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# Subauroral red arcs as a conjugate phenomenon: comparison of OV1-10 satellite data with numerical calculations

A. V. Pavlov

Institute of Terrestrial Magnetism, Ionosphere and Radio-Wave Propagation, Russia Academy of Science. (IZMIRAN), Troitsk, Moscow Region, 142092, Russia

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**Abstract.** This study compares the OV1-10 satellite measurements of the integral airglow intensities at 630 nm in the SAR arc regions observed in the northern and southern hemisphere as a conjugate phenomenon, with the model results obtained using the time-dependent one-dimensional mathematical model of the Earth ionosphere and plasmasphere (the IZMIRAN model) during the geomagnetic storm of the period 15–17 February 1967. The major enhancements to the IZMIRAN model developed in this study are the inclusion of  $\text{He}^+$  ions (three major ions:  $\text{O}^+$ ,  $\text{H}^+$ , and  $\text{He}^+$ , and three ion temperatures), the updated photochemistry and energy balance equations for ions and electrons, the diffusion of  $\text{NO}^+$  and  $\text{O}_2^+$  ions and  $\text{O}(^1\text{D})$  and the revised electron cooling rates arising from their collisions with unexcited  $\text{N}_2$ ,  $\text{O}_2$  molecules and  $\text{N}_2$  molecules at the first vibrational level. The updated model includes the option to use the models of the Boltzmann or non-Boltzmann distributions of vibrationally excited molecular nitrogen. Deviations from the Boltzmann distribution for the first five vibrational levels of  $\text{N}_2$  were calculated. The calculated distribution is highly non-Boltzmann at vibrational levels  $v > 2$  and leads to a decrease in the calculated electron density and integral intensity at 630 nm in the northern and southern hemispheres in comparison with the electron density and integral intensity calculated using the Boltzmann vibrational distribution of  $\text{N}_2$ . It is found that the intensity at 630 nm is very sensitive to the oxygen number densities. Good agreement between the modeled and measured intensities is obtained provided that at all altitudes of the southern hemisphere a reduction of about factor 1.35 in MSIS-86 atomic oxygen densities is included in the IZMIRAN model with the non-Boltzmann vibrational distribution of  $\text{N}_2$ . The effect of using of the  $\text{O}(^1\text{D})$  diffusion results in the decrease of 4–6% in the calculated integral intensity of the northern hemisphere and 7–13% in the calculated integral intensity of the southern hemisphere. It is found that the modeled intensities of the southern hemisphere are more sensitive to the assumed values of the rate coefficients of  $\text{O}^+(^4\text{S})$

ions with the vibrationally excited nitrogen molecules and quenching of  $\text{O}^+(^2\text{D})$  by atomic oxygen than the modeled intensities of the northern hemisphere.

## 1 Introduction

During large magnetic storms characterized by a geomagnetic planetary index  $K_p$  of 5 or greater, stable auroral red (SAR) arcs are observed optically on the equatorward edge of the mid-latitude ionospheric trough in both the northern and southern hemispheres on approximately the same L-shell value of the geomagnetic field (Rees and Roble, 1975). A major part of ground and satellite observations that determine the properties of the SAR arcs have been obtained in the northern hemisphere. The observed SAR-arc properties in magnetically conjugate hemispheres were discussed only by LaValle and Elliott (1972) using OV1-10 satellite data on 16 February 1967, and Reed and Blamont (1968), who determined the positions of the SAR arcs on 28–29 September 1967 using data from the satellite Ogo 4. In this paper we present the results of the first calculations and studies of the SAR arc properties in magnetically conjugate hemispheres for 16–17 February 1967.

The ring current is a prime candidate as the energy source responsible for maintenance of 630-nm emission (Cole, 1965; Rees and Roble, 1975; Kozyra and Nagy, 1991). Several excitation mechanisms have been proposed to account for the heating of the electron gas in the SAR arc region (Cole, 1965; Cornwall et al., 1971; Rees and Roble, 1975; Kozyra et al., 1987; Kozyra and Nagy, 1991; Thorne and Horne, 1992; Horne and Thorne, 1993). Unfortunately, the theory of these mechanisms is very complicated and we cannot use simple equations for the additional heating rate of the electron gas. To model the SAR arc during the 18 December 1971 magnetic storm, a simple model, without any additional heating rate of the electron gas

during the main phase and with the constant additional heating rate during the recovery phase of geomagnetic storm, was employed by Pavlov (1996). This approach with some modification is used in our model to carry out a comprehensive study of the response of the ionosphere and the integral airglow intensities,  $I_{630}$ , at 630 nm in the SAR arc regions to the 15–17 February 1967 geomagnetic storm.

In earlier theoretical studies, Newton *et al.* (1974) and Pavlov (1989) found that deviation from a Boltzmann distribution of  $N_2(v)$  was large in the SAR-arc region at vibrational energy levels  $v > 2$ , and this deviation significantly affected the loss rate of the  $O^+(^4S)$  ions at F2 region altitudes. The model used in this study gives an opportunity to examine the effect of a non-Boltzmann distribution of  $N_2(v)$  on the electron densities and temperatures in the SAR-arc regions of magnetically conjugate hemispheres with the result in the integral airglow intensities at 630 nm.

It has been suggested in a large number of modeling studies that semi-empirical models of the neutral atmosphere (e.g., MSIS) misrepresent true atomic oxygen and molecular nitrogen number densities and neutral temperatures in SAR-arc regions (Kozyra *et al.*, 1982, 1990; Slater *et al.*, 1987; Pavlov, 1996). In this paper we examine the sensitivity of the model results to important atmospheric parameters of the northern and southern hemispheres.

Previous calculations of  $O(^1D)$  number densities have assumed the formation of  $O(^1D)$  in the atmosphere due to production and loss reactions of  $O(^1D)$ , without taking into account the diffusion of  $O(^1D)$ . An important aim of this paper is to study the role of the diffusion of  $O(^1D)$  in the SAR arc regions.

## 2 Theoretical model

The mathematical model of the thermal plasma in the ionosphere and plasmasphere (the enhanced and improved IZMIRAN model) employed in our calculations is presented in the Appendix. This model includes the option to use models A, B, or C of the rate coefficients of  $O^+(^4S)$  ions with the vibrationally excited nitrogen molecules and quenching of  $O^+(^2D)$  by atomic oxygen described in the Appendix. In this study we employ model A and compare results with those from models B and C to evaluate the effects of the uncertainties in these rate coefficients on calculated integral airglow intensities at 630 nm in the SAR-arc regions. We have no information on hmF2 and thus can use neither the method developed by Richards (1991) nor our modification of it (Pavlov and Buonsanto, 1997) to calculate an equivalent plasma drift velocity. To overcome this problem, we use the approach of Eqs. (A4–A7) described in the Appendix. In this paper we will examine the effect of the non-Boltzmann distribution of  $N_2(v)$  (by solving the system Eq. (A19–A20) of the Appendix) on the integral airglow intensities at 630 nm in the SAR-arc regions in comparison with the Boltzmann distribution of  $N_2(v)$  given by Eq. (A23) of the Appendix.

One of the inputs for the IZMIRAN model is the additional heating rate,  $P_{rc}$ , of the electron gas in Eq. (A10) of the Appendix. At the present time, wave-particle interactions (Thorne and Horne, 1992; Erlandson *et al.*, 1993; Horne and Thorne, 1993) and Coulomb collisions between ring-current ions (Cole, 1965; Kozyra *et al.*, 1987; Slater *et al.*, 1987; Kozyra and Nagy, 1991; Fok *et al.*, 1995a,b, 1993) and plasmaspheric electrons and ions are the most plausible mechanisms used to explain electron temperature enhancements and the associated SAR arcs. The present work addresses the contribution from this latter process.

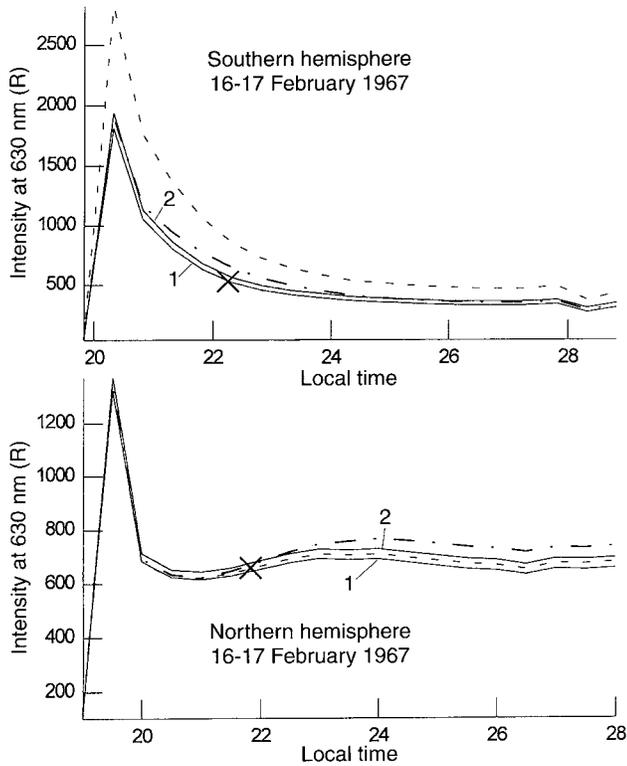
The rate of energy transfer from the energetic ring-current ions to plasmaspheric electrons due to Coulomb collisions has been discussed extensively by Kozyra *et al.* (1987). There is no information available on the ring-current population for this storm, and we are therefore forced to assume a heating rate function. It can be found from Eq. (4) in Kozyra *et al.* (1987) that the value of  $P_{rc}$  is proportional to  $N_e \cdot T_e^{-1/2}$ , where  $N_e$  is the electron number density and  $T_e$  is the electron temperature of the plasmasphere. Therefore we selected the additional heating rate as

$$P_{rc}(s, t) = f(s, t) \cdot C_0 \cdot N_e(s, t) \cdot T_e(s, t)^{-1/2}, \quad (1)$$

where  $f(s, t) = 0$  for  $s > s_0$  and  $s < -s_0$ ;  $f(s, t) = 0$  for  $t < t_0$ ;  $f(s, t) = 1$  for  $-s_0 < s < s_0$  and  $t \geq t_0$ ;  $t$  is local time;  $s$  is the distance along the magnetic field line, positive in the direction north to south. The values of  $C_0$ ,  $t_0$ , and  $s_0$  are selected so that agreement between the measured and the modeled integral airglow intensities at 630 nm is obtained.

