence with a drop to very low β_p values (β_p is calculated from the proton parameters only) at 01 UT on day 61. After the field vector has reached its most northern dip at 01 UT on day 62 it relaxes back to the average Parker spiral direction of the IMF. Note also the high total internal plasma pressure P_T inside the MC. For shock-associated MCs we found the P_T -values usually to be highest between the shock and the leading edge of the cloud where the plasma turbulence is high. This region is commonly called the cloud's sheath (in analogy to the Earth's magnetosheath, see also McComas *et al.*, 1988). Maximum pressure often occurred right at the nose of the MCs, thus indicating that MCs were indeed the drivers of the associated shocks. The large decrease of the magnetic field strength from the front to the rear portion of the cloud certainly stems from compression of the cloud's leading edge while plowing through slower moving ambient solar wind ahead. The in situ observations show that MCs are certainly not ideal force-free structures in reality.

The boundaries of the cloud have been determined from two parameters: the β_p -slope and the time interval of rotation of the magnetic field vector. Note that unique criteria to determine MC-boundaries have not been established in the literature. The size of the MC as calculated from the cloud's average plasma speed and its

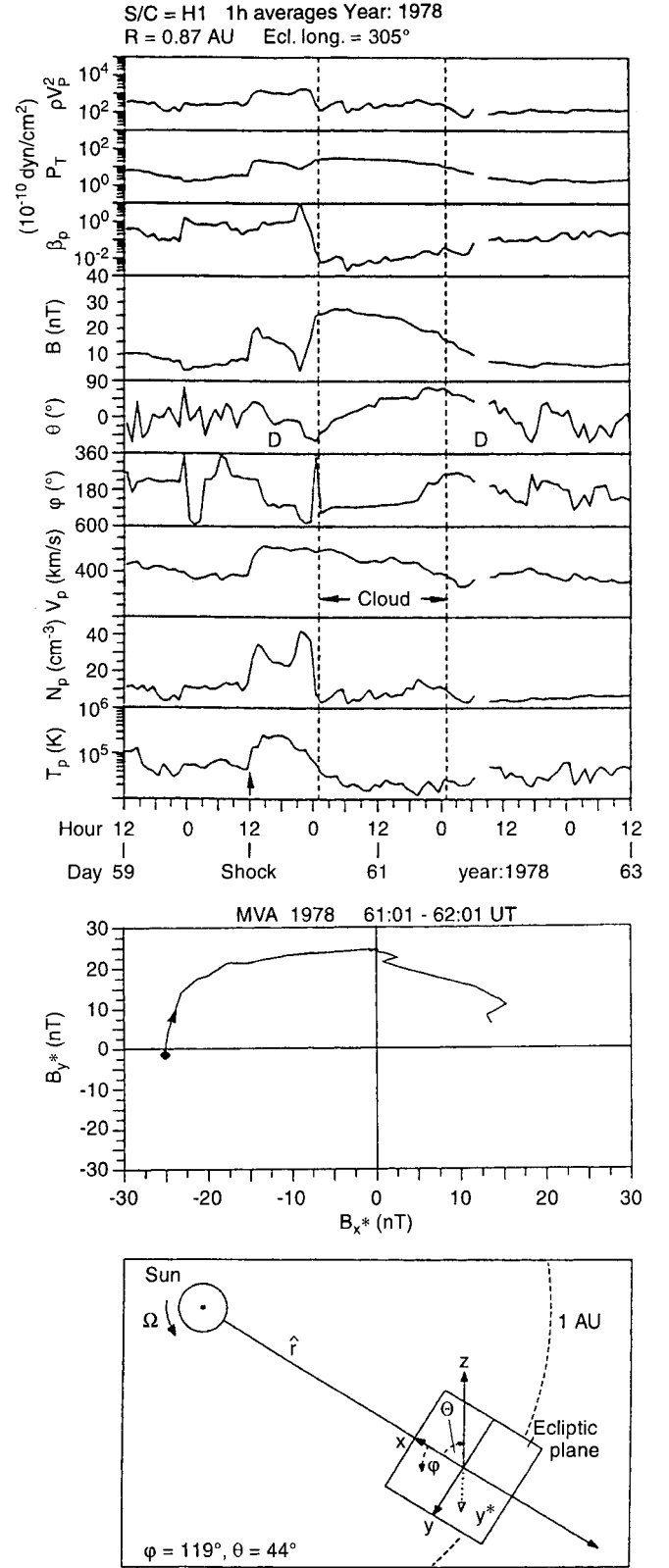


Fig. 4. Solar wind parameters (top panel) and results from the minimum variance analysis (middle panel) for a MC (dashed lines) observed with Helios 1 at 0.9 AU, day 61, 01 UT – day 62, 01 UT in 1978. The MC was associated with a shock on day 60, 12 UT. The draping of the IMF in the cloud's sheath region and behind the cloud's rear portion is labelled (D). The calculated orientation of the cloud's axis, sketched in the bottom panel, was $\phi = 119^\circ$, $\theta = +44^\circ$. The direction of minimum variance was $\phi = 189^\circ$, $\theta = -20^\circ$

duration was 0.27 AU, a typical size for MCs observed at 1 AU (Klein and Burlaga, 1982).

Results of the MVA of the magnetic field data are presented in Fig. 4 in the lower diagrams. During the time interval of the MC, B_z^* was small compared to the magnitude of the total magnetic field $B(\langle |B_z^*| \rangle / \langle B \rangle = 0.09)$. The ratio of the intermediate eigenvalue λ_2 to the minimum eigenvalue λ_3 was $\lambda_2/\lambda_3 = 16.8$, and the angle between the first and the last value of the field vector was 150° (see Table 1 which includes the results of the MVA for all 46 MCs). The variance directions were well determined according to the error criteria given by Siscoe and Sney (1972) and Lepping and Behannon (1980). Plotted in the bottom panel in Fig. 4 is the component of intermediate variance B_y^* versus the component of maximum variance B_x^* . The hodogram shows that the magnetic field vector rotates smoothly through a large angle in the plane of maximum variance.

The comparison of the in situ observed magnetic characteristics of this MC with that expected from the cylindrical flux tube model proposed by Goldstein (1983) is in excellent agreement. The calculated direction of the cloud's axis, i.e., the direction of intermediate variance, was $\varphi = 119^\circ$ (azimuthal direction in the ecliptic), $\theta = +44^\circ$ (inclination relative to the ecliptic). In contrast to the in-ecliptic orientation of the flux tube axis sketched in Fig. 3, the axis of this MC had a considerable inclination relative to the ecliptic. A sketch of the axis direction at Helios 1 is given in Fig. 4 at the bottom. The calculated minimum variance direction was pointing southward of the ecliptic ($\theta = -20^\circ$), slightly towards the west ($\varphi = 189^\circ$), close to the probable direction of propagation, which is radially outwards from the Sun.

Although more sophisticated cloud models have been developed, including also a possible offset of the spacecraft relative to the cloud's axis (Burlaga, 1988; Lepping *et al.*, 1990; Vandas *et al.*, 1993), a MVA of the magnetic field data is appropriate to estimate the cloud's axis orientation to first order.

5 Configurations and orientations of MCs in interplanetary space

The concept of cylindrical flux tubes allows different magnetic configurations. In the context of the investigation of the orientations of MCs in interplanetary space, it will be shown that all the predicted types of flux-tubes actually occur in the solar wind. Assume as in Fig. 3 that the axis of a MC lies in the ecliptic plane, normal to the sun-spacecraft line. A flux tube (MC) cloud possess one of the four different magnetic configurations presented in Table 2, which differ in the orientation of the magnetic field lines at the cloud's outer boundaries and on its axes. These MCs would produce different characteristic magnetic signatures when passing over a spacecraft. According to Table 2, MCs can be classified into SEN (SWN) clouds where the magnetic field vector turns from south (S) to east (E) (west, W) on the cloud's axis and finally to the north (N)

at its rear boundary, and vice versa into NES (NWS) clouds. More generally we can define $\Delta\theta > 0^\circ$ for SN-clouds ($\Delta\theta < 0^\circ$ for NS clouds) with $180^\circ < \varphi < 0^\circ$ ($360^\circ > \theta > 180^\circ$) for SEN, NES (SWN, NWS) clouds. The classification introduced here involves the property of magnetic helicity (see Burlaga, 1988; Lepping *et al.*, 1990). SEN and NWS clouds possess left-handed (LH), SWN and NES clouds, right-handed (RH) helicity as viewed by an observer looking towards the Sun. In situ observations of these four different MC-types will be presented in the following section.

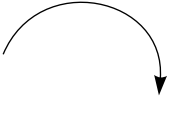
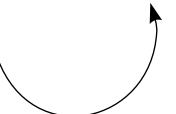
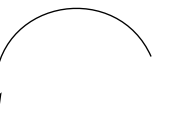
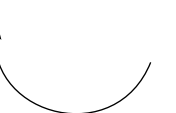
5.1 Magnetic cloud types

5.1.1 SEN-type MC. Figure 5 shows a shock-associated SEN-cloud observed by Helios 1 in August 1977. The magnetic field direction changes smoothly from south to north during the cloud's passage over the spacecraft. Note that the southward dips of the magnetic field vector in the cloud's sheath region (see the variations of θ , labelled D, in Fig. 5) were caused by draping of the IMF ahead of the MC (Gosling and McComas, 1987; McComas *et al.*, 1988). At the cloud's center the magnetic field had an eastward direction. In the plane of maximum variance we find a rotation of the magnetic field vector from negative to positive B_x^* -values and $+B_y^*$ -values. In this case, the component B_z^* along the direction of minimum variance had a constant, but substantial non-zero value (see Table 1) which might be indicative of twisted rather than pure circular outer magnetic field lines. The ratio of λ_2/λ_3 was 120, and the direction of minimum variance was $\varphi = 205^\circ$, $\theta = +17^\circ$ in SECs. The cloud's axis was lying almost parallel to the ecliptic plane ($\theta = 12^\circ$), normal to the sun-spacecraft-line ($\varphi = 112^\circ$). Note also the high internal total plasma pressure P_T and the very low plasma- β_p inside the cloud. Other interesting features of this MC were its high magnetic field strength with a very symmetric time profile and the extreme drop of the plasma temperature near the center of the cloud.

5.1.2 SWN-type MC. A shock-associated SWN-cloud observed in December 1978 is presented in Fig. 6. The magnetic field changes of θ and φ look relatively moderate, but a unique rotation of the field vector is present in the plane of maximum variance. Note that this cloud might have gone undetected had the data simply been inspected by eye. The axis of the cloud was lying close to the sun-spacecraft line ($\varphi = 177^\circ$), pointing slightly towards the south ($\theta = -30^\circ$). The minimum variance direction was $\varphi = 230^\circ$, $\theta = +47^\circ$. Inside the cloud β_p is low and P_T is high, as is typical. The strong decrease of the magnetic field strength from the cloud's front to its rear has probably been caused by interaction of this MC with the ambient solar wind.