## 3 Storm of 15–17 February 1967

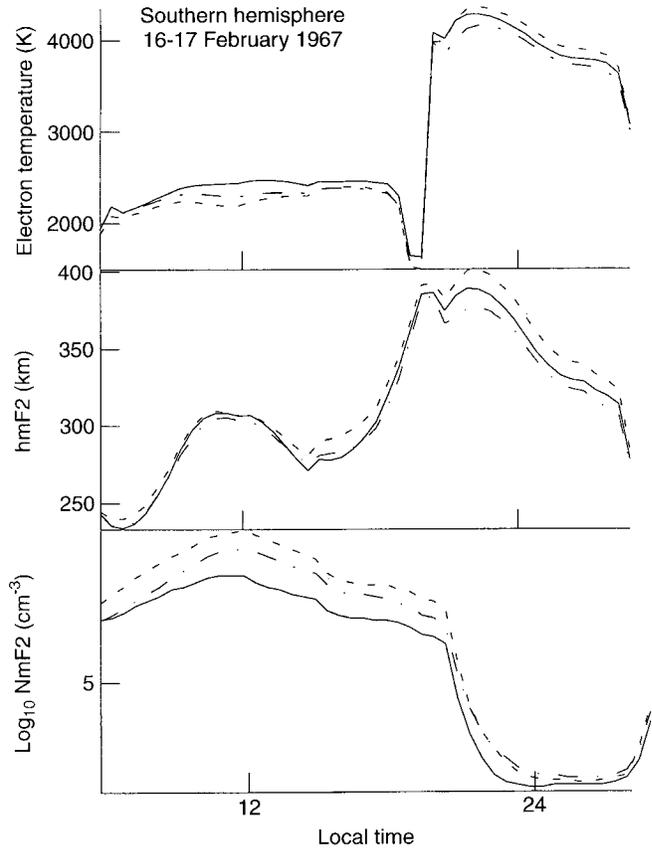
The 15–17 February 1967 geomagnetic storm occurred during medium solar activity ( $F10.7 = 126$  and  $125$ ). The period of magnetic disturbance started with a sudden commencement at 23:48 UT on 15 February. The recovery phase of this geomagnetic storm began near 14:00 UT 16 February with the index  $Dst = -130$   $\gamma$ . The observations of the intensities (as an excess above a background nightglow level),  $I_{630}$ , of the 630-nm oxygen emission in the SAR-arc regions were made by the OV1-10 photometer on 16 February 1967 at 16:26 UT (21:47 LT) in the northern hemisphere at geographic latitude  $62.95^\circ$  and geographic longitude  $80^\circ$  (L-shell value of the geomagnetic field is 2.72,  $I_{630} = 645$  R) and at 16:01 UT (22:15 LT) in the southern hemisphere at geographic latitude  $-42.37^\circ$  and geographic longitude  $93^\circ$  ( $L = 2.83$ ,  $I_{630} = 510$  R) (LaValle and Elliott, 1972). Although the L-shells were not precisely coincident, the difference  $\Delta L = 0.11$  is very small and is the order of accuracy of the location determination (LaValle and Elliott, 1972). Therefore, our calculations were carried out along the geomagnetic field line at  $L = 2.72$  from 10:00 LT on 15 February to 04:00 LT on 17 February (this is local time for the northern hemisphere).



**Fig. 1.** The time variations of the modeled integral airglow intensities at 630 nm in the SAR-arc regions in the northern (*bottom panel*) and southern (*top panel*) hemispheres at  $L = 2.72$  for the recovery phase of the 16–17 February 1967 magnetic storm. *Crosses* are the integral intensities at 630 nm in the SAR-arc regions measured by the OV1–10 photometer. The *solid* and *dashed-dotted* lines represent the integral intensities from the IZMIRAN model when the MSIS-86 atomic oxygen densities were decreased by a factor of 1.35 during the storm period in the southern hemisphere, and *dashed* lines are the intensities with the original MSIS-86 atomic oxygen densities. The modeled results were obtained using the Boltzmann vibrational distribution of  $N_2(j)$  (*dashed-dotted lines*) and the non-Boltzmann populations of the first five vibrational levels by solving the  $N_2(j = 1 - 5)$  time-dependent continuity equations (*solid* and *dotted lines*). The *solid line 1* and the *solid* and *dotted lines* show results when the diffusion of  $O(^1D)$  was included in the IZMIRAN model, and the *solid line 2* shows those without the effects of the diffusion of  $O(^1D)$  on  $[O(^1D)]$ . The model A of the rate coefficients of  $O^+(^4S)$  ions with the vibrationally excited nitrogen molecules and quenching of  $O^+(^2D)$  by atomic oxygen was used in the IZMIRAN model

It seems that red arc emission is usually detected during the recovery phase of geomagnetic storms (Rees and Roble, 1975) and we believe that  $t_0$  in Eq. (1) is close to the starting time of the recovery phase. Previous work concerning the rate of energy transfer from the energetic ring-current ions to plasmaspheric electrons has assumed that the region  $-s_0 < s < s_0$  was in plasmasphere above 5000 km (Pavlov, 1996). For the purposes of this paper this region is taken to be above 5000 km too and the values of  $C_0 = 1480 \text{ K}^{3/2} \text{ s}^{-1}$  and  $t_0 = 13:39 \text{ UT}$  (19:00 LT in the northern hemisphere at geographic latitude  $62.95^\circ$  and geographic longitude  $80^\circ$ ) were selected for calculations of the additional heating rate given by Eq. (1).

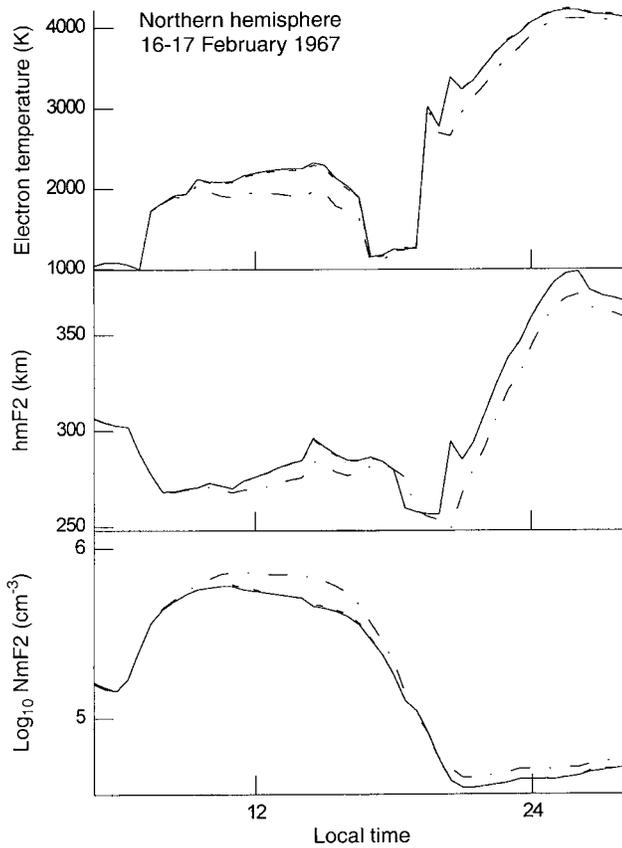
Figures 1–3 show the time variations of the modeled integral airglow intensities at 630 nm and the modeled



**Fig. 2.** The time variations of the modeled NmF2 (*bottom panel*), hmF2 (*middle panel*), and the electron temperature (*top panel*) at the F2 peak altitude in the SAR-arc region at  $L = 2.72$  in the southern hemisphere for the 16–17 February 1967 magnetic storm. The *solid* and *dashed-dotted* lines are hmF2 and NmF2 from the IZMIRAN model when the MSIS-86 atomic oxygen densities were decreased by a factor of 1.35 during the storm period in the southern hemisphere, and *dashed* lines are hmF2 and NmF2 with the original MSIS-86 atomic oxygen densities. The modeled results were obtained using the Boltzmann vibrational distribution of  $N_2(j)$  (*dashed-dotted lines*) and the non-Boltzmann populations of the first five vibrational levels by solving the  $N_2(j = 1 - 5)$  time-dependent continuity and energy equations (A19, A20) (*solid* and *dotted lines*) and the model A of the rate coefficients of  $O^+(^4S)$  ions with the vibrationally excited nitrogen molecules and quenching of  $O^+(^2D)$  by atomic oxygen

F2 peak parameters NmF2, hmF2, and  $T_{em}$  in the SAR arc regions at  $L = 2.72$  in the northern and southern hemispheres for the 16–17 February 1967 magnetic storm which will be explained later. Crosses shown in Fig. 1 represent the integral intensities at 630 nm in the SAR-arc regions measured by the OV1-10 photometer at 21:47 LT (northern hemisphere) and 22:15 LT (southern hemisphere) on 16 February 1967.

The excitation of atomic oxygen by thermal electrons provides the main contribution to the values of the integral intensity at 630 nm in the SAR-arc region. The rate,  $P_t$ , of  $O(^1D)$  formation in collision of atomic oxygen with thermal electrons determined by Eq. (A18) of the Appendix depends on  $T_e$ . Due to this dependence, the production of  $O(^1D)$  increases with increasing  $T_e$ . However, as given by Eq. (A18), the production rate decreases linearly with decreasing electron density. A

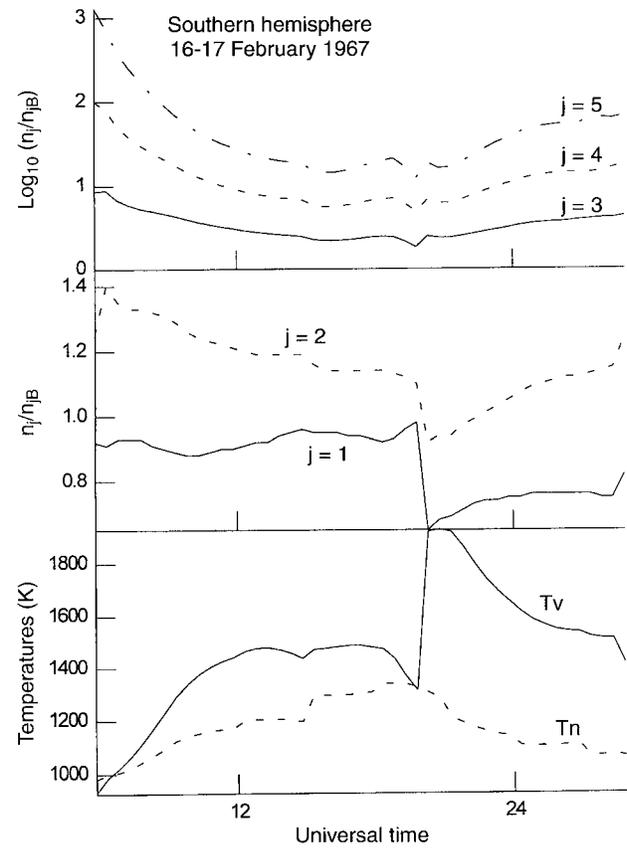


**Fig. 3.** The time variations of the modeled F2 peak parameters in the SAR-arc region at  $L = 2.72$  the northern hemisphere for 15–17 February 1967 magnetic storm. *Top:*  $T_e$ . *Middle:* hmF2. *Bottom:* NmF2. The curves are the same as in Fig. 1 (the dashed lines are very close to the solid lines and are not seen in the Fig. 2)

decrease in  $N_e$ , causes an increase in  $T_e$  (the electron-ion cooling rate of the thermal electrons is proportional to  $N_e^2$ ) and an increase in  $T_e$  causes a decrease in  $N_e$  (as a result of the reactions of vibrationally excited nitrogen molecules  $N_2(j = 1-5)$  with  $O^+(^4S)$  ions shown in Table 1 in the Appendix). Therefore, the resulting dependence of  $P_t$  (and, hence, the volume emission rate,  $V_{630}$ , and  $I_{630}$ ) on  $N_e$  and  $T_e$  appears to be very complex. The main causes of the variations of the integral intensity at 630 nm in the SAR-arc region are the variations in  $N_e$  and  $T_e$  in the region close to and above the F2 region, which are determined by the variations in the modeled F2 peak altitude, hmF2, and density, NmF2, and the modeled electron temperature,  $T_{em}$ , at the F2 peak altitude.