5.1.3 NES-type MC. A shock-associated NES-cloud observed by Helios 1 in January 1977, directly at a

Table 2. Sketch showing the different magnetic configurations of MCs and their magnetic helicity (left-handed (LH), right-handed (RH)) based on the magnetic flux-tube concept and the field rotation that a s/c would observe during the cloud's passage. The number at the bottom indicates how often each MC-type was observed by Helios 1/2 between 0.3–1 AU during 1974–1981

MC Type	Magnetic helicity	Variation of magnetic field vector	Direction of magnetic field on flux tube axis	Rotation of magnetic field vector in Bz - By -plane (Bx - By -plane)
SEN	Left-handed	South ($-Bz$) \rightarrow north ($+Bz$)	East ($+By$)	
SWN	Right-handed	South ($-Bz$) \rightarrow north ($+Bz$)	West ($-By$)	
NES	Right-handed	North ($+Bz$) \rightarrow south ($-Bz$)	East ($+By$)	
NWS	Left-handed	North ($+Bz$) \rightarrow south ($-Bz$)	West ($-By$)	
Orientations for high inclinations to the ecliptic SEN, NWS, SWN, NES		East ($+By$) \rightarrow west ($-By$) West ($-By$) \rightarrow east ($+By$)	North ($+Bz$) \rightarrow south ($-Bz$) South ($-Bz$) \rightarrow north ($+Bz$)	Rotations in By - Bz - (By - Bx -) plane

sector boundary (see the changes of φ) is shown in Fig. 7. The peak of the magnetic field strength at the cloud's front suggests strong interaction with the ambient plasma, which certainly also deformed the magnetic field lines. Although the field vector changes direction from 90° N to 0° , with $\varphi \approx 90^\circ$ at the cloud's center, the plane of maximum variance revealed a rather complex rotation, presumably a result of compression and distortion of the field lines due to the cloud's interaction with the ambient slower moving solar wind. An extremely high overabundance of He^+ has been measured inside this cloud (Schwenn *et al.*, 1980) being most probably cold chromospheric material caused by an erupting prominence (Cane *et al.*, 1986). The cloud's axis was lying in the ecliptic ($\theta = +8^\circ$), normal to the sun-spacecraft line ($\varphi = 103^\circ$). The direction of minimum variance was near the ecliptic, approximately radially outwards ($\varphi = 191^\circ$, $\theta = -15^\circ$). Schwenn *et al.* (1980) explicitly stated that they did not find any evidence for the end-time of the driver gas within 48 h following the arrival of the shock wave at ~ 01 UT on January 29. However, the rotation of the magnetic field vector ends 15 h earlier, around ~ 10 UT on January 30. In this case, the MC might have been the leading portion

of a much larger plasma cloud from a SME, consistent with observations of intermittent intervals of BDEs following the MC (private communication, K. Ivory, Max-Planck-Institut für Aeronomie, Katlenburg-Lindau, Germany), as was the case for an IPC observed recently with the Ulysses spacecraft (Gosling *et al.*, 1995a; Bothmer *et al.*, 1996b). Events like these need further investigation to clarify their nature.

5.1.4 NWS-type MC. A shock-associated NWS-cloud passed Helios 1 in June 1980 (Fig. 8, see also Burlaga *et al.*, 1982). It was directly associated with a Helios-directed SME detected with the Solwind-coronagraph onboard the P78/1 satellite (Sheeley *et al.*, 1985). The MC shows low β_p and high P_T . The magnetic field vector turns from north (N) over west (W) finally to the south (S) at the cloud's rear. Note also the draping of the IMF in the sheath region between the shock and the MC, associated with high β_p values, and the field variations following the cloud when the surrounding IMF drapes back to its undisturbed orientation. The cloud's axis was lying close to the EW-direction ($\varphi = 87^\circ$), slightly

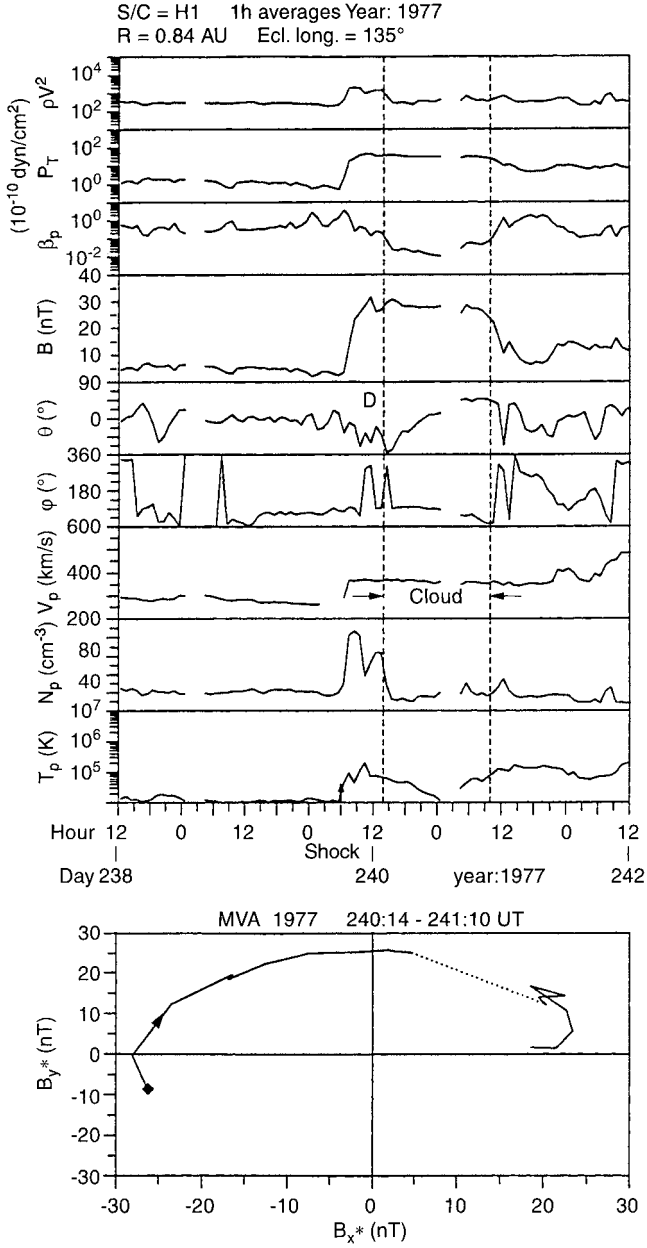


Fig. 5. A MC (dashed lines) of type SEN observed with Helios 1 at 0.8 AU on days 240, 14 UT – 241, 10 UT in 1977. The MC was associated with a shock on day 240, 06 UT. *D* denotes the draping of the IMF in the cloud's sheath region. The bottom panel shows the clockwise rotation with $+B_y$ -values in the plane of maximum variance. The calculated axis orientation was $\varphi = 112^\circ$, $\theta = +12^\circ$, with a minimum variance direction of $\varphi = 205^\circ$, $\theta = +17^\circ$

inclined towards the ecliptic ($\theta = -17^\circ$). The angles for the direction of minimum variance were $\varphi = 168^\circ$, $\theta = +26^\circ$.

5.2 Statistical results

For all 46 MCs we determined their orientation and magnetic configuration (see Table 1). The frequency distribution of the calculated axis orientations in terms

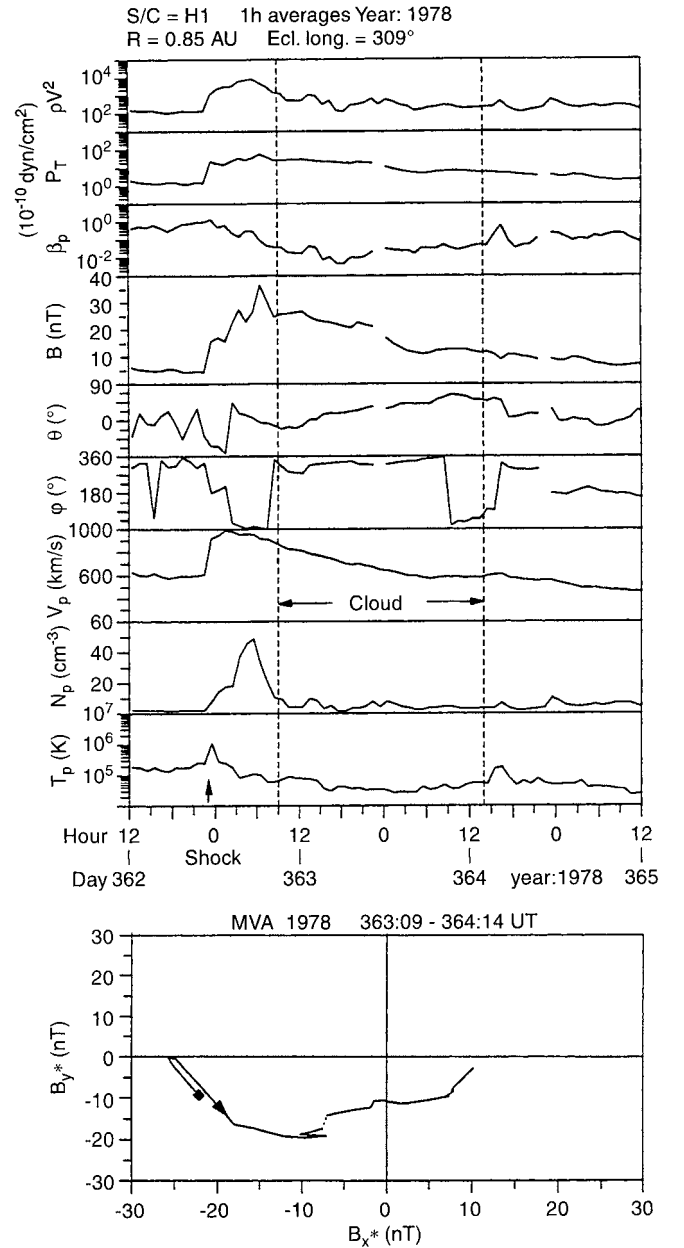


Fig. 6. A MC (dashed lines) of type SWN observed with Helios 1 at 0.9 AU on days 363, 09 UT – 364, 14 UT in 1978, associated with a shock on day 362, 23 UT. The bottom panel shows the counter-clockwise rotation with $-B_y$ -values in the plane of maximum variance. The cloud's axis direction was $\varphi = 117^\circ$, $\theta = -30^\circ$, the minimum variance direction was $\varphi = 230^\circ$, $\theta = +47^\circ$

of θ and φ is shown in Fig. 9. In the top diagram of Fig. 9 we compared the frequency distribution of the axes inclinations to the ecliptic between 90°S – 90°N found for the Helios 1/2 MCs with a random distribution of orientations for the same number of events in bins of 20° . Solid bars represent the number of MCs in each bin, open bars represent the number of cases as expected from the random distribution. Assume that the sky is divided into bins of equal areas of one square each, then for $\theta = 0^\circ$ there are 360 possibilities for φ , but there is only one possibility for $\theta = 90^\circ$. Thus random orientations at 0° should be 360 times more likely than those at