#### 4 The role of non-Boltzmann distribution of vibrationally excited molecular nitrogen

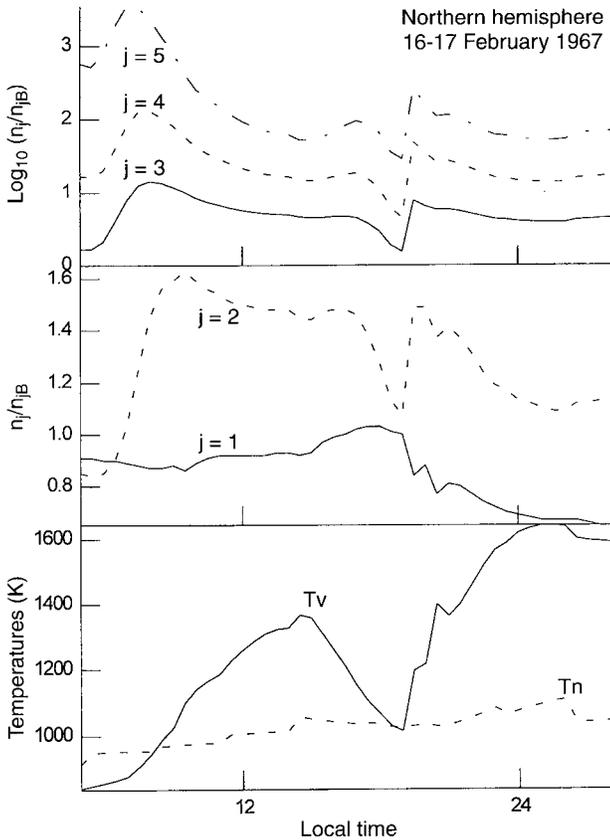
The vibrational distribution of  $N_2$  at levels does not follow a Boltzmann distribution in the F region of the ionosphere (Newton *et al.*, 1974; Torr and Torr, 1982; Pavlov, 1988, 1989; Pavlov and Namgaladze, 1988; Pavlov and Buonsanto, 1997). These deviations can affect the recombination rate of  $O^+(^4S)$  ions and the



**Fig. 4.** The time variations of the vibrational and neutral temperatures and populations of the first five vibrational levels of  $N_2(j)$  in comparison with the Boltzmann distribution during the 15–17 February period at the F2 peak altitude within the SAR-arc regions at  $L = 2.72$  in the northern hemisphere. The solid lines show the modeled  $T_v$ ,  $n_1/n_{1B}$ , and  $n_3/n_{3B}$ , and the dashed lines show the modeled  $T_n$ ,  $n_2/n_{2B}$  and  $n_4/n_{4B}$ , and the dashed-dotted line is the modeled  $n_5/n_{5B}$ . The model A of the rate coefficients of  $O^+(^4S)$  ions with the vibrationally excited nitrogen molecules and quenching of  $O^+(^2D)$  by atomic oxygen was used in the IZMIRAN model

heating rate of electrons which influence the electron temperature and density and the integral airglow intensities at 630 nm in the SAR-arc regions.

The calculated fractional deviation of the  $N_2$  vibrational populations from a Boltzmann one,  $n_j/n_{jB}$ , the vibrational temperature,  $T_v$  of  $N_2(j)$ , and the neutral temperature,  $T_n$ , at the F2 peak altitude are presented in Fig. 4 (southern hemisphere) and Fig. 5 (northern hemisphere). The present study suggests that the calculated distribution is highly non-Boltzmann at vibrational levels  $j > 2$ . The deviations from the Boltzmann distribution are not significant at the first and second vibrational levels of  $N_2$ . Decreases in  $n_j/n_{jB}$  and an increase in  $T_v$  are seen at the recovery phase, after the beginning of the increase in  $T_e$  due to the additional heating rate of electrons. Unlike the northern hemisphere, the nighttime minimum values of  $n_j/n_{jB}$  ( $j = 1-2$ ) and the nighttime maximum value of  $T_v$  are realized near  $t_0$  at 20.19–20.49 LT on 16 February in the southern hemisphere. In the northern hemisphere the deviations  $n_j/n_{jB}$  ( $j = 2-5$ ) and the vibrational temper-



**Fig. 5.** The time variations of the vibrational and neutral temperatures and populations of the first five vibrational levels of  $N_2$  in comparison with the Boltzmann distribution during the 15–17 February period at the F2 peak altitude within the SAR-arc regions at  $L = 2.72$  in the southern hemisphere. The curves are the same as in Fig. 4

ature have their minimum values at 19:00 LT on 16 February. From the diurnal variations of the calculated vibrational (solid line) and neutral (dashed line) temperatures shown in Figs. 4 and 5, it follows that  $T_v < T_n$  in the atmosphere during the night without the additional heating rate of electrons where the production frequency of  $N_2(j)$  is low [details are given by Pavlov (1994)]. This means that for these periods the populations of  $N_2(j)$  are less than the populations for a Boltzmann distribution with temperature  $T_n$ . In the SAR-arc regions  $T_v$  is larger than  $T_n$  due to the enhanced thermal excitation of  $N_2$  as a result of high thermal electron temperatures at F2-region altitudes, but this difference between  $T_v$  and  $T_n$  is not more than 712 K in the either hemisphere.

The modeled results shown in Figs. 1–3 were obtained from the IZMIRAN model using the Boltzmann vibrational distribution of  $N_2(j)$  (dashed-dotted lines) by solving the time-dependent energy equation, Eq. (A20), for the vibrational quanta of  $N_2(j)$  and the non-Boltzmann populations of the first five vibrational levels of  $N_2(j)$  (solid and dotted lines) by solving the  $N_2(j = 1-5)$  time-dependent continuity and energy equations, Eqs. (A19–A20).

Figure 1 shows that the use of the non-Boltzmann vibrational distribution of  $N_2$  and the choice of  $C_0 = 1480 \text{ K}^{3/2} \text{ s}^{-1}$  in Eq. (1) brings the measured and modeled integral intensity at 630 nm into agreement only in the northern hemisphere. It seems likely that the intensity discrepancy between the measured and modeled  $I_{630}$  of the southern hemisphere (the disagreement between the dotted lines and the cross at the top panel of Fig. 1) is caused by inaccuracies in key inputs such as the composition and temperature of the neutral atmosphere (the MSIS-86 model does not accurately predict the neutral atmosphere during a magnetic storm). We found that the integral intensity at 630 nm of the SAR arc region is much more sensitive to the variations in the oxygen number densities than the number densities of  $N_2$ ,  $O_2$  and the neutral temperature. Solid lines in Fig. 1 show the model results when the MSIS-86 neutral atomic oxygen density of the southern hemisphere was decreased by the factor of 1.35 at all altitudes during the geomagnetic storm period. No adjustments were made to the  $[O]$  values before a sudden storm commencement at 23:48 UT on 15 February. This MSIS-86 correction is enough to reproduce very well the observed  $I_{630}$  on 16 February.

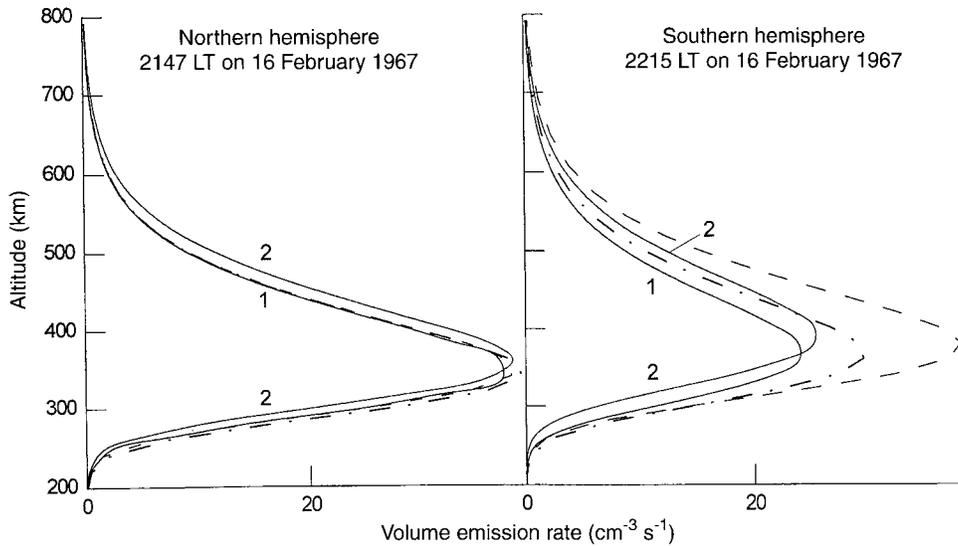
It can be seen from the comparison between the solid line (marked 1) and the dashed-dotted lines of Fig. 1 and those of Figs. 2 and 3, that in the SAR-arc regions the non-Boltzmann distribution assumption leads to the decrease in the calculated NmF2 up to a factor of 1.2–1.4 and brings about reductions by factors of 1.1–1.2 in the integral intensity at 630 nm in comparison with NmF2 and the integral intensity calculated using the Boltzmann vibrational distribution of  $N_2$ .

From our calculations it follows that the use of the Boltzmann vibrational distribution of  $N_2$  gives good agreement between measured and modeled  $I_{630}$  of the northern hemisphere if we choose the value of  $C_0 = 1450 \text{ K}^{3/2} \text{ s}^{-1}$  in Eq. (1). However, in order to reproduce the observed  $I_{630}$  of the southern hemisphere it is necessary to decrease by a factor of 1.55 the MSIS-86 neutral atomic oxygen density in the southern hemisphere at all altitudes during the geomagnetic storm period. Therefore, the use of the non-Boltzmann distribution assumption improves the agreement between the data and the theoretical results giving lower the ratio of the southern-hemisphere intensity to the northern-hemisphere intensity in comparison with the Boltzmann distribution assumption.