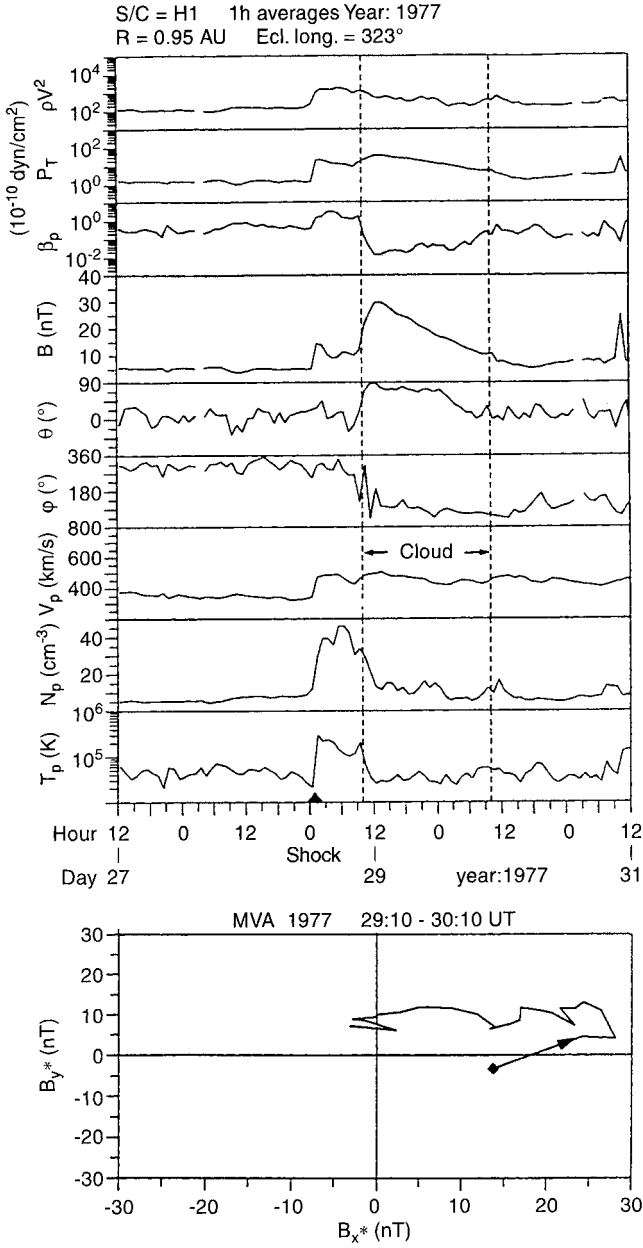


Fig. 7. A MC (dashed lines) of type NES observed with Helios 1 on days 29, 10 UT – 30, 10 UT in 1977. It was preceded by a shock on day 29, 01 UT. The *bottom panel* shows the counter-clockwise “rotation” with $+B_y$ -values in the plane of maximum variance. The cloud’s axis was $\phi = 103^\circ$, $\theta = +8^\circ$, the minimum variance direction was $\phi = 191^\circ$, $\theta = -15^\circ$

90°. A more sophisticated approach shows that the area is proportional to $\int \vartheta$ for the measured range of ϑ , where ϑ is the colatitude. Figure 9 shows that the distributions are roughly comparable, i.e., MCs can have in principal any inclination with respect to the ecliptic. However, there is an indication for an overabundance of axes inclinations in the range 10°S–10°N. Since MC axis do not spread randomly over 180° in azimuth the number of cases in the interval 10°S–10°N should be too peaked in the random distribution in Fig. 9 supporting the assumption that the axes of MCs commonly lie rather close to the ecliptic than highly inclined to it.

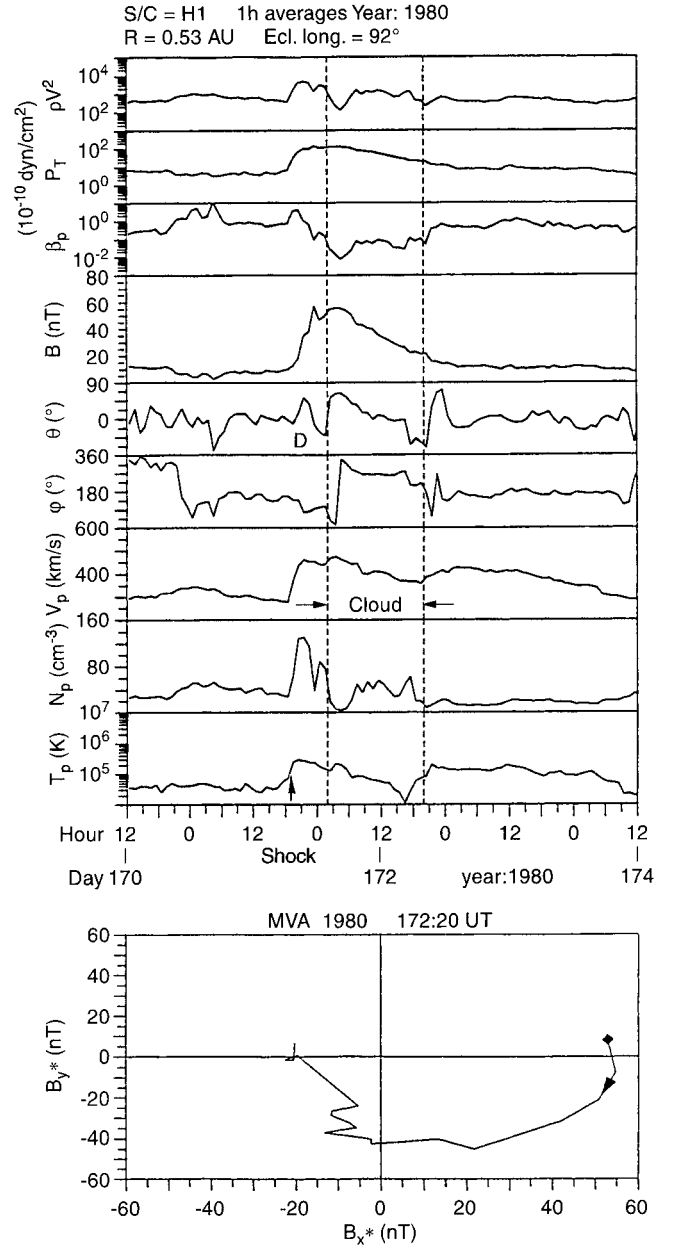


Fig. 8. A MC (dashed lines) of type NWS that passed Helios 1 at 0.5 AU on day 172, 02 – 20 UT in 1980. It was associated with a shock on day 171, 19 UT. *D* denotes the draping of the IMF in the cloud’s sheath region. The *bottom panel* shows the clockwise rotation with $-B_y$ -values in the plane of maximum variance. The calculated axis direction was $\phi = 87^\circ$, $\theta = -17^\circ$, the minimum variance direction was $\phi = 168^\circ$, $\theta = +26^\circ$

The distribution for the azimuthal axes direction of MCs in the ecliptic between 0°–180° is shown in the lower diagram of Fig. 9, also in bins of 20°. The axes scattered about the east-west direction, i.e., normal to the sun-spacecraft line. The average azimuthal orientation of the flux tube axes was $\langle \phi \rangle = 91^\circ$, $\text{rms} = 44^\circ$, and the average orientation of the minimum variance direction was $\langle \phi \rangle = 188^\circ$, $\text{rms} = 38^\circ$ and $\langle \theta \rangle = +3^\circ$, $\text{rms} = 27^\circ$. Our findings confirm the results for IPCs and shocks found by Smith and Phillips (1997).

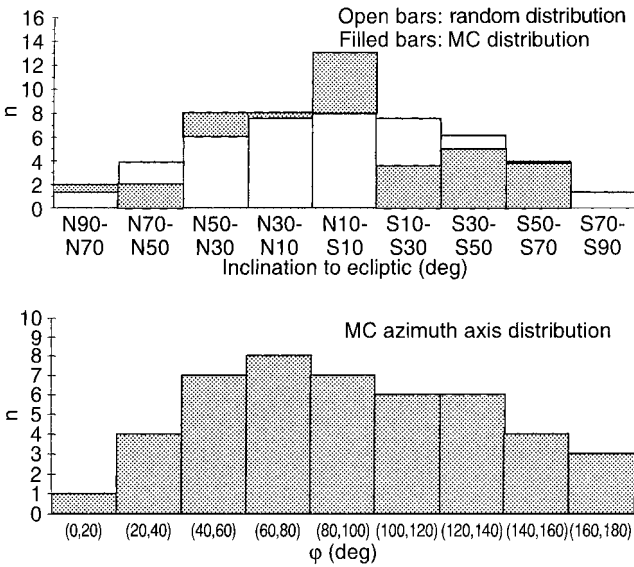


Fig. 9. Frequency distribution for the calculated axis orientations of all 46 MCs observed with Helios 1/2 between 0.3 and 1 AU. *Top diagram:* comparison of the distribution for the cloud's axis inclinations relative to the ecliptic plane in 20° bins between 90°S–90°N with a random distribution. *Solid bars* represent the number of MCs in each bin, *open bars* represent the number of cases expected from the random distribution. *Bottom diagram:* distribution for the cloud's axis azimuthal angles in the ecliptic in 20° bins between 0°–180°

Most of the clouds' axes were lying inside a cone angle of $\pm 45^\circ$ with respect to the east-west direction in the ecliptic, distinctive from the average Parker spiral direction of the IMF ($\varphi \approx 135^\circ$ at 1 AU). The large scatter of the axes orientations around the east-west direction may result from passage through the individual MCs under different aspect angle relative to the center of the MCs. Often left or right wings of the flux tubes may have been encountered by the Helios probes. BDEs seem to be best explained through heat fluxes of supra-thermal electrons along magnetic field lines that remain connected back at both ends to the solar corona (e.g., Gosling, 1990) so that expanding flux tubes should in reality be curved obstacles, which could explain the observed scatter of the axes orientations in the azimuthal direction.

Table 2 shows the frequencies of the different flux rope types observed by Helios 1/2. SEN- and SWN-clouds were observed approximately three times more frequently than NES- and NWS-type MCs during 1974 to 1981. On the other hand, comparing the frequency of MCs by magnetic helicity, the number of left-handed (SEN, NWS) and right-handed (SWN, NES) MCs was the same, as one might expect if these types occur with equal probability. Zhang and Burlaga (1988) found that 79% of the MCs they investigated for the years 1978–1982 were associated with SN-rotations of the magnetic field vector, but they did not give a physical explanation for this preference.