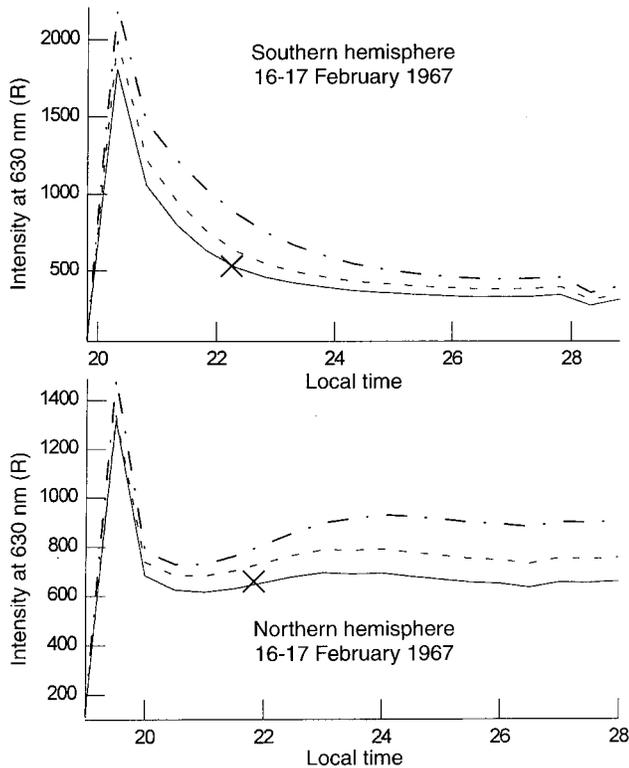
### 5 Effects of the $O(^1D)$ diffusion and the uncertainties in the rate coefficients of $O^+(^4S)$ and $O^+(^2D)$ ions on the integral intensity at 630 nm

We use one of the important changes of the IZMIRAN model given in the Appendix to take into account the effect of the  $O(^1D)$  diffusion on the volume emission and the integral intensity at 630 nm. To model these effects on  $[O(^1D)]$ , the diffusion term in the  $O(^1D)$  continuity Eq. (A24) of the IZMIRAN model was added.

Figure 6 shows calculated altitude profiles of volume emission rate at 630 nm for the 16 February 1971



**Fig. 6.** Computed altitude profiles of volume emission rate at 630 nm for 16 February 1971 magnetic storm within the SAR arc regions at  $L = 2.72$  in the northern (left panel) and southern (right panel) hemispheres at 2147 and 2215 LT. The curves are the same as in Fig. 1



**Fig. 7.** The time variations of the modeled integral intensities at 630 nm in the SAR-arc regions in the northern (bottom panel) and southern (top panel) hemispheres at  $L = 2.72$  for the recovery phase of the 16–17 February 1967 magnetic storm. Crosses represent the integral intensities at 630 nm in the SAR-arc regions measured by the OVI-10 photometer at 21:47 LT (northern hemisphere) and 22:15 LT (southern hemisphere) on 16 February 1967. The solid, dashed and dashed-dotted lines are the calculated integral intensities from model A, B, and C of the rate coefficients of  $O^+(^4S)$  ions with the vibrationally excited nitrogen molecules and quenching of  $O^+(^2D)$  by atomic oxygen described in the Appendix. The modeled results were obtained using the non-Boltzmann populations of the first five vibrational levels by solving the  $N_2(j = 1-5)$  time-dependent continuity equations A19, A20 when the diffusion of  $O(^1D)$  was included in the IZMIRAN model and the MSIS-86 atomic oxygen densities were decreased by the factor 1.35 during the storm period in the southern hemisphere

magnetic storm within the SAR-arc regions at  $L = 2.72$  in the northern (left panel) and southern (right panel) hemispheres at 21:47 LT and 22:15 LT on 16 February 1967 when the measurements of  $I_{630}$  were made by the OVI-10 photometer. In Figs. 1 and 6 the solid lines marked by 1 show the results when we use in the IZMIRAN model the  $O(^1D)$  continuity Eq. (A24) with the diffusion term, whereas the solid lines marked 2 were obtained without the diffusion of  $O(^1D)$ . It can be seen from the comparison of lines 1 and 2 that the effect of including the  $O(^1D)$  diffusion results in a decrease of 4–6% in the calculated integral intensity of the northern hemisphere and of 7–13% in the calculated integral intensity of the southern hemisphere. These changes in the integral intensities are caused by volume emission rate changes. Figure 6 shows that the effect of using the  $O(^1D)$  diffusion in the IZMIRAN model results in a 20–26-km decrease in the modeled peak altitude of the volume emission rate.

The updated IZMIRAN model includes the option to use models A, B, or C of the rate coefficients of  $O^+(^4S)$  ions with the vibrationally excited nitrogen molecules and quenching of  $O^+(^3D)$  by atomic oxygen described in the Appendix. The results shown in Figs. 1–6 were obtained using model A. Figure 7 shows the time variations of the modeled integral airglow intensities at 630 nm from models A (solid lines), B (dashed lines), and C (dashed-dotted lines) in the SAR-arc regions at  $L = 2.72$  in the northern (bottom panel) and southern (top panel) hemispheres for the 16–17 February 1967 magnetic storm. Crosses shown in Fig. 7 represent the integral intensities at 630 nm in the SAR-arc regions measured by the OVI-10 photometer. The used values of  $C_0$ ,  $t_0$ , and  $s_0$  in Eq. (1) were the same as for results shown in Figs. 1–6 and the MSIS-86 neutral atomic oxygen density of the southern hemisphere was decreased by the factor of 1.35 in models A, B, and C at all altitudes during the geomagnetic storm period.

The modeled  $N_e$ ,  $T_e$ , and  $I_{630}$  are sensitive to the assumed values of the rate coefficients of  $O^+(^4S)$  ions with the vibrationally excited nitrogen molecules and

quenching of  $O^+(^2D)$  by atomic oxygen. The modeled  $N_e$  are larger in model B than in model A, and in model C than in model B in the SAR-arc regions (at altitudes close and above the F2 region). As a result, the modeled  $I_{630}$  increases in the SAR-arc regions by 1–15% (northern hemisphere) and 10–21% (southern hemisphere) in model B, and by a factor 1.1–1.4 (northern hemisphere) and 1.2–1.65 (southern hemisphere) in model C, in comparison with model A.

## 6 Conclusions

This paper presented a comparison between measured and calculated integral airglow intensities at 630 nm in the SAR-arc regions of the northern and southern hemispheres, corresponding to the time-period of the 15–17 February 1967 magnetic storm. The model used is an enhanced and updated version of the IZMIRAN model that has been steadily developed over the years. This time-dependent mathematical model of the Earth's ionosphere and plasmasphere includes interhemispheric coupling, the  $H^+$ ,  $O^+(^4S)$ ,  $He^+$ ,  $O^+(^2D)$ ,  $O^+(^2P)$ ,  $NO^+$ ,  $O_2^+$ , and  $N_2^+$  ions, electrons, photoelectrons, the electron and ion temperatures, vibrationally excited nitrogen molecules, the components of thermospheric wind,  $O(^1D)$  atoms and the MSIS-86 neutral atmosphere model. The major enhancements to the IZMIRAN model added for this study are the inclusion of the  $He^+$  ions (three major ions:  $O^+$ ,  $H^+$ , and  $He^+$  and three ion temperature), the updated photochemistry and energy balance equations for ions and electrons, the diffusion of  $NO^+$  and  $O_2^+$  ions and  $O(^1D)$  and the revised electron cooling rates arising from their collisions with unexcited  $N_2$ ,  $O_2$  molecules and  $N_2$  molecules at the first vibrational level. The updated model includes the option to use the models of the Boltzmann or non-Boltzmann distributions of vibrationally excited molecular nitrogen. In order to model the electron temperature, the additional heating rate of the electron gas providing the heat flux to the SAR-arc region was added in the energy balance equation at the altitude above 5000 km during the recovery phase of the geomagnetic storm.

The deviations from the Boltzmann distribution for the first five vibrational levels of  $N_2$  were calculated. The present study suggests that the calculated distribution is highly non-Boltzmann at the vibrational levels  $j > 2$  and the deviations from the Boltzmann distribution are not significant at the first and second vibrational levels of  $N_2$ . It was found that the non-Boltzmann distribution assumption leads to the decrease in the calculated NmF2 up to a factor of 1.2–1.4 and brings about factors of 1.1–1.2 reductions in the integral intensity at 630 nm in comparison with NmF2 and the integral intensity calculated using the Boltzmann vibrational distribution of  $N_2$ . The MSIS-86 oxygen number densities of the southern hemisphere were decreased by a factor of 1.35 at all altitudes for the storm period 15–17 February in order to decrease the modeled integral intensity to match observations better. The preceding results indicate that the use of the non-Boltzmann distribution

assumption improves the agreement between the data and the theoretical results decreasing the ratio of the calculated southern-hemisphere intensity to the northern-hemisphere intensity.

The enhanced IZMIRAN model used for this study takes into account the effect of the  $O(^1D)$  diffusion on the volume emission and the integral intensity at 630 nm. We found that in the SAR-arc region the inclusion  $O(^1D)$  diffusion results in a 20–26-km decrease in the modeled peak altitude of volume emission rate, in a decrease of 4–6% in the calculated integral intensity of the northern hemisphere and 7–13% in the calculated integral intensity of the southern hemisphere. The modeled intensities of the southern hemisphere are more sensitive to the assumed values of the rate coefficients of  $O^+(^4S)$  ions with the vibrationally excited nitrogen molecules and quenching of  $O^+(^2D)$  by atomic oxygen than the modeled intensities of the northern hemisphere.

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## Appendix IZMIRAN model description

The model used is an enhanced and updated version of one that we have steadily developed over the years (Pavlov, 1991, 1994, 1996; Pavlov and Buonsanto, 1997). In the model, coupled time-dependent equations of continuity, momentum, and energy balance for  $O^+(^4S)$ ,  $H^+$ , and  $He^+$  ions, and electrons are solved along a centered-dipole magnetic field line for the concentrations, field-aligned diffusion velocities, and temperatures of ions and electrons from a base altitude (160 km) in the northern hemisphere through the plasmasphere to the same base altitude in the southern hemisphere. In the northern and southern hemispheres the model solves time-dependent continuity equations for  $O^+(^2D)$ ,  $O^+(^2P)$ ,  $NO^+$ ,  $O_2^+$ ,  $N_2^+$ , vibrationally excited molecular nitrogen,  $N_2(v)$ , at the first five vibrational levels  $v = 1–5$ , and the momentum equations for the horizontal components of thermospheric wind within the altitude range 120–700 km and the time-dependent continuity equation for  $O(^1D)$  in the region between 120 and 1500 km in altitude.