A time-period of the solar cycle similar to the years 1974–1981 is 1984–1991. For these years we inspected the near Earth solar wind data (OMNI-database at

Table 3. Frequencies of SN- and NS-MCs at Earth's orbit during 1984–1991. The dates and types of the individual MCs are given below the Table

Years	SN-MCs	NS-MCs
1984–1991	2	7

(dates and types: 1 May 85, NS; 12 Jul 85, SN; 24 Nov 86, NS; 14 Jan 88, NS; 21 Feb 88, NS; 14 Jun 89, NS; 29 Aug 89, NS; 25 Apr 91, NS; 9 Jul 91, SN)

NSSDC, for a description of this data set see King, (1991), for MCs showing a unique coherent internal SN- or NS-rotation of the magnetic field vector according to the identification criteria of MCs at 1 AU (Burlaga, 1991). Our search yielded only nine clearly defined MCs, since the data are sparse and we did not investigate events with complex structure or time durations shorter than ~ 1 day. Table 3 shows that seven of the nine MCs identified between 1984–1991 had NS-rotations, contrary to the years 1974–1981 where SN-MCs were dominant.

For the years 1974–1981 (solar cycle 21), magnetic field lines in the Sun's northern hemisphere had predominantly northern magnetic polarity, i.e. they were pointing away from the Sun, whereas magnetic field lines in the Sun's southern hemisphere were predominantly directed sunwards (southern magnetic polarity). During 1984–1991 (cycle 22) the magnetic polarities in the Sun's hemispheres were reversed. The observations are in agreement with the findings of Bothmer and Rust (1997) who suggested a solar cycle dependence of the magnetic configuration of MCs. This topic will be further discussed in Sec. 8, 9.

6 Large-scale structure of MCs

MCs are thought to be large-scale solar wind structures. Thus, they should be observable over a substantial range in solar longitude/latitude with the same overall magnetic flux tube structure. One spectacular event in which the same MC was observed by Helios, Voyager and near-Earth spacecraft at distances between 1–2 AU has been studied in detail by Burlaga *et al.* (1990). The authors concluded that the spacecraft observations are in agreement with the model of a MC as a large-scale magnetic flux-rope.

To further clarify the structure of MCs, we investigated near-Earth solar wind data collected in the OMNI-database at NSSDC to see if the MCs observed by Helios 1/2 also passed the Earth (see Table 1). In a few cases, when only plasma data were available and when the MC observed by the Helios s/c had a substantial southward magnetic field, we additionally investigated geomagnetic activity which sensitively responds to large southward components of the IMF (see, e.g., Bothmer and Schwenn, 1995). Doubtful events due to major gaps in the solar wind data and other uncertainties were discarded, leaving 9 cases (note in Table 1 that Helios 1 and Helios 2 observed the same

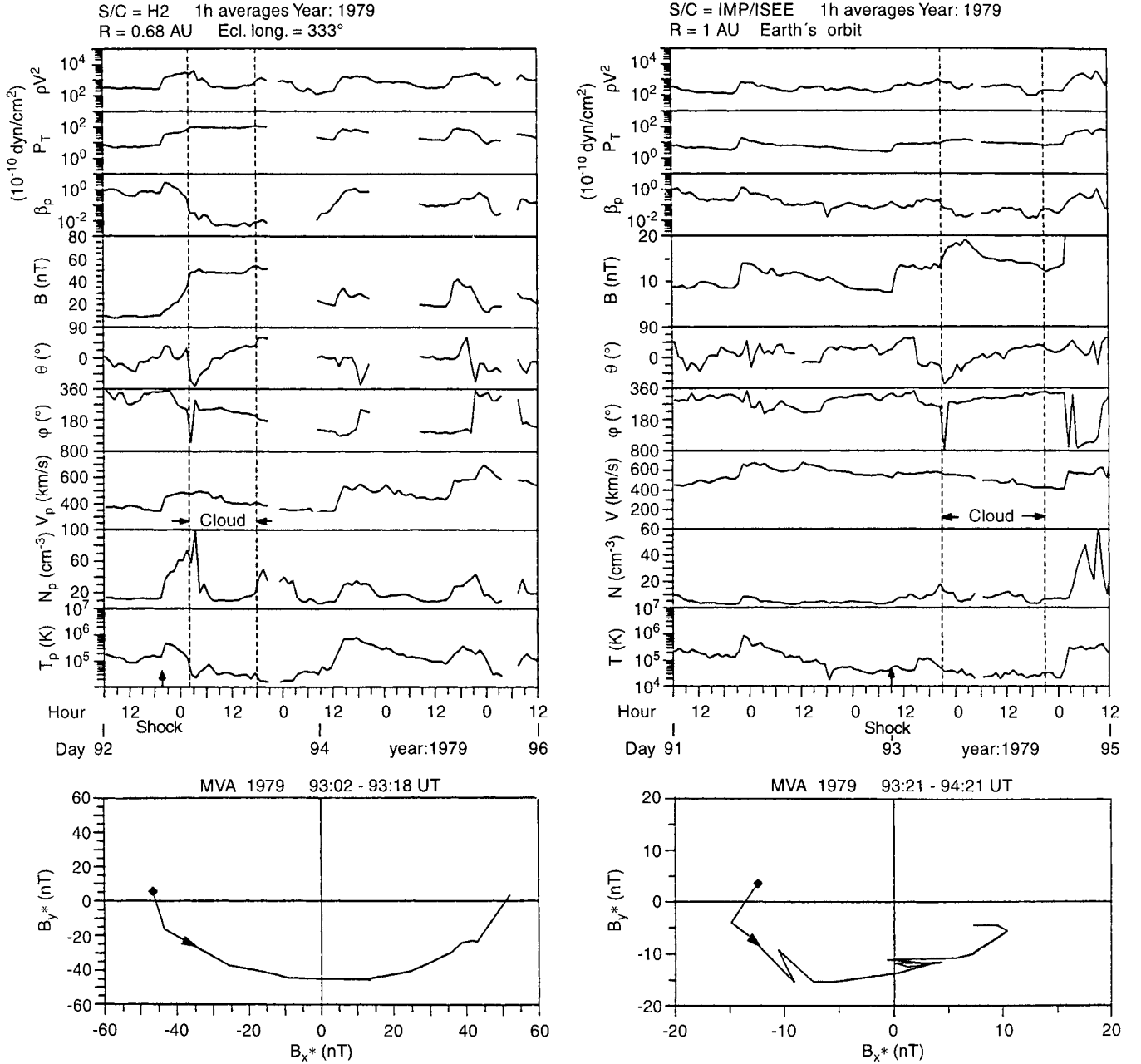


Fig. 10. Solar wind parameters of a MC (dashed lines) that passed Helios 2 at 0.7 AU, 27°E of Earth on day 93, 02–18 UT (left panel) and Earth's orbit on days 93, 21 UT – 94, 21 UT (right panel). Times

of the associated shocks: day 92, 20 UT at Helios 2; day 93, 09 UT at IMP/ISEE. Results from the MVA of the magnetic field data are displayed in the bottom panels

MC in the events in January and February 1978). Figure 10 gives an example. The left panel shows a MC that was observed by Helios 2 in April 1979 at 0.7 AU, 27° east of the Earth. Due to data gaps it was difficult to define the rear boundary of the MC unambiguously. The time interval that we identified as the MC is comparable to the start- and end-times given for this event by Burlaga *et al.* (1987).

This MC later passed Earth's orbit as shown by IMP/ISEE observations in the right panel. The MC was associated with an interplanetary shock and caused a strong geomagnetic storm with $K_p \geq 8$ - (see also Burlaga *et al.*, 1987). From the results of the MVA the cloud

was identified as a flux tube of type SWN. The spacecraft measurements of IMP/ISEE and that of Helios 2 show a striking similarity, although the MC was observed in different ambient solar wind streams. At Earth's orbit the MC was propagating through fast ambient wind, whereas it plowed through low speed wind at Helios 2. This observation is consistent with the slightly higher plasma speed of the MC at Earth's orbit ($V \sim 500$ km/sec), compared with the plasma speed measured at Helios 2 ($V \sim 450$ km/sec). Note that Helios 2 observed the MC at a sector boundary (see the change in ϕ from $\sim 315^\circ$ to $\sim 135^\circ$ in Fig. 10). The orientation of the cloud's axis as calculated from the minimum

variance method was $\varphi = 79^\circ$, $\theta = +14^\circ$ (time interval: 93:02–93:18 UT) at Helios 2 and $\varphi = 88^\circ$, $\theta = -12^\circ$ (time interval: 93:21–94:21 UT) at Earth's orbit with minimum variance directions of $\varphi = 164^\circ$, $\theta = -19^\circ$ at Helios 2 and $\varphi = 173^\circ$, $\theta = +22^\circ$ at Earth's orbit.

Figure 11 is a simple sketch for the possible large-scale configuration of this MC at the arrival time of the MC at Helios 2 based on the calculated axis orientations at the positions of Helios 2 and at Earth's orbit. The axis orientation of the MC at the heliolongitude of the Earth at the arrival time of the MC's center at Helios 2 was determined back by using the speed of the MC measured at 1 AU and the time delay between the MC's arrival at Helios 2 and 1 AU, assuming that the MC propagated along the Sun-Earth line with constant speed. Note that Helios 1 at 0.7 AU, 43° east of Helios 2, did observe the shock, but no signatures of a MC or ejecta in general could be found in the data.