### *Ion continuity equations*

The ion continuity equations for the three major ions,  $O^+(^4S)$ ,  $H^+$ , and  $He^+$  and for the minor ions,  $O^+(^2D)$ ,  $O^+(^2P)$ ,  $NO^+$ ,  $O_2^+$ , and  $N_2^+$  are given by

$$\frac{\partial}{\partial t} N_j + B \frac{\partial}{\partial s} C_i N_j B^{-1} = Q_j + \Delta Q_j + P_j - L_j, \quad (A1)$$

where  $N_j$  is the ion concentration,  $B$  is the geomagnetic field,  $t$  is the local time,  $s$  is the distance along the magnetic field line, positive in the direction north to south,  $P_j$  and  $L_j$  are the production rates and the ion loss rates by the chemical reactions of Table 1,  $Q_j$  and  $\Delta Q_j$  are the

**Table 1.** Chemistry of ions (<sup>a</sup>The effective temperature  $T = (m_i T_n + m_n T_i)/(m_i + m_n) + m_i m_n V_d^2/[3k(m_i + m_n)]^{-1}$  where  $V_d$  is plasma drift velocity)

No.	Reaction	Rate coefficient (cm <sup>3</sup> s <sup>-1</sup> ) or rate (s <sup>-1</sup> )	Reference
1.	$O^+(^4S) + N_2(0) \rightarrow NO^+ + N$	$1.533 \cdot 10^{-12} - 5.92 \cdot 10^{-13} (T/300)$ $+ 8.60 \cdot 10^{-14} (T/300)^2$ for $300 \leq T \leq 1700$ K, $2.73 \cdot 10^{-12} - 1.155 \cdot 10^{-12} (T/300)$ $+ 1.483 \cdot 10^{-13} (T/300)^2$ for $1700 < T < 6000$ K	Albritton <i>et al.</i> (1977) <sup>a</sup> Viehland and Mason (1977) St.-Maurice and Torr (1978)
2.	$O^+(^4S) + N_2(1) \rightarrow NO^+ + N$	see text	
3.	$O^+(^4S) + N_2(2) \rightarrow NO^+ + N$	see text	
4.	$O^+(^4S) + N_2(3) \rightarrow NO^+ + N$	see text	
5.	$O^+(^4S) + N_2(4) \rightarrow NO^+ + N$	see text	
6.	$O^+(^4S) + N_2(5) \rightarrow NO^+ + N$	see text	
7.	$O^+(^4S) + O_2 \rightarrow O_2^+ + N$	$[2.82 \cdot 10^{-11} - 7.74 \cdot 10^{-12} (T/300)$ $+ 1.073 \cdot 10^{-12} (T/300)^2$ $- 5.17 \cdot 10^{-14} (T/300)^3$ $+ 9.65 \cdot 10^{-16} (T/300)^4]$ $\exp(-8.53 \cdot 10^{-4} T_n + 0.315)$	Albritton <i>et al.</i> (1977) <sup>a</sup> Chen <i>et al.</i> (1978) Viehland and Mason (1977) St.-Maurice and Torr (1978) Pavlov (1993)
8.	$O^+(^2D) + O \rightarrow O^+(^4S) + O$	see text	
9.	$O^+(^2P) + O \rightarrow O^+(^4S) + O$	$4 \cdot 10^{-10}$	Chang <i>et al.</i> (1993)
10.	$O^+(^2P) + N_2 \rightarrow O + N_2^+$	$3.4 \cdot 10^{-10}$	Chang <i>et al.</i> (1993)
11.	$O^+(^2P) \rightarrow O^+(^4S) + h\nu$	0.047	Seaton and Osterbrock (1957)
12.	$O^+(^2P) \rightarrow O^+(^4D) + h\nu$	0.172	Seaton and Osterbrock (1957)
13.	$O^+(^2D) + N_2 \rightarrow N_2^+ + O$	$8 \cdot 10^{-10}$	Johnsen and Biondi (1980)
14.	$O^+(^2D) + O_2 \rightarrow O_2^+ + O$	$7 \cdot 10^{-10}$	Johnsen and Biondi (1980)
15.	$O^+(^2P) + e \rightarrow O^+(^4S) + e$	$4.7 \cdot 10^{-8} (300/T_e)^{0.5}$	Torr and Torr (1982)
16.	$O^+(^2P) + e \rightarrow O^+(^4S) + e$	$1.5 \cdot 10^{-7} (300/T_e)^{0.5}$	Torr and Torr (1982)
17.	$O^+(^2D) + e \rightarrow O^+(^4S) + e$	$7.8 \cdot 10^{-8} (300/T_e)^{0.5}$	Torr and Torr (1982)
18.	$N_2^+ + e \rightarrow N + N$	$2.6 \cdot 10^{-7} (300/T_e)^{0.37}$	Canosa <i>et al.</i> (1991) Guberman (1991)
19.	$N_2^+ + O_2 \rightarrow N_2 + O_2^+$	$5.0 \cdot 10^{-11} (300/T)^{0.8}$	Lindinger <i>et al.</i> (1974) <sup>a</sup>
20.	$N_2^+ + O \rightarrow NO^+ + N$	$1.4 \cdot 10^{-10} (300/T_i)^{0.44}$	Mc Farland <i>et al.</i> (1974)
21.	$N_2^+ + O \rightarrow O^+(^4S) + N_2$	$9.8 \cdot 10^{-12} (300/T_i)^{0.23}$	Mc Farland <i>et al.</i> (1974)
22.	$N_2^+ + NO \rightarrow NO^+ + N_2$	$3.3 \cdot 10^{-10}$	Fehsenfeld <i>et al.</i> (1970)
23.	$NO^+ + e \rightarrow N + O$	$4.3 \cdot 10^{-7} (300/T_e)$	Torr and Torr (1979) Bates (1988)
24.	$O_2^+ + e \rightarrow O + O$	$1.6 \cdot 10^{-7} (300/T_e)^{0.55}$ for $T_e \geq 1200$ K, $2.0 \cdot 10^{-7} (300/T_e)^{0.7}$ for $T_e < 1200$ K	Mehr and Biondi (1969) Walls and Dunn (1974) Bates (1988)
25.	$O_2^+ + NO \rightarrow NO^+ + O_2$	$4.4 \cdot 10^{-10}$	Lindinger <i>et al.</i> (1974)
26.	$O_2^+ + N \rightarrow NO^+ + O$	$1.2 \cdot 10^{-10}$	Fehsenfeld (1977)
27.	$O^+(^4S) + H \rightarrow H^+ + O$	$4.3 \cdot 10^{-11} (T_n + T_i/16$ $+ 1.2 \cdot 10^{-8} V_d^2)^{0.5}$	Banks and Kockart (1973) Barakat and Schunk (1984)
28.	$O^+ + H^+ \rightarrow O^+(^4S) + H$	$3.44 \cdot 10^{-10} (T_n + T_i/16$ $+ 1.2 \cdot 10^{-8} V_d^2)^{0.5}$ $/[5 + 3 \cdot \exp(-227.8/T_n$ $+ \exp(-325.6/T_n))]$	Banks and Kockart (1973) Barakat and Schunk (1984)
29.	$He^+ + N_2 \rightarrow He + N + N^+$	$1.0 \cdot 10^{-9}$	Adams and Smith (1976)
30.	$He^+ + N_2 \rightarrow He + N_2^+$	$6.5 \cdot 10^{-10}$	Adams and Smith (1976)
31.	$He^+ + O_2 \rightarrow He + O + O^+(^4S)$	$7.0 \cdot 10^{-10}$	Mauclaire <i>et al.</i> (1979)
32.	$He^+ + O_2 \rightarrow He + O + O^+(^2D)$	$4.0 \cdot 10^{-10}$	Mauclaire <i>et al.</i> (1979)

production rates of ions by photoionization and due to photoelectrons,  $C_j = V_j + W$ ,  $V_j$  is the field-aligned diffusion velocity,  $W$  is the field-aligned wind-induced plasma drift velocity. In our model we neglect the  $C_j$  effects on the  $O^+(^2D)$ ,  $O^+(^2P)$ , and  $N_2^+$  ion concentrations.

In the present model, the solar EUV fluxes,  $\Phi_\infty(\lambda)$ , are obtained from the EUVAC model (Richards *et al.*, 1994) or EUV94X model (Tobiska, 1993, 1994). The reactions of photoionization and the ionization due to photoelectrons of the neutral species,  $N_2$ ,  $O_2$ ,  $O$ , and  $He$

are included in the model. We use Eqs. (A3) and (A5) in Pavlov (1994) to calculate  $Q_j$  and  $\Delta Q_j$ . The revised photoionization and photoabsorption cross sections are taken from Richards *et al.* (1994); these authors did not publish the photoionization cross section for  $He$ , and for these we have used values from Torr and Torr (1982). The Chapman function approximation is taken from Smith and Smith (1972). The model results are sensitive in the key rate coefficients  $K_2$ – $K_6$ , and  $K_8$  of the reactions 2–6 and 8 presented in Table 1.

The recombination rate coefficients of  $O^+(^4S)$  ions with  $N_2(v > 0)$  were measured by Schmeltekopf *et al.* (1968) and Ferguson *et al.* (1984) for  $T_n = T_i = 300$  K. Van Zandt and O'Malley (1973) determined the dependence of these rate coefficients on the neutral temperature by using the results of Schmeltekopf *et al.* (1968) and a simple theoretical model of these reactions. Pavlov (1988) enlarged the use of this approach by taking into account the dependence of these rate coefficients on the effective temperature, and changed the rate coefficient of  $O^+(^4S)$  ions with  $N_2(0)$  given by Van Zandt and O'Malley (1973) to the rate coefficient  $K_1$  of Table 1.

The rate coefficient for the quenching of  $O^+(^2D)$  by atomic oxygen has not been measured and was assumed to be  $10^{-10} \text{ cm}^3 \text{ s}^{-1}$  (Oppenheimer *et al.*, 1977; Fox and Dalgarno, 1985) and  $5 \cdot 10^{-12} \text{ cm}^3 \text{ s}^{-1}$  (Abdou *et al.*, 1984). These rate coefficients were obtained from agreement between the measured and model densities of  $N_2^+$  and the different points of view were discussed by Torr and Torr (1982), Abdou *et al.* (1984), Ferguson (1984), and Fox and Dalgarno (1985).

Our model of the thermal plasma in the ionosphere and plasmasphere includes the option to use three sets for models of these rate coefficients. In model A the rate coefficient  $K_5 = 10^{-10} \text{ cm}^3 \text{ s}^{-1}$  and the rate coefficients  $K_2$ – $K_6$  are determined by the equations given by Pavlov (1988) and Pavlov and Namgaladze (1988). This approach was used in the IZMIRAN model (Pavlov, 1994, 1996; Pavlov and Buonsanto, 1997). Model B uses  $K_8 = 5.10^{-12} \text{ cm}^3 \text{ s}^{-1}$  and the recombination rate coefficients of  $O^+(^4S)$  ions with  $N_2(v > 0)$  measured by Schmeltekopf *et al.* (1968) for  $T_n = T_i = 300$  K:  $K_2 = 1.3 \cdot 10^{-12} \text{ cm}^3 \text{ s}^{-1}$ ,  $K_3 = 5.0 \cdot 10^{-11} \text{ cm}^3 \text{ s}^{-1}$ ,  $K_4 = 1.2 \cdot 10^{-10} \text{ cm}^3 \text{ s}^{-1}$ ,  $K_5 = 2.6 \cdot 10^{-10} \text{ cm}^3 \text{ s}^{-1}$ ,  $K_6 = 3.6 \cdot 10^{-10} \text{ cm}^3 \text{ s}^{-1}$ . In model C the rate coefficient  $K_8 = 5.0 \cdot 10^{-12} \text{ cm}^3 \text{ s}^{-1}$  and the ratios of the recombination rate coefficients of  $O^+(^4S)$  ions with  $N_2(v > 0)$  to the rate coefficients of  $O^+(^4S)$  ions with  $N_2(0)$  are the same as measured by Schmeltekopf *et al.* (1968) for  $T_n = T_i = 300$  K:  $K_2(T)/K_1(T) = 1$ ,  $K_3(T)/K_1(T) = 38$ ,  $K_4(T)/K_1(T) = 85$ ,  $K_5(T)/K_1(T) = 220$ ,  $K_6(T)/K_1(T) = 270$ . The model C assumes that the temperature dependence of these reactions rates are the same as for unexcited  $N_2$ , i.e., are as given by St. Maurice and Torr (1978). This approach was used in the FILP model (Torr *et al.*, 1990)