In nine cases where a MC was detected from both Helios and Earth-bound satellites, longitudinal separations of up to $\sim 60^\circ$ occurred. We find that in cases with separations $> 60^\circ$, the MC was observed by only one spacecraft. Figure 12 shows the histogram obtained from our study. In two cases, in which both Helios s/c observed the same MC, we used the value of the spacecraft which was separated largest in longitude from Earth. Based on these observations the azimuthal broadness of a MC (60° in heliolongitude $\simeq 1$ AU at 1 AU) should be larger than its radial size (the typical diameter of a MC is ~ 0.25 AU at 1 AU). Figure 12 might imply a smaller extent for MCs in heliolongitude compared to the value of $\sim 100^\circ$ found previously for shock-associated IPCs (Richardson and Cane, 1993), but we have not found any case where a MC was observed with two spacecraft separated by more than 60° . A possible explanation for this discrepancy could be that the events studied by Richardson and Cane (1993)

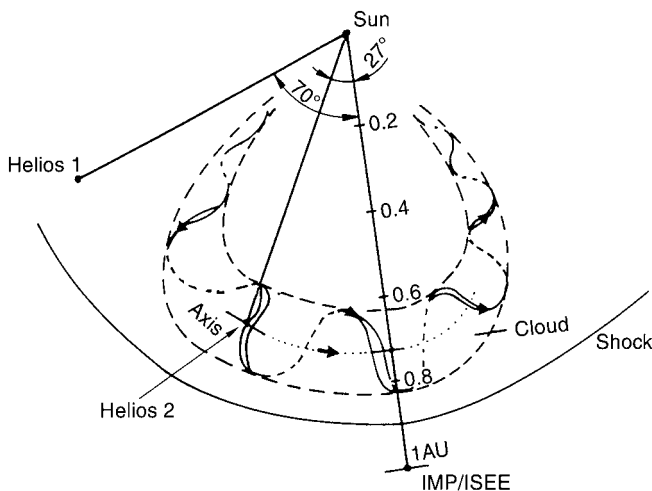


Fig. 11. Sketch showing the possible large-scale geometry of the MC observed by Helios 2 and IMP/ISEE in April 1979 (see Fig. 10) based on results of the MVA of the magnetic field data. Helios 1 did observe the shock, but not the MC. Arrows denote the orientation of the magnetic field lines at the cloud's outer boundaries and on its axis

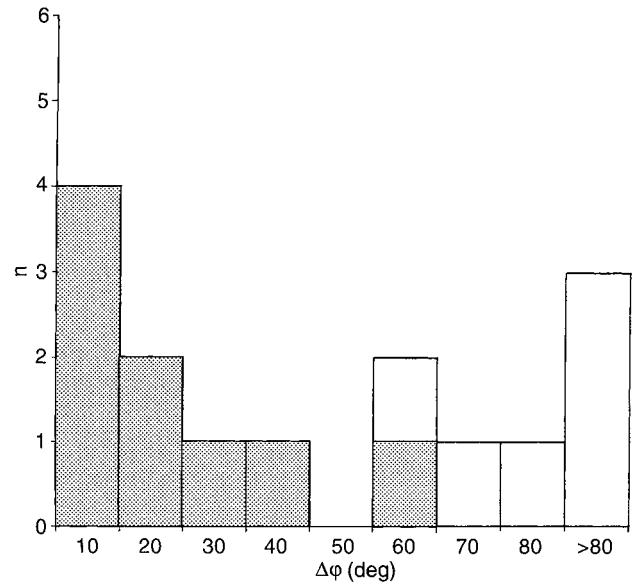


Fig. 12. Frequency distribution for the longitudinal span of MCs in the ecliptic based on correlated Helios and near Earth spacecraft observations. The spacecraft's longitudinal separation, $\Delta\varphi$, is given in intervals of 10° . Observations for longitude separations of more than 80° have been added together. Solid bars mark the number of events in which a MC passed both spacecraft separated by the given longitude interval in the ecliptic. Empty portions of the bars denote cases in which MCs were not observed by the s/c at Earth's orbit at the given angular separations from the Helios s/c

were all flare-associated and thus probably more energetic and larger in their spatial extent.

7 Expansion of Magnetic clouds

Burlaga and Behannon (1982) investigated MCs in the outer heliosphere using Voyager data. They found larger radial sizes for MCs observed further away from the Sun. To investigate the expansion of MCs in the heliosphere, we used the MCs observed by Helios 1/2 between 0.3–1 AU and investigated whether they had passed also the Voyager spacecraft in the outer heliosphere (s/c data other than Helios were taken from the COHO data base at NSSDC, described, e.g., in NSSDC News vol. 8(2), 1992). We found a number of cases in which the same MC could be uniquely identified in both data sets with time delays corresponding to the outward propagation of the MCs from the Helios to the Voyager spacecraft with approximately constant speed.

Here we present an example observed by Helios 1 at 0.9 AU and by Voyager 1/2 at 2.6 AU in March 1978 (for the identification of the MC in the Voyager data see also Burlaga and Behannon, 1982). The left side of Fig. 13 shows magnetic field data for Helios 1 (top panel) and Voyager 1/2 (middle, bottom panel), and the right side shows the corresponding results of the MVA. The spacecraft were separated by 18° in solar longitude and by 1.6 AU in solar distance. No signatures of a solar wind transient were found in the plasma data at Earth's orbit, 55° west of Helios 1, i.e., 73° west of Voyager 1/2,

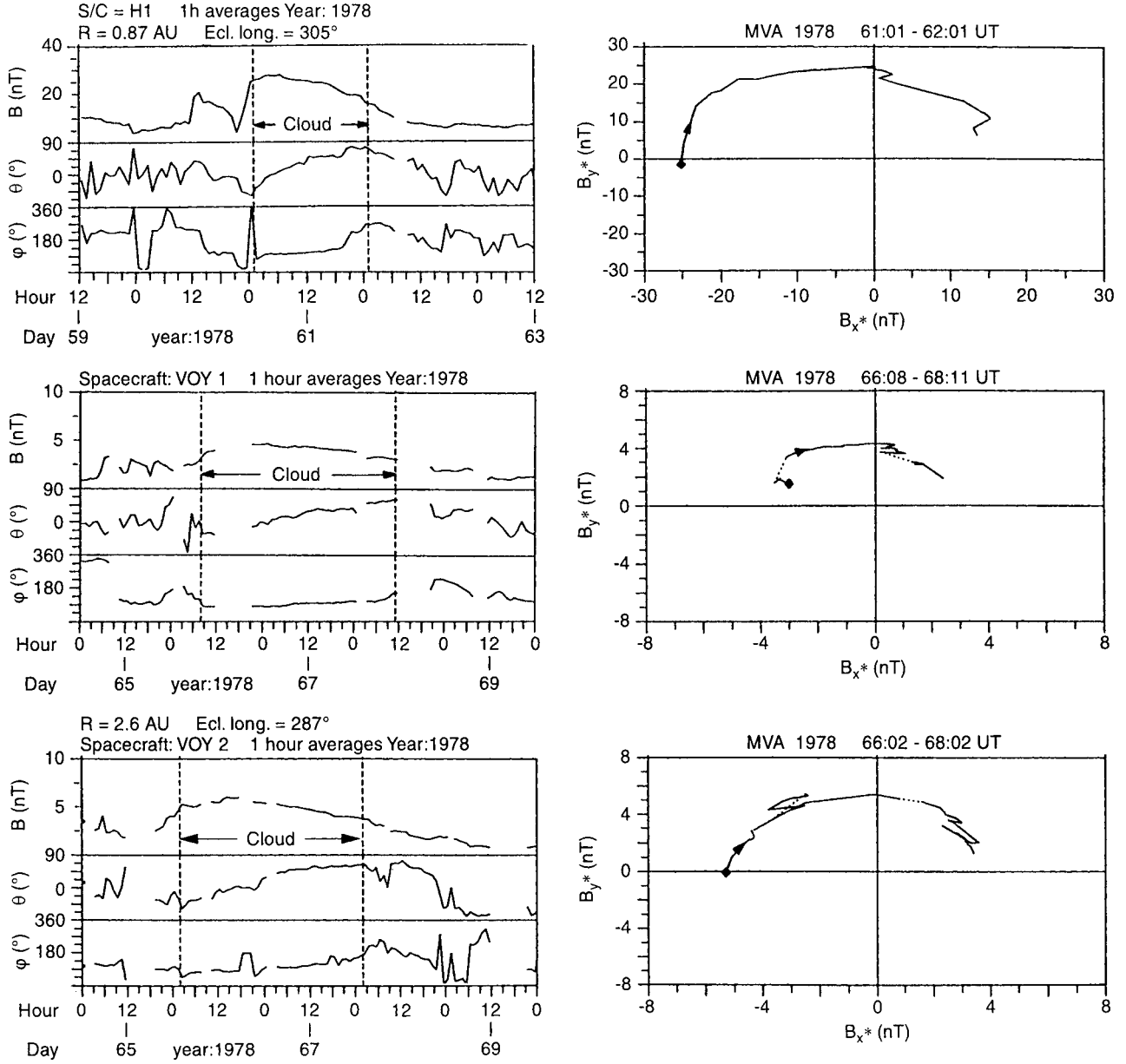


Fig. 13. Magnetic field data and results of the MVA for a MC (dashed lines) that passed Helios 1 at 0.9 AU on days 61, 01 UT – 62, 01 UT, Voyager 1 at 2.6 AU on days 66, 08 UT – 67, 11 UT and

Voyager 2 on days 66, 02 UT – 68, 02 UT in 1978. The Helios and Voyager spacecraft were separated by 18° in solar longitude. Note the longer time duration of the MC at Voyager 1/2

so that we can assume that there was also no MC at the Earth although there was a gap in the magnetic field data. Note the striking similarity of the directional changes of the magnetic field vector at all three spacecraft, which all detected a SEN-type MC. The arrival time of the MC at Voyager 1, at 2.6 AU (08 UT on day 66) and at Voyager 2 (02 UT on day 66) agreed well with the propagation time from Helios 1 to Voyager 1/2 as estimated from the plasma speed of ~ 446 km/sec observed during the cloud's passage over Helios 1 at 0.9 AU. The cloud showed a much longer time-duration at the position of Voyager 1 ($t = 50$ h) and Voyager 2 ($t = 48$ h) compared to Helios 1 ($t = 24$ h), thus giving direct evidence for the expansion of this MC. The MVA

yielded similar results for the orientation of the cloud's axis at all three spacecraft, with values of $\varphi = 119^\circ$, $\theta = +44^\circ$ at Helios 1, $\varphi = 91^\circ$, $\theta = +30^\circ$ at Voyager 1, and $\varphi = 83^\circ$, $\theta = +28^\circ$ at Voyager 2.

Figure 14 presents a log-log plot showing the growth of the radial sizes of MCs between 0.3 and 4.2 AU obtained from Helios 1/2, Voyager 1/2, and Pioneer 10 data (Bothmer and Schwenn, 1994). The smallest MCs had radial sizes < 0.1 AU and were observed close to the Sun, the largest MCs had radial sizes > 0.4 AU and were observed beyond 2 AU.

From linear regression it is found that the radial size, $s(R)$ in AU, of MCs increases with radial distance, R , from the Sun as:

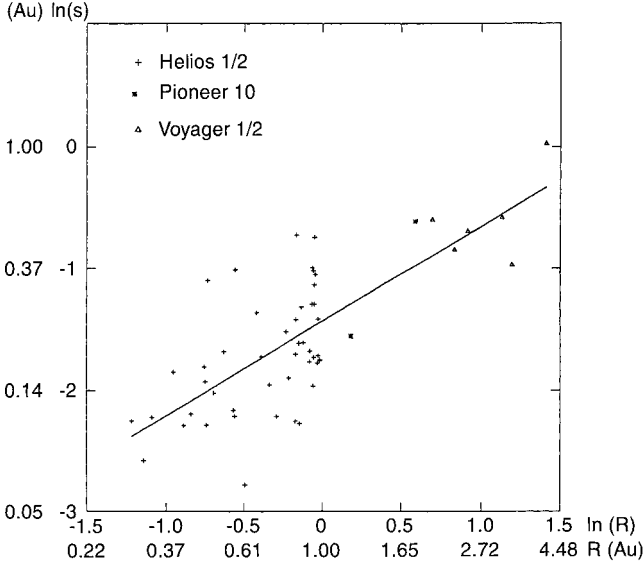


Fig. 14. Double logarithmic plot showing the radial sizes, s , of MCs observed by Helios 1/2 (+), Voyager 1/2 (Δ), Pioneer 10 (*) between 0.3 and 4.2 AU, versus solar distance, R (from Bothmer and Schwenn, 1994). The *straight line* was obtained by linear regression: $s(R) = (0.24 \pm 0.01) \times R^{0.78 \pm 0.10}$ [AU]

$$s(R) = (0.24 \pm 0.01) \times R^{(0.78 \pm 0.10)} (R \text{ in AU}) \quad (1)$$

There is a clear trend for an increase in the radial sizes of MCs with distance from the Sun, confirming that radial expansion is a common feature of MCs in the heliosphere.