### Transport formulation

The momentum equations for determination of the ion field-aligned velocities were formulated by Conrad and Schunk (1979). In updated “low-speed” model the nonlinear acceleration terms and stress tensor divergences are neglected and the ion diffusion velocities are calculated by solving the system equations

$$\sum_j v_{ij}(1 - \Delta_{ij})(V_j - V_i) - V_i \sum_n v_{in}(1 - \Delta_{in}) \quad (\text{A2})$$

$$= \left[ \frac{\partial}{\partial s} p_i + N_i N_e^{-1} \frac{\partial}{\partial s} P_e + m_i N_i g + p_i \gamma_i \frac{\partial}{\partial s} \ln T_i \right] (m_i N_i)^{-1},$$

where  $N_e$  is the electron concentration,  $p_e = kN_e T_e$ ,  $P_i = kN_i T_i$ ,  $k$  is Boltzmann's coefficient,  $T_e$  is the electron temperature,  $T_i$  is the  $i$ -th ion temperature,  $m_i$  denotes the mass of the  $i$ th ion,  $v_{ij}$  and  $v_{in}$  are the collision frequencies for momentum transfer between ions and between ion and neutral, the thermal diffusion coefficients  $\gamma_i$  and the correction factors  $\Delta_{ij}$  and  $\Delta_{in}$  are given by Conrad and Schunk (1979),  $g = \sin I (g_0 L^{-2} \sin^4 \theta - 1.5 \Omega^2 L R_E \sin^4 \theta)$ ,  $L = r R_E^{-1} \sin^2 \theta$  is the McIlwain parameter,  $\Omega = 7.3 \cdot 10^{-5} \text{ s}^{-1}$ ,  $g_0 = 980 \text{ cm s}^{-2}$ ,  $r$  denotes radial distance from the Earth's center,  $\theta$  is the geomagnetic colatitude,  $I$  is the magnetic field dip angle,  $\sin I = 2 \cos \theta (1 + 3 \cos^2 \theta)^{-0.5}$ .

The wind-induced plasma drift velocity

$$W_d = U_\theta \cos I, \quad (\text{A3})$$

where  $U_\theta$  is the component of thermospheric wind in spherical polar coordinates  $(r, \theta, \phi)$ ,  $\phi$  is the geomagnetic longitude.

For determination of the value of  $U_\theta$  it is necessary to solve the system of momentum Eqs. (A13, A14) given by Pavlov (1994). The large uncertainties in the calculated  $U_\theta$  lead to large errors in the modeled electron densities and peak height altitude, hmF2, of the F2-layer (Richards, 1991; Pavlov, 1994). This problem was largely overcome when Richards (1991) developed a technique for deriving equivalent neutral winds from the observed hmF2. Ideally, the observed hmF2 should be available for several hours before the time of the comparison, but very often only a limited set of hmF2 measurements is available or we have not got the observed hmF2 and thus cannot use the method developed by Richards (1991) or our modification of this method (Pavlov and Buonsanto, 1997) to calculate an equivalent plasma drift velocity. To overcome this problem, we can try to use hmF2<sub>IRI</sub> given by the IRI-90 model (Bilitza, 1990) to calculate an equivalent plasma drift velocity,  $W_{\text{IRI}}$ , for some hypothetical undisturbed conditions with the same solar and time conditions as during the studied conditions:

$$W_{\text{IRI}}(t + \Delta t) = [\text{hmF2}_{\text{IRI}}(t + \Delta t) - \text{hmF2}(t)] \beta(t)^{-1} + W_{\text{IRI}}(t), \quad (\text{A4})$$

$$\beta(t) = 1.76 \cdot 10^{-5} [T_n(T_i + T_e)^{-1} + 0.46] (T_n v_{in})^{2/3} L(O^+)^{-1/3}, \quad (\text{A5})$$

where hmF2( $t$ ) and  $W_{\text{IRI}}(t)$  are the calculated heights of the F2 layer and the equivalent drift velocity of the plasma at time  $t$ ,  $L(O^+)$  is the recombination rate of the  $O^+(^4S)$  ions, the values of  $T_n$ ,  $T_e$ ,  $T_i$ ,  $v_{in}$ , and  $L(O^+)$  are given at 300-km altitude (the dependence of  $\beta$  on altitude is small at the F-region altitudes because of  $\beta(t) \sim \{[O]^2 L^{-1}\}^{1/3}$ ).

After that we calculate the wind-induced plasma drift velocities

$$W_q = U_\theta \text{ (quiet)} \cos I, \quad W_m = U_\theta \text{ (modeled)} \cos I, \quad (\text{A6})$$

where  $U_\theta$  (quiet) and  $U_\theta$  (modeled) are the horizontal components of the thermospheric wind for hypothetical quiet and real studied conditions obtained by solving the two momentum Eqs. (A13, A14) of Pavlov (1994) for these conditions.

Finally, in the northern and southern hemispheres the plasma drift velocities in the ion continuity equations are given by Pavlov (1996) as

$$W_d = W_m - W_q + W_{IRI}. \quad (\text{A7})$$

Such an approach allows us to avoid some serious errors in the calculations of electron density during magnetic storms (Pavlov and Buonsanto, 1997). During quiet conditions the modeled hmF2 are very close to hmF2<sub>IRI</sub>.

If we have the observed peak height altitude, hmF2<sub>exp</sub>, of the F2-layer in the northern or southern hemispheres then the algorithm developed by Richards (1991) with our modification (Pavlov and Buonsanto, 1997) is used to calculate  $W_d$  for this hemisphere

$$W_d(t + \Delta t) = [\text{hmF2}_{\text{exp}}(t + \Delta t) - \text{hmF}(t)]\beta(t)^{-1} + W_d(t), \quad (\text{A8})$$

where hmf( $t$ ) and  $W_d(t)$  are the calculated heights of the F2 layer and the equivalent drift velocity of the plasma at time  $t$ ,  $\beta(t)$  is determined by Eq. (A5).

#### Energy balance equations for ions and electrons

The energy balance equations for ions and electrons

$$\begin{aligned} 1.5 \frac{\partial}{\partial t} T_i &= B(kN_i)^{-1} \frac{\partial}{\partial s} \lambda_i B^{-1} \frac{\partial}{\partial s} T_i - T_i B \frac{\partial}{\partial s} C_i B^{-1} \\ &- 1.5 C_i \frac{\partial}{\partial s} T_i + v_{ie} (T_e - T_i) \\ &+ 3m_i \left\{ \sum_j v_{ij} (m_i + m_j)^{-1} (T_j - T_i) \right. \\ &\left. + (T_n - T_i) \sum_n v_{in} (m_i + m_n)^{-1} \right\} \\ &+ m_i k^{-1} V_d^2 \sum_n v_{in} m_n (m_i + m_n)^{-1}, \end{aligned} \quad (\text{A9})$$

$$\begin{aligned} 1.5kN_e \frac{\partial}{\partial t} T_e &= B \frac{\partial}{\partial s} \lambda_e B^{-1} \frac{\partial}{\partial s} T_e \\ &+ P_e + P_{rc} - \sum_l L_{el}, \end{aligned} \quad (\text{A10})$$

where  $V_d^2 = W^2 + (cE_\perp B^{-1})^2 \{1 + \omega_i^{-2} (\sum v_{in})^2\}^{-1}$ ,  $\omega_i$  is the ion cyclotron frequency,  $E_\perp$  is the perpendicular component of the electric field with respect to the geomagnetic field,  $c$  is the speed of light,  $\lambda_i$  and  $\lambda_e$  are the thermal conductivities of ions and electrons,  $L_{el}$  is the electron cooling rate in the process “ $l$ ”,  $P_e$  is the heating rate of the electron gas by photoelectrons; to calculate  $P_e(s)$  along a centered-dipole magnetic field line the numerical method of Krinberg and Tachilin (1984) for the determination of the photoelectron fluxes within a

plasmaspheric field tube was used on the same field line grid that is used in solving for the densities and temperatures;  $P_{rc}$  is the additional heating rate of the electron gas due to Coulomb collisions between ring-current ions and plasmaspheric electrons and wave-electron interactions.

To calculate the electron thermal conductivity coefficient we use the new mixture rule given by Pavlov (1996). The expressions used for  $\lambda_i$ ,  $v_{ei}$ ,  $v_{ij}$ , and  $L_{el}$  are given by Bailey and Selleg (1990). For the electron cooling rate,  $L_{el}(N_2^*)$ , arising from their collisions with  $N_2$  molecules, with the excitation of vibrational levels  $v = 1-5$  of  $N_2(0)$  and  $v = 2-5$  of  $N_2(1)$ , and the  $N_2^*$  de-excitation effect, the revised expression was used

$$\begin{aligned} L_{el}(N_2^*) &= E_1 \sum_{v=1}^5 v I_{0v} \{ [N_2(0)] - [N_2(v)] \exp(vE_1 T_e^{-1}) \} \\ &+ E_1 \sum_{v=2}^5 (v-1) I_{1v} \{ [N_2(1)] - [N_2(v)] \exp((v-1)E_1 T_e^{-1}) \}, \end{aligned} \quad (\text{A11})$$

where

$$I_{0v} = N_e \{8k T_e (\pi m_e)^{-1}\}^{0.5} \int_0^\infty \sigma_{0v}(x) x \exp(-x) dx, \quad (\text{A12})$$

$$I_{1v} = N_e \{8k T_e (\pi m_e)^{-1}\}^{0.5} \int_0^\infty \sigma_{1v}(x) x \exp(-x) dx, \quad (\text{A13})$$

and where  $x = E(kT_e)^{-1}$ ,  $E$  is electron energy  $E_1 = 3353$  K is the energy of the first level of  $N_2$  given by Radzig and Smirnov (1980),  $\sigma_{0v}(E)$  are the partial cross sections for excitation of  $N_2(0)$  by electrons,  $\sigma_{1v}(E)$  are the partial cross sections for excitation of  $N_2(1)$  by electrons.

The results presented in this paper were obtained using  $\sigma_{01}(E)$  compiled by Itikawa *et al.* (1986) for the region of electron energy from 0.3 to 1.8 eV. We used  $\sigma_{0v}(E)$  recommended by Schulz (1976) for  $E > 1.8$  eV with the normalization factor of 0.7 given by Haddad (1984) and the absolute value of  $\sigma_{01}(E) = 4.45 \cdot 10^{-16}$  cm<sup>2</sup> at  $E = 2$  eV. The partial cross sections for excitation of  $N_2(1)$  to  $N_2(v > 1)$  by electrons calculated by Dube and Herzenberg (1979) for  $v = 2-4$  were adjusted to give the best fit of the calculated  $\sigma_{01}(E)$  also from Dube and Herzenberg (1979), to the measured  $\sigma_{01}(E)$ . The values of  $\sigma_{1v}(E)$  calculated by Chen (1964) for  $v = 5$  were modified by taking into account the new values of  $\sigma_{0v}(E)$ .

For the electron cooling rate,  $L_{el}(O_2^*)$ , arising from collisions of electrons with unexcited  $O_2$  molecules with the excitation of vibrational levels  $v = 1-4$  of  $O_2(0)$  and the  $O_2^*$  de-excitation effect revised expression was used

$$L_{el}(O_2^*) = E'_1 N_e \sum_{v=1}^4 v J_v \{ [O_2(0)] - [O_2(v)] \exp(vE'_1 T_e^{-1}) \}, \quad (\text{A14})$$

**Table 2.** Coefficients for calculations of the electron cooling rate due to  $O_2^*$ 

$j$	1	2	3	4	5	6	7	8	9	10	11	12	13
$E_j^*$	3829	5222	6603	7961	9295	10610	11895	13171	14413	15620	16815	17987	19136
$S_{j1}$	95.7	428	726	796	567	371	202	119	72.1	0	0	0	0
$S_{j2}$	0	0	112	398	595	594	452	316	167	75.0	33.6	0	0
$S_{j3}$	0	0	0	0	36.3	175	261	277	251	155	90.8	54.0	0
$S_{j4}$	0	0	0	0	0	0	0	52.6	110	125	114	71.9	76.5

where  $J_v = \{8kT_e(\pi m_e)^{-1}\}^{0.5} \int_0^\infty \Omega_{0v}(x) x \exp(-x) dx$ ,  $x = E(kT_e)^{-1}$ ,  $E$  is an energy of electrons,  $E_1^* = 2239$  K is the energy of the first level of  $O_2$  given by Radzig and Smirnov (1980),  $\Omega_{0v}(E)$  are partial cross sections for excitation of  $O_2(0)$  by electron impact.

The vibrational-excitation cross sections of  $O_2(0)$ , as a function of impact energy, exhibit very sharp and narrow spikes (Schulz, 1976). The use of the measured integral vibrational-excitation cross sections in these narrow spikes given by Schulz (1976) leads to the new electron cooling rate due to  $O_2^*$ :

$$L_{el}(O_2^*) = 1.61 \cdot 10^{-4} N_e T_e^{-1.5} \sum_{v=1}^4 \{[O_2(0)] - [O_2(v)] \exp(vE_1^* T_e^{-1})\} \sum_{j=1}^{13} S_{jv} \exp(-E_j^* T_e^{-1}), \quad (A15)$$

where the values of  $S_{jv}$  and  $E_j^*$  are given in Table 2, the units of  $[O_2(0)]$ ,  $[O_2(v)]$  and  $N_e$  are  $\text{cm}^{-3}$ , the units of  $T_e$  are K, the units of  $L_{el}(O_2^*)$  are  $\text{K cm}^{-3} \text{s}^{-1}$ .

To calculate the number densities of  $O_2(v)$  we use a Boltzmann distribution with the vibrational temperature  $T_v(O_2^*) = T_n$

$$[O_2(v)] = [O_2(0)] \exp(-vE_1^* T_n^{-1}) \quad (A16)$$

From Eq. (A16) it may similarly be shown that the sum of the vibrationally unexcited and excited molecular oxygen concentrations  $[O_2] = [O_2(0)] \{1 - \exp(-E_1^* T_n^{-1})\}$ . There is no possibility to differentiate the vibrationally unexcited  $O_2$  from the vibrationally excited  $O_2$  during the mass-spectrometer measurements of the neutral species in the atmosphere, and we therefore believe that the MSIS-86 neutral model (Hedin, 1987) used in our model calculates the value of  $[O_2]$ . The difference between  $[O_2]$  and  $[O_2(0)]$  is not negligible during solar maximum. For example  $[O_2]/[O_2(0)] = 1.29$  for  $T_n = 1500$  K.

We use the revised cooling rate in collision of  $O(^3P)$  with thermal electrons

$$L_{el}(O(^1D)) = E_0 P_t, \quad (A17)$$

where the rate of  $O(^1D)$  formation in collision of  $O(^3P)$  with thermal electrons is calculated as (Pavlov, 1990, 1996)

$$P_t = CN_e \{ [O] - 1.8 [O(^1D)] \exp(X) \} X^{0.5} (1 + a_1 X - a_2 X^2) (b_1 + b_2 X + X^2)^{-1} \exp(-X), \quad (A18)$$

$$X = E_0 T_e^{-1}, E_0 = 22861 \text{ K}; C = 2.72 \cdot 10^{-8} \text{ cm}^3 \text{ s}^{-1}, a_1 = 7.8035 \cdot 10^{-2}, a_2 = 3.4534 \cdot 10^{-3}, b_1 = 1.6882, b_2 =$$

3.331, the units of  $[O]$ ,  $[O(^1D)]$ , and  $N_e$  are  $\text{cm}^{-3}$ , the units of  $T_e$  are K, the units of  $L_{el}(O(^1D))$  are  $\text{K cm}^{-3} \text{ s}^{-1}$ .

The loss rate of the  $O^+(^4S)$  ions and  $L_{el}(N_2^*)$  are functions of number densities,  $n_j$ , of vibrationally excited molecular nitrogen  $N_2(j)$ . The model includes the option to use the models of the Boltzmann and non-Boltzmann distributions of vibrationally excited molecular nitrogen.

#### Continuity and energy equations for the vibrationally excited nitrogen molecules

Pavlov (1988, 1989) treated the vibrationally excited nitrogen molecules as a set of harmonic and anharmonic oscillator energy levels. The Pavlov (1989) study suggests that the effect of anharmonicity on  $n_j$  is small at the energy levels  $j = 1-5$ . Therefore our model uses the system of equations given by Pavlov (1988) and Pavlov and Namgaladze (1988) to determine  $n_j (j = 1-5)$  using the harmonic oscillator energy level approximation, the vibrational-vibrational and vibrational-translational energy exchange of  $N_2(j)$ , the diffusion of  $N_2(j)$  in the mixture of  $N_2(0)$ ,  $O_2$ , and  $O$ , the production rates,  $Q_j$ , of  $N_2(j)$  as follows:

$$\begin{aligned} \frac{\partial}{\partial t} n_j + \frac{\partial}{\partial z} F_j(N_2^*) &= \{(j+1) n_{j+1} - [(j+1)\Theta + j] n_j \\ &+ j \Theta n_{j-1}\} \tau_{vt}^{-1} + \{(j+1)(1+\alpha) n_{j+1} \\ &- [(j+1)\alpha + j(1+\alpha)] n_j + j \alpha n_{j-1}\} \tau_{vv}^{-1} + Q_j, \end{aligned} \quad (A19)$$

$$\begin{aligned} \frac{\partial}{\partial t} \alpha - \frac{1}{n} \frac{\partial}{\partial z} D n \frac{\partial}{\partial z} \alpha &= \frac{1}{n} \sum_j j Q_j - \left( \alpha - \frac{\Theta}{1-\Theta} \right) \tau_{vt}^{-1} \\ &- \alpha K_m [CO_2], \end{aligned} \quad (A20)$$

where  $F_j(N_2^*) = -D \left[ \frac{\partial}{\partial z} n_j + n_j \left( \frac{1}{H} + \frac{1}{T_n} \frac{\partial}{\partial z} T_n \right) \right]$ ,  $n = \sum_j n_j$ ,  $\alpha = \frac{1}{n} \sum_j j n_j$ ,  $\Theta = \exp(-E_1 T_n^{-1})$ ,  $\tau_{vt}^{-1} = \{K_{vt}^{(1)} [O] + K_{vt}^{(2)} [O_2]\} (1 - \Theta)$ ,  $\tau_{vv}^{-1} = K_{vv} n$ ,  $H = kT_n (m g)^{-1}$ ,  $D$  is the diffusion coefficient of the vibrationally excited nitrogen molecules,  $m$  denotes the mass of  $N_2$ ,  $g$  is the acceleration due to gravity,  $E_1 = 3353$  K, the rate coefficients of the vibrational-translational energy exchange  $N_2(j=1)-O$  is measured by McNeal *et al.* (1974) and can be approximated as  $K_{vt}^{(1)} = 1.07 \cdot 10^{-10} \exp(-69.9 T_n^{-1/3}) \text{ cm}^3 \text{ s}^{-1}$  by use of the theory of this energy exchange, the rate coefficients of the vibrational-translational energy

exchange  $N_2(j=1)-O_2$  is given by Gilmore *et al.* (1969) as  $K_{vt}^{(2)} = 3.6 \cdot 10^{-12} T_n^{0.5} \exp(-110 T_n^{-1/3}) \text{ cm}^3 \text{ s}^{-1}$ , for the vibrational-vibrational energy exchange of  $N_2(j=1)$  with  $CO_2$  Taylor (1974) presented  $K_m = 1.7 \cdot 10^{-6} \exp(-175 T_n^{-1/3}) + 6.2 \cdot 10^{-14} \exp(-15.13 T_n^{-1/3}) \text{ cm}^3 \text{ s}^{-1}$ , the rate coefficients of the vibrational-vibrational energy exchange of  $N_2(j=1)$  with  $N_2(j=0)$  is measured by Suchkov and Schebeko (1981) as  $K_{vv} = 5.2 \cdot 10^{-10} T_n^{-3/2} \text{ cm}^3 \text{ s}^{-1}$ .

It is required to know the value of  $n_6$  to calculate the value of  $n_5$  from Eqs. (A19, A20), and the condition  $n_6 = 0$  was assumed. The diffusion coefficients of vibrationally excited molecules are not sensitive to the number of energy level (Fujimoto *et al.*, 1976) and we can use

$$D = Y\{[O](D(N_2, O))^{-1} + [N_2](D(N_2, N_2))^{-1} + [O_2](D(N_2, O_2))^{-1}\}^{-1}, \quad (\text{A21})$$

where  $Y = [O] + [N_2] + [O_2]$ , the binary diffusion coefficients of the unexcited  $N_2$  molecules (Pavlov, 1981; Pavlov and Namgaladze, 1988)  $D(N_2, O) = 9.69 \cdot 10^{16} \cdot T_n^{0.774} Y^{-1}$ ,  $D(N_2, N_2) = 9.76 \cdot 10^{16} \cdot T_n^{0.69} Y^{-1}$ ,  $D(N_2, O_2) = 8.29 \cdot 10^{16} \cdot T_n^{0.724} Y^{-1}$ , the units of  $[O]$ ,  $[O_2]$  and  $[N_2]$  are  $\text{cm}^{-3}$  and the units of  $D$  are  $\text{cm}^{-3} \text{ s}^{-1}$ .

The production rates of  $N_2(j)$  are given as

$$Q_j = I_{0j}\{n_0 - n_j \exp(jE_1 T_e^{-1})\} + I_{1j}\{n_1 - n_j \exp[(j-1)E_1 T_e^{-1}]\} + Q_{j1} + Q_{j2}, \quad (\text{A22})$$

where the first two terms describe the thermal electron excitation of  $N_2(0)$  and  $N_2(1)$  and the deexcitation of  $N_2(j)$ , the values of  $I_{0j}$  and  $I_{1j}$  are determined by Eqs. (A12 and A13),  $Q_{j1}$  and  $Q_{j2}$  are the  $N_2(j)$  production rates in the reactions  $O(^1D)$  with  $N_2(0)$  and  $N(^4S)$  with  $NO$ .