MCs observed considerably closer to the Sun compared to 1 AU had often higher plasma densities than the surrounding solar wind (see Fig. 1). The density inside the flux tubes might fall off rapidly with increasing distance from the Sun as indicated by the radial expansion found for MCs. We determined the average proton density, $\langle N_p \rangle$, inside each of the MCs observed by the Helios probes. Figure 15 shows in a log-log plot $\ln(\langle N_p \rangle)$ versus $\ln(R)$. For comparison the decrease of the proton density measured in slow ($V < 400$ km/s; dashed line) and fast ($V > 600$ km/s; dotted line) solar wind streams over the range 0.3–1 AU has been included. These slopes were calculated according to results summarized in Schwenn (1990) with $\langle N_p \rangle = N_p \times R^{-2.1} \text{ cm}^{-3}$ (R in AU).

From linear regression we found that the density of MCs decreases with increasing radial distance, R , as:

$$N_p = (6.47 \pm 0.85) \times R^{(-2.4 \pm 0.3)} \text{ cm}^{-3} (R \text{ in AU}) \quad (2)$$

Thus the density inside MCs decreases stronger with radial distance from the Sun than that of the average solar wind which falls off as $R^{-2.1}$ in agreement with what one would expect for expanding structures. Figure 15 shows that the density in MCs is generally higher than that of the average fast solar wind and that it is also generally higher than that of the slow solar wind at closer distances to the Sun. Note that this result implies, without taking density enhancements into account that have been caused by compression due to

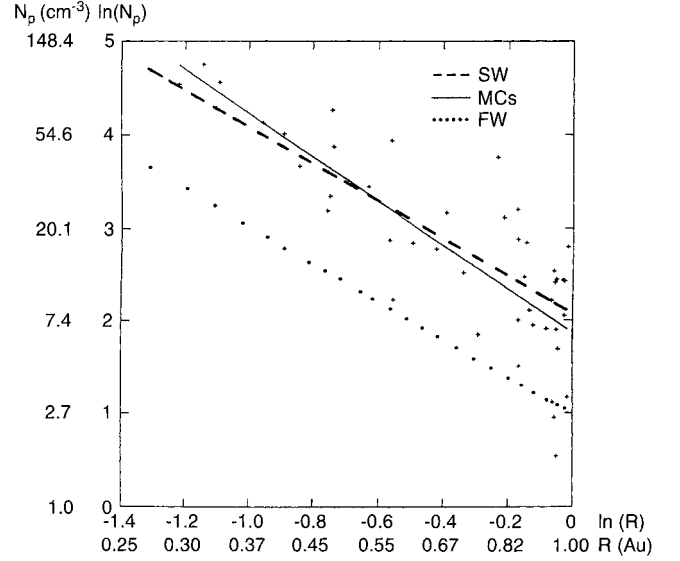


Fig. 15. Double logarithmic plot showing the average proton density inside MCs observed by Helios 1/2 between 0.3 and 1 AU during 1974–1981. The *straight line* was obtained by linear regression: $N_p(R) = (6.47 \pm 0.85) \times R^{-2.4 \pm 0.3} \text{ cm}^{-3}$ (R in AU). *Dashed* and *dotted lines* represent slopes for proton densities measured by the Helios probes in slow (S) and fast (F) solar wind-streams in the inner heliosphere (taken from Schwenn, 1991)

the interaction of a MC with the ambient solar wind, that IPC type MCs should be observed as transient white-light depletions rather than enhancements in the interplanetary medium at larger ($\gtrsim 0.5$ AU) distances from the Sun, in agreement with the suggested higher preference for MC associations with photometer IPCs in Table 1 for MCs observed closer to the Sun.

MCs in which the densities are found to be considerably lower compared to that of the ambient slow solar wind should have undergone strong expansion on their way out from the Sun. Consistent with this result, recent observations of the Ulysses spacecraft showed that IPCs including MCs were over-expanding at high latitudes in high speed solar wind due to their higher internal pressure compared to the surrounding flow (Gosling *et al.* 1994a,b).

To test whether our results are consistent with the flux-tube model for MCs, we assume that a MC has the topology of an ideal cylindrical flux tube of length l and diameter s , so that its volume V is proportional to $s^2 \times l$. We assume further mass is conserved inside the flux tube, i.e. $V \sim 1/N_p$. From Eq. (1) the diameter of the flux tube increases due to the expansion as $R^{0.78}$. Based on observations of prominence eruptions and white-light observations from space-coronagraphs (see e.g. Webb and Hundhausen, 1987; Webb, 1988) we assume that l increases linearly with R . Thus, the volume of the flux tube should increase as $R^{2.56}$, i.e., the density should decrease as $R^{-2.56}$. If we consider the simplicity of the model assumptions, we find this value to be in qualitative agreement with the $R^{-2.4}$ dependence in Eq. (2). The present results obtained for MCs observed in the ecliptic by Helios 1/2 are in agreement with the possible large-scale topology of a MC sketched in Fig. 11.

8 Magnetic clouds and disappearing filaments

Studies of the relationship of SMEs with other features of solar activity have shown that SMEs are better associated with prominence eruptions (disappearing filaments) than with solar flares (Harrison, 1986; Webb and Hundhausen, 1987). Yet, we do not know what causes SMEs nor which processes lead to the specific topology of the subset of SMEs associated with MCs in the interplanetary medium. A possible explanation is that a SME is caused by magnetic nonequilibrium leading to the rise of a prominence (e.g., Hundhausen *et al.* 1994; Priest, 1988; Feynman and Martin, 1995). The typical three-part structure of SMEs seen in white-light images (see, e.g., Webb, 1988 and references therein) is interpreted as a leading bright front comprised of coronal field lines with a cavity beneath it, followed by the rising prominence in the trailing portion of the SME.

Erupting quiescent prominences often show a magnetic flux-rope topology, so that it may seem plausible to speculate whether they are related to MCs (the term “eruption” means here that a prominence disappears in the H α -observations, i.e., an erupting prominence might remain attached to the Sun). The concept of force-free magnetic flux tubes was first introduced by Lüst and Schlüter (1954) for quiescent solar prominences and was later applied to MCs (Goldstein, 1983; see also the review on magnetic flux ropes given by Priest, 1990 and the paper on “theory of prominence eruption and propagation” by Chen, 1996).

The ‘Catalog of Solar Filament Disappearances 1964–1980’ (Wright, 1991) offered us the most useful data set to investigate whether quiescent filaments

(disparition brusques, DBs) and MCs can be associated and to compare their magnetic structures (see also Bothmer, 1993; Rust, 1994). In order to establish unique associations between MCs that were observed in situ in interplanetary space by the Helios probes and DBs observed from Earth we used the following criteria: (1) Since the observations of DBs were taken from Earth we did not include any MC that was observed by the Helios probes when it was positioned further away than 60° in solar longitude from the Sun’s central meridian (as seen from Earth), in order to avoid DB-events which might have gone undetected due to limb darkening or because the DB might have occurred behind the Sun’s limbs which reduced the number of events to 25 cases. (2) We regarded only those associations of MCs and DBs as safe when spacecraft and filament were separated not more than $\sim 20^\circ$ in solar longitude, and when the DBs had disappeared at the Sun in time windows (including the error in the given DB disappearance) that were calculated from the maximum and minimum proton speeds measured inside the MCs.

These criteria led to the 10 MC events listed in Table 4. In five cases, MCs could be confidentially associated with a DB at the Sun and we compared their magnetic structures (see also Bothmer and Schwenn, 1994). The solar latitude of the filaments was ~ 30 – 60° (see Table 4). We remark that for the first two associations in Table 4 more than one filament disappeared during the estimated time windows for the solar events. In these cases the DB closest in longitude to the Helios probes has been selected. It is interesting to note that the daily rate of DBs between 45°E and 45°W was only 0.1–0.4 during 1975–1980.

Table 4. Associations between MCs observed with Helios 1/2 and disappearing filaments (DBs). Left to right: date of the MC, s/c, s/c coordinates, minimum and maximum measured proton speeds inside the MC, calculated time-window for the solar event, disappearance of the DB at the Sun, solar position of the DB, magnetic polarity & inferred flux tube type of the DB, in-situ observed type of the MC (after Bothmer and Schwenn, 1994)

MC (year: month: day)	S/C	$R[\text{AU}]$ Long, Lat. [deg]	V_{\min}, V_{\max} [km/s]	Time-window start-end (month: day: hour)	DB (month: day: hour)	Solar position [degree]	Magnetic polarity, type	MC type
75:03:04	H1	0.39 W30, S7	418, 507	03:03:00 03:03:07	no			
77:12:01	H1	0.75 E36, N5	401, 431	11:28:07 11:28:12	11:28:06	N60E65 N50E33	–/+ SEN	SEN
78:01:17	H2	0.98 E09, S4	305, 340	01:11:10 01:12:00	01:11:09	S37W01 S10W17	–/+ NES	NES
78:01:29	H2	0.98 E13, S5	475, 549	01:25:21 01:26:09	01:26:06	N55E25 N39E08	–/+ SEN	SEN
78:02:06	H2	0.98 E16, S5	367, 444	02:02:00 02:02:19	no			
78:02:16	H2	0.95 E19, S6	385, 491	02:11:19 02:12:17	02:11:15	S39E45 S37E26	–/+ NES	SWN
78:03:02	H1	0.87 E55, S4	376, 502	02:26:00 02:27:00	no			
78:04:02	H2	0.61 E17, S7	420, 441	03:30:12 03:30:15	03:30:03	S60E39 S45E09	+/- SWN	SWN
79:04:03	H2	0.68 E27, S7	389, 496	03:31:00 03:31:16	no			
79:12:31	H2	0.84 E15, S1	375, 442	12:27:02 12:27:16	no			

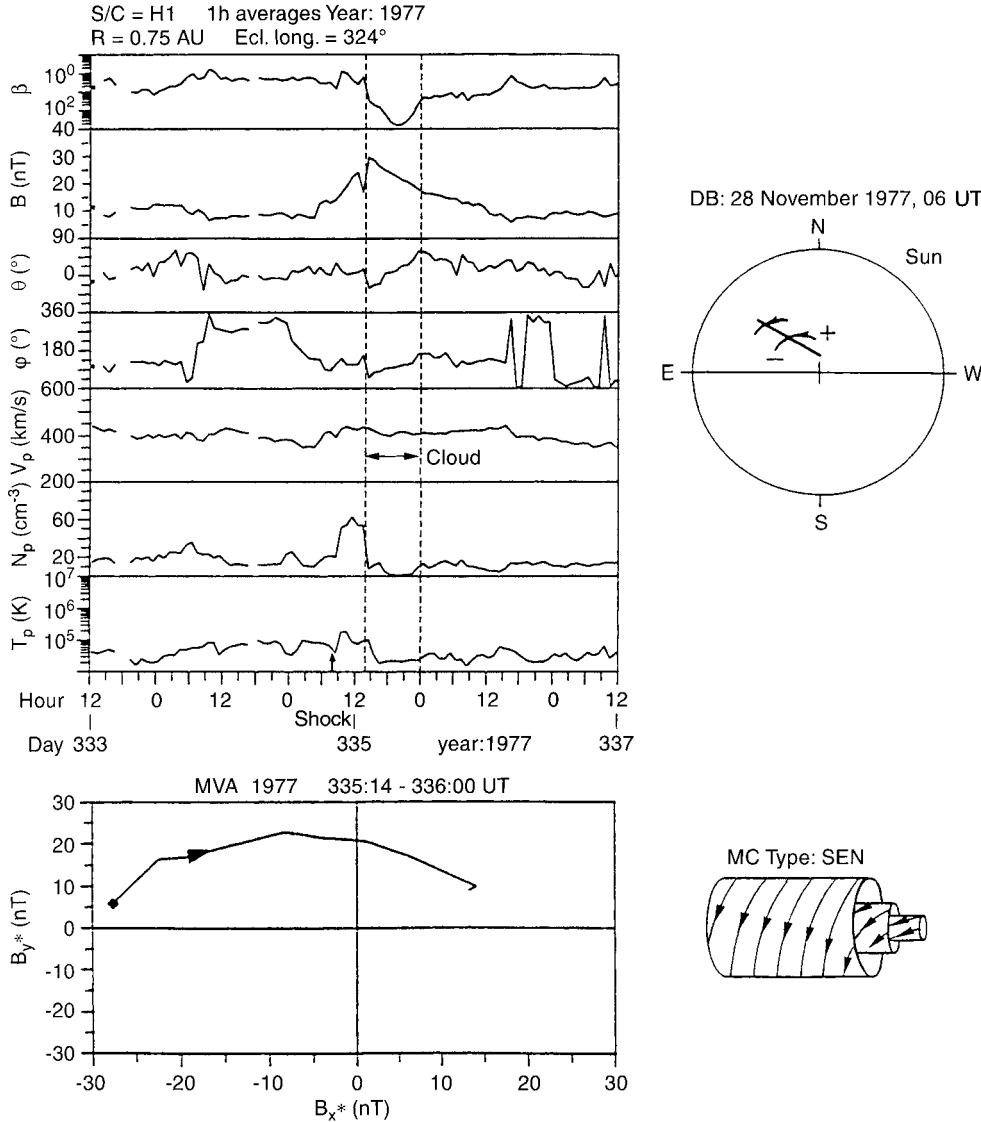


Fig. 16. An MC observed by Helios 1 at 0.8 AU (day 335, 14 UT – day 336, 0 UT) in December 1977 and the associated filament at the Sun. *Upper left part:* solar wind and plasma parameters for the MC. *Upper right part:* orientation and polarity of the filament (DB) at the Sun (see Table 4). *Left bottom panel:* results from the MVA of the magnetic field data inside the MC. *Right bottom panel:* in situ observed flux tube type of the MC

Figure 16 shows one of the five MCs of Table 4 that was associated with a DB. This MC was observed by Helios 1 in December 1977 at 0.8 AU. The figure shows the magnetic flux tube type as inferred from the MVA results and the inferred magnetic structure of the associated DB. The magnetic structure of the filament has been inferred from the polarity of the magnetic field with respect to the prominence axis based on the observation that the field structure in quiescent prominences is composed of loop-like arcades overlying the prominence axis (neutral line) (see, e.g., Dynamics and Structure of Quiescent Solar Prominences, edited by Priest, 1989), by taking into account the hemispheric magnetic patterns found for mid- to high-latitude filaments (Rust, 1967; Martin *et al.*, 1994; Rust and Kumar, 1994a) and by applying the helicity rule found by Rust and Kumar (1994a) who inferred that for prominences in the northern solar hemisphere, the twist of field lines is preferentially left-handed, whereas it is preferentially right-handed for prominences in the southern hemisphere (compare with Fig. 18). This

finding is consistent with sunspot observations which show that anti-clockwise whorls are three times as common in the north and clockwise ones in the south (Richardson, 1941; Yang *et al.*, 1988). Anti-clockwise whorls are related to dextral and clockwise ones to sinistral filaments. It is interesting to note that the Helios observations showed SN-MCs to be three times more frequent than NS-MCs which is similar to the ratio of clockwise and anti-clockwise whorls in sunspots (Richardson, 1941).

The agreement of the magnetic topologies of the MC and the DB in Fig. 16 is striking. The individual flux-tube types given for the DBs in Table 4 have been labeled according to the notation that we introduced for MCs in Section 4 of this study (to compare with Table 2). Similar agreements between the magnetic structure of the filament and that of the associated MC were found in four of the five events listed in Table 4, with three different types of flux-tubes. In the one event in which we have not found an agreement, the MC had a complex internal magnetic field structure.

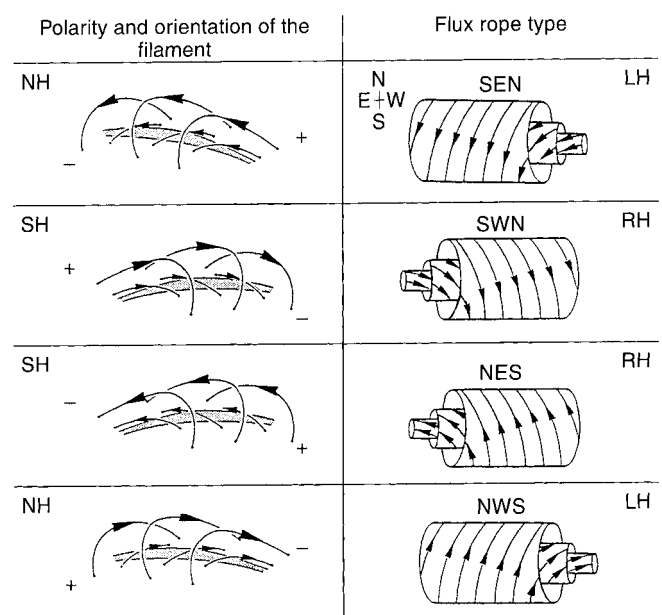
Table 5. Associations between MCs observed with near Earth spacecraft (OMNI database) and DBs during 1978–1980. Left to right: date of MC, speed range of MC, date of associated DB, solar position of DB, magnetic polarity of DB, inferred DB flux tube type, MC flux tube type

MC (y.mo.:d)	V (km/s)	DB (month:day)	Solar position	Magnetic polarity	Type	MC type
78:01:17	300–350	01:11	S37W01 S10W17	−/+	NES	NES
78:04:03	450–500	03:30	S60E39 S45E09	+/-	SWN	SWN
78:08:27	400–450	08:23	N16E18 N11E04	+/-	NWS	NWS
79:04:25	500–600	04:22	S40E29 S36E07	+/-	SWN	SWN
80:02:16	350–450	02:11	N19E02 N17W22	−/+	SEN	SEN
80:03:19	300–400	03:16	S48W10 S53W32	−/+	NES	NES

In the same way as for the Helios data, we have searched near Earth solar wind data from the OMNI and COHO data-bases at NSSDC for the years 1978–1980 when the data coverage was very good. We identified 12 MCs during these years. In four events DB observations were missing, one MC was not associated with a DB and in one event the filament was separated substantially in longitude from central meridian. Table 5 lists the six events that were found to be associated with a DB using the same criteria as for the Helios observations. Note that the first two events in Table 5 are MCs that have also been observed with the Helios probes so that Table 5 yields a total set of four new associations between MCs and DBs. We found an agreement between the magnetic flux tube type of the MC and that of the associated DB in all cases. Similar to the Helios observations, the filaments associated with the MCs, were observed typically at higher latitudes, in agreement with the recent findings by Bothmer and Rust (1997) based on an extended study of MCs observed at 1 AU and DBs.

Our results are summarized in Fig. 17 (adapted from Bothmer and Schwenn, 1994): four different magnetic flux tubes are expected in interplanetary space from the magnetic structure of disappearing filaments in the Sun's southern and northern hemispheres. The results indicate that the magnetic helicity of MCs observed in interplanetary space is the same as that of the associated filaments, i.e., the magnetic helicity inherent in magnetic fields at the Sun might be conserved outward into interplanetary space, as has been pointed out by Rust and Kumar (1994b) in agreement with findings of helicity conservation in laboratory plasmas. Note that on the contrary, Martin and McAllister (1997) claimed that the helicity of the overlying arches above a prominence axis which may evolve into flux tube is opposite to that of the filament beneath it.

Magnetic flux ropes/tubes may be formed through magnetic reconnection of coronal field lines in a rising SME as has been pointed out by Gosling (1990) and Gosling *et al.* (1995c,d). In support of this view, the Ulysses spacecraft has directly observed MCs and IPCs after formation of new coronal loops in solar regions that map to the spacecraft, as observed by the Japanese Yohkoh-satellite (Gosling *et al.*, 1995c; Bothmer *et al.*, 1996b; Lemen *et al.*, 1996; Weiss *et al.*, 1996). We believe that our findings are further evidence for an evolution of arcades of coronal loops into magnetic flux tubes (MCs). Weiss *et al.* (1996) found no characteristic

**Fig. 17.** Inferred magnetic structure of filaments in the Sun's northern and southern hemispheres (NH, SH) and that of associated MCs in interplanetary space (adapted from Bothmer and Schwenn, 1994). The abbreviations for the magnetic field structures have been chosen in analogy to the different flux tube types of MCs presented in Table 2. LH (left-handed) and RH (right-handed) denote the corresponding magnetic helicity

differences in the coronal soft X-ray signatures of IPC- and MC-type IPC events, but they pointed out that the interplanetary plasma- β of SMEs appears to be crucial in determining the magnetic field evolution. This is consistent with our finding of low plasma- β values in MCs between 0.3–1 AU. It is further interesting to remark that the magnetic fields of quiescent filaments are presumably low plasma- β regions in the solar atmosphere (see, e.g., Priest, 1989).

We found evidence for a solar cycle variation of the magnetic structure of MCs in Sec. 5.2. In Fig. 18 (Bothmer and Rust, 1997) we propose an extended association between the polarity of sunspots, the magnetic structure of filaments and that of MCs, that could explain our observations.

It is well known that H α filaments commonly occur in regions where sunspots are no longer visible. Usually the preceding sunspot (in the sense of solar rotation) lies closer to the solar equator than the following one which is of opposite magnetic polarity. The left column of

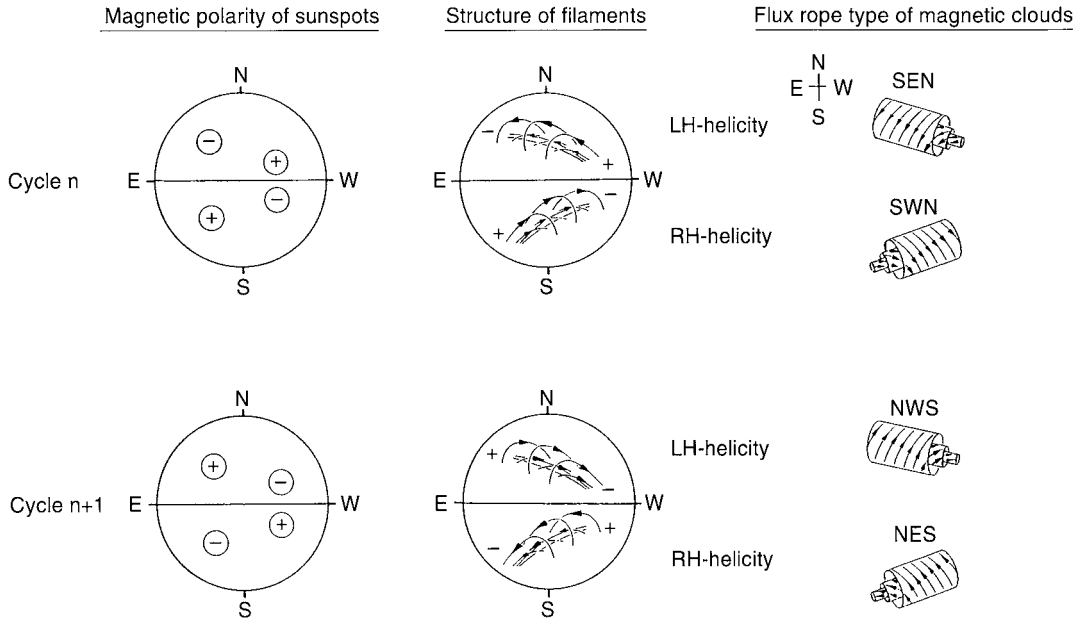


Fig. 18. Sketch showing the extended association between sunspots, filaments and MCs. The *left panel* shows the polarity of sunspots during two subsequent cycles. Note that the solar cycle number, n , is assumed to be an odd number. The middle panel shows the expected preferential magnetic structure of filaments in the two solar

hemispheres and the *right panel* shows the expected structure of MCs in interplanetary space. During odd (n) cycles SN-MCs are dominant whereas NS-MCs dominate during even ($n+1$) cycles (Bothmer and Rust, 1997)

Fig. 18 shows the magnetic polarities of sunspots in both solar hemispheres during two subsequent cycles. The top row of Fig. 18 shows the magnetic polarities for odd cycles, e.g., for cycle 21, when the Helios observations were taken. The bottom row shows the polarities during even cycles. The middle column sketches the magnetic structures for mid- to high-latitude filaments during the two cycles based on the hemispheric helicity rule found by Rust and Kumar (1994) who inferred that in filaments in the northern solar hemisphere the twist of the magnetic field lines is preferentially left-handed, whereas it is preferentially right-handed in the southern hemisphere. The right column of Fig. 18 finally presents the magnetic structure of MCs based on the interpretation of our observations (see also Bothmer and Schwenn, 1994).

Figure 18 implies that during odd cycles, e.g., during solar cycle 21, filaments should preferentially evolve in association with SMEs into SN flux tubes (MCs) in the interplanetary medium, whereas in the following cycle preferentially NS flux tubes (MCs) should evolve, in agreement with our findings. Further evidence for the solar cycle dependence of MCs and their agreement with the magnetic structure of filaments based on the helicity rule has been recently presented by Bothmer and Rust (1997).

9 Conclusions and summary

Using the complete set of plasma and magnetic field data of the Helios 1/2 spacecraft, we have identified 46 MCs between 0.3–1 AU during the period 1974–1981

between 0.3 and 1 AU. Additionally, we used solar wind data from the IMP, Pioneer and Voyager spacecraft for multi-spacecraft studies of MCs and to investigate the expansion of MCs in the outer heliosphere. Based on these investigations, we summarize the main properties of MCs in the heliosphere:

1. Our results support the close relationship between MCs and SMEs. 85% of the MCs observed by Helios 1/2 were associated with a fast forward shock. Seven of the MCs were observed after a Helios-directed SME was detected by the Solwind-coronagraph on the satellite P78/1. Without taking the duty cycle of the instruments into account, half of the MCs were found to be associated with passage of an IPC detected by the Helios-photometer as a white-light brightness enhancement in the interplanetary medium.
2. Between 0.3 and 1 AU, MCs have higher total plasma pressure and lower plasma- β values ($\beta \ll 1$) compared to the normal solar wind.
3. MCs can best be described as large-scale cylindrical magnetic flux tubes with observed longitudinal extents up to $\sim 60^\circ$. The axis of the flux tube commonly has a small inclination to the ecliptic, with its azimuthal direction being roughly directed near the EW-direction, i.e., normal to the sun-spacecraft line.
4. MCs can occur in four different magnetic configurations. The individual types are distinguished by the magnetic orientation at the center (the axis) of the MC and on its outer boundaries. The four configurations result from two possible directions of the field axis and two possible values of magnetic helicity.
5. MCs expand during their outward propagation. Their radial diameter, s , was found to increase between 0.3

and 4.2 AU proportional to $R^{0.8}$. Consistent with this expansion, the plasma density inside MCs decreases, faster than the normal solar wind, proportional to $R^{-2.4}$ between 0.3 and 1 AU. According to this result, on average, an IPC type MC should be identifiable as a transient white-light brightness enhancement in the interplanetary medium up to a distance of ~ 0.5 AU from the Sun, whereas, without taking brightness enhancements into account that have been caused by compression due to the interaction of a MC with the ambient solar wind, a MC should at greater distances from the Sun lead to a transient white-light brightness depletion.

6. We found nine MCs identified in Helios and near-Earth solar wind data to be uniquely associated with quiescent filament disappearances (DBs) at mid to high solar latitudes outside active regions. In eight of the nine cases we found agreement between the magnetic flux tube structure of the MC and that of the associated filament as inferred from the orientation of the filament axis and its magnetic polarity, on the assumption of left-handed (right-handed) magnetic helicity dominance of filaments in the northern (southern) solar hemisphere. Further, the preferred east-west orientation of mid- to high-latitude quiescent prominences caused by differential rotation may explain the preference of MCs to have small axis inclinations to the ecliptic in the interplanetary medium. However, we note that some MCs may well be related to other solar activity features, such as flares. Also, other results do not necessarily imply that the filament itself evolves into the MC in interplanetary space. From a comparison of the spatial sizes of the features seen in H α -filtergrams and white-light coronagraph observations it seems likely that the filament may be either the bottom part of a large flux tube or just an indicator of the overlying coronal field structure.
7. There is a solar cycle dependence of the magnetic structure of MCs. Although MCs of opposite magnetic helicity occurred with equal frequency during 1974–1981, $\sim 3/4$ of the MCs were associated with SN-rotations of the magnetic field vector, but only $\sim 1/4$ had NS-rotations. During the subsequent cycle predominately NS-rotations were found.
8. We suggest that the solar cycle variation of the magnetic structure of MCs can be explained by the magnetic field structure of disappearing filaments at the Sun and their solar cycle changes.
9. It is likely that magnetic reconnection occurs in the coronal loops overlying the prominence axis of a disappearing filament in association with a SME.

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Appendix: minimum variance analysis for MCs

During the passage of a MC, the magnetic field vector B was measured at N successive times t . The mean value of B in Cartesian components is:

$$\langle B \rangle = \frac{1}{N} \sum_{i=1}^N B^i \quad (A1)$$

with: $B^i = (B_x^i, B_y^i, B_z^i)$, $i = 1, \dots, N$

A normal-vector n for the direction of minimum variance of B is to be determined. This can be achieved by calculating the mean quadratic deviation of the individual products $B^i \cdot n$ from $\langle B^i \rangle \cdot n$:

$$\sigma^2 = \frac{1}{N} \sum_{i=1}^N (B^i \cdot n - \langle B^i \rangle \cdot n)^2 \quad (A2)$$

Optimizing Eq. (A2) is equivalent to finding the smallest eigenvalue of the covariance-matrix $M_{\alpha,\beta}$:

$$M_{\alpha,\beta} = (\langle B_\alpha B_\beta \rangle - \langle B_\alpha \rangle \langle B_\beta \rangle) \quad (A3)$$

with: $\alpha, \beta \in \{x, y, z\}$

The individual $B_{\alpha,\beta}$ in Eq. (A3) are the Cartesian components of B for a single measurement. The three calculated eigenvectors and eigenvalues of $M_{\alpha,\beta}$ correspond to the directions of minimum, intermediate and maximum variance of B , i.e. to the directions of a new principal axis (the minimum variance) system. The common nomenclature in the literature for the eigenvectors \underline{e}_i and eigenvalues λ_i is:

- e_1, λ_1 : eigenvector and eigenvalue for the direction of maximum variance
- e_2, λ_2 : eigenvector and eigenvalue for the direction of intermediate variance
- e_3, λ_3 : eigenvector and eigenvalue for the direction of minimum variance

The variance directions are well determined if the error criteria given by Lepping and Behannon (1980) and Siscoe and Suey (1972) are satisfied, i.e. if:

$$\frac{\lambda_2}{\lambda_3} \geq 2 \quad (A4)$$

$$\angle(B_1, B_N) \geq 30^\circ$$

The approximate error of the minimum variance direction is $\sim 10^\circ$ (Burlaga and Behannon, 1982). The calculated variance directions in solar ecliptic coordinates are:

$$\varphi_k = \text{atan} \frac{y_k}{x_k} \quad (A5)$$

$$\theta_k = \text{atan} \frac{\sqrt{x_k^2 + y_k^2}}{z_k}$$

with: $|e| = 1$; $e_k = (x_k, y_k, z_k)$; $k = 1, 2, 3$

Finally, the measured components of B in the minimum variance system are:

$$B_k^i = e_k \cdot B^i \quad (\text{A6})$$

with:

$$\begin{aligned} B_1^* &= B_x^* \equiv \text{component of maximum variance} \\ B_2^* &= B_y^* \equiv \text{component of intermediate variance} \\ B_3^* &= B_z^* \equiv \text{component of minimum variance} \end{aligned}$$

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