The rate coefficients,  $\gamma_j$ , of the reactions  $O(^1D) + N_2(0) \rightarrow O + N_2(j)$  are not measured. Streit *et al.* (1976) reported that the measured rate coefficient  $\gamma = \sum_j \gamma_j$ . Tully (1974) developed a classic collision complex model to describe production of  $N_2(j)$  in these reactions. This theoretical approach defines  $\gamma = \int_0^{E^*} K(E) dE$ , where  $E^*$  is the critical energy of the potential surface intersection of  $N_2O$  with separated  $O(^3P) + N_2$ ,  $K(E)$  is a partial rate constat for formation of  $N_2(j)$  vibrational energy between  $E$  and  $E + dE$  calculated by Tully (1974) for  $E^* = 1.18$  and  $1.97$  eV.

The predicted value of  $\gamma$  given by Tully (1974) agrees with the experimental results of Streit *et al.* (1976). Although there is considerable uncertainty in the theoretical results, they suggest that the amount of energy channeled into  $N_2$  vibration can be an average of  $30 \pm 10\%$  of the 1.97-eV exothermicity of quenching. This is in excellent agreement with the measurements of Slanger and Black (1974) who concluded that an average of  $33 \pm 10\%$  of the 1.97-eV electronic energy of  $O(^1D)$  converts into  $N_2$  vibrational energy.

Using  $K(E)$  we calculate the branching ratios as  $\frac{\gamma_j}{\gamma} = \frac{\int_{E_{j-1}}^{E_j} K(E) dE}{\int_0^{E^*} K(E) dE}^{-1}$ , where  $E_j$  is the energy of the  $j$ -th vibrational level of  $N_2(j)$ . Our model use the values of  $\frac{\gamma_j}{\gamma} = 0.293, 0.253, 0.212, 0.159, 0.083$  for  $j = 1-5$  and  $\frac{\gamma_j}{\gamma} = 0$  for  $j > 5$ , the measured rate coeffi-

cient  $\gamma$  given by Streit *et al.* (1976) and the production rate of  $N_2(j)$  in the reaction  $O(^1D)$  with  $N_2(0)$  as  $Q_{j1} = \gamma_j [O(^1D)] n_0$ . For these values of  $\frac{\gamma_j}{\gamma}$  an average of 36% of the 1.97-eV electronic energy of  $O(^1D)$  converts into  $N_2$  vibrational energy, in excellent agreement with the measurements of Slanger and Black (1974).

The vibrationally excited nitrogen molecules may be also produced in the reactions  $N(^4S)$  with  $NO$  with the rate  $Q_{j2} = \varepsilon_j \beta [N(^4S)][NO]$ , where  $\beta = 3.4 \cdot 10^{-11} \text{ cm}^3 \text{ s}^{-1}$  (Lee *et al.*, 1978),  $\varepsilon_j$  is the probability vibrational distribution given by Phillips and Schiff (1962).

Solving the system Eqs. (A19), (A20) the model can calculate the number densities of vibrationally excited molecular nitrogen determining the real non-Boltzmann distribution of  $N_2(j)$ . The model includes the option to use the model of the Boltzmann distributions of vibrationally excited molecular nitrogen

$$n_j = n_{jB} = n_{0B} [\alpha(1 + \alpha)^{-1}]^j = n_{0B} \exp(-jE_1 T_v^{-1}), \quad (\text{A23})$$

where  $T_v$  is the vibrational temperature and the value of  $\alpha$  is calculated from Eq. (A20). In this case from Eq. (A23) it may be similarly be shown that the sum of the vibrationally unexcited and excited molecular nitrogen concentrations  $[N_2] = n = n_{0B}(1 + \alpha)$ . There is no possibility to differentiate the vibrationally unexcited  $N_2$  from the vibrationally excited  $N_2$  during the mass-spectrometer measurements of the neutral species in the atmosphere, and therefore we believe that the MSIS-86 model calculates the value of  $[N_2]$ . The difference between  $[N_2]$  and  $n_{0B}$  is not negligible during solar maximum conditions. For example  $[N_2]/n_{0B} = 1.12$  for  $T_v = 1500$  K.

#### Continuity equation for $O(^1D)$

We calculate  $[O(^1D)]$  from the continuity equation using the chemical reactions of formation and loss of  $O(^1D)$  presented in Table 3, the production of  $O(^1D)$  in the photodissociation reaction of oxygen molecules, the  $O(^1D)$  formation in collision of  $O(^3P)$  with thermal electrons and photoelectrons, and the diffusion of  $O(^1D)$  in the mixture of  $N_2(0)$ ,  $O_2$ , and  $O$ :

$$\begin{aligned} \frac{\partial}{\partial t} [O(^1D)] + \frac{\partial}{\partial z} F(O(^1D)) &= Q_d + P_{ph} + P_t + K_1 [O_2^+] N_e \\ &- \{K_2 [N_2] + K_3 [O_2] + K_4 [O] \\ &+ K_5 + K_6 + K_7\} [O(^1D)], \end{aligned} \quad (\text{A24})$$

where  $F(O(^1D)) = -D(O(^1D)) \left[ \frac{\partial}{\partial z} [O(^1D)] + [O(^1D)] \left( \frac{1}{H} + \frac{1}{T_n} \frac{\partial}{\partial z} T_n \right) \right]$ ,  $Q_d = [O_2] \sum_{\lambda} \sigma_d(\lambda) \Phi_{\alpha}(\lambda) \exp[-\tau(\lambda)]$ ,  $H = kT_n (\text{mg})^{-1}$ ,  $D$  is the diffusion coefficient of  $O(^1D)$ ,  $m$  denotes the mass of  $O$ ,  $\tau(\lambda)$  is the optical depth at wavelength  $\lambda$ ,  $\sigma_d(\lambda)$  is the generally accepted photodissociation cross section of  $O_2$  with the production of  $O(^1D)$  (Solomon *et al.*, 1988),  $P_t$  is determined by Eq. (A.18),  $P_{ph}$  is the rate of the  $O(^1D)$  formation in

**Table 3.** The chemical processes of formation and loss of O(<sup>1</sup>D)

No.	Reaction	Rate coefficient (cm <sup>3</sup> s <sup>-1</sup> ) or rate (s <sup>-1</sup> )	Reference
1.	O <sub>2</sub> <sup>+</sup> + e → O( <sup>1</sup> D) + O	1.76 · 10 <sup>-7</sup> (300/T <sub>e</sub> ) <sup>0.55</sup> for T <sub>e</sub> ≥ 1200 K, 2.2 · 10 <sup>-7</sup> (300/T <sub>e</sub> ) <sup>0.7</sup> for T <sub>e</sub> < 1200 K	Mehr and Biondi (1969) Walls and Dunn (1974) Bates (1988)
2.	O( <sup>1</sup> D) + N <sub>2</sub> → N <sub>2</sub> <sup>*</sup> + O	2 · 10 <sup>-11</sup> exp(107.8 T <sub>n</sub> <sup>-1</sup> )	Solomon <i>et al.</i> (1988) Streit <i>et al.</i> (1976)
3.	O( <sup>1</sup> D) + O <sub>2</sub> → O <sub>2</sub> <sup>*</sup> + O	3.2 · 10 <sup>-11</sup> exp(67 T <sub>n</sub> <sup>-1</sup> )	Streit <i>et al.</i> (1976)
4.	O( <sup>1</sup> D) + O → O + O	2.5 · 10 <sup>-12</sup>	Sobral <i>et al.</i> (1993)
5.	O( <sup>1</sup> D) → O( <sup>3</sup> P <sub>2</sub> ) + hν	5.63 · 10 <sup>-3</sup>	Solomon <i>et al.</i> (1988)
6.	O( <sup>1</sup> D) → O( <sup>3</sup> P <sub>1</sub> ) + hν	1.82 · 10 <sup>-3</sup>	Solomon <i>et al.</i> (1988)
7.	O( <sup>1</sup> D) → O( <sup>3</sup> P <sub>0</sub> ) + hν	8.92 · 10 <sup>-7</sup>	Solomon <i>et al.</i> (1988)

collision of O(<sup>3</sup>P) with photoelectrons, the value of  $P_{ph}$  is calculated by using the semi-empirical formula (Pavlov, 1990; Konikov and Pavlov, 1991),  $K_i$  ( $i = 1-6$ ) are the rate coefficients of the chemical reactions of Table 3.

We believe that the diffusion coefficient of O(<sup>1</sup>D) is the same as the diffusion coefficient of O(<sup>3</sup>P) in the mixture of N<sub>2</sub>, O<sub>2</sub>, and O(<sup>3</sup>P):

$$D = Y\{[O](D(O,O))^{-1} + [N_2](D(O,N_2))^{-1} + [O_2](D(O,O_2))^{-1}\}^{-1}, \quad (A25)$$

where  $Y = [O] + [N_2] + [O_2]$ , the binary diffusion coefficients (Pavlov, 1981),  $D(O,N_2) = D(O,O_2) = 9.69 \cdot 10^{16} \cdot T_n^{0.774} Y^{-1}$ , the self-diffusion coefficient of O is the function of the oxygen viscosity coefficient (Ferziger and Kaper, 1972) and can be calculated as  $D(O,O) = 1.94 \cdot 10^{17} \cdot T_n^{0.69} Y^{-1}$ , the units of [O], [O<sub>2</sub>] and [N<sub>2</sub>] are cm<sup>-3</sup> and the units of  $D$  are cm<sup>2</sup> s<sup>-1</sup>.

Important characteristics of airglow that are calculated by the model are the volume emission rate at 630 nm and the integral intensity:

$$V_{630} = K_5[O(<sup>1</sup>D)], I_{630} = \int_0^{\infty} V_{630}(z) dz. \quad (A26)$$

### Model inputs

The inputs for the model are a neutral atmosphere, the solar EUV fluxes, the perpendicular electric fields, and the additional heating rate of the electron gas due to Coulomb collisions between ring-current ions and plasmaspheric electrons and wave-electron interactions. The solar EUV fluxes are obtained from the EUVAC model (Richards *et al.*, 1994) or EUV94X model (Tobiska, 1993, 1994). To specify the neutral temperature and the number densities of N<sub>2</sub>, O<sub>2</sub>, O, and N we utilize the MSIS-86 model (Hedin, 1987) using the 3-h  $A_p$  index as an indicator of the geomagnetic activity. To calculate the density of NO the same scheme of the processes of formation and loss of NO as in Iwagami and Ogawa (1987) is used, so it is not presented here: The ratio of [CO<sub>2</sub>] to ([O<sub>2</sub>] + [N<sub>2</sub>] + [O]) can be taken as  $3 \cdot 10^{-4}$  at 120 km (Banks and Kockarts, 1973). Above 120 km the model uses  $[CO_2](z) = [CO_2](z_0)$

$\exp\{-(z - z_0)H^{-1}\}$ , where  $z_0 = 120$  km,  $H = kT_n(mg)^{-1}$ ,  $m$  denotes the mass of CO<sub>2</sub>.

### Lower and upper boundary conditions

The lower boundary 160 km is set for determination of [O<sup>+</sup>(<sup>4</sup>S)], [H<sup>+</sup>], [He<sup>+</sup>], field-aligned diffusion velocities, and temperatures of ions and electrons from Eqs. (A1, A2, A9, A10) in the northern and southern hemispheres. At this base altitude the diffusion processes of O<sup>+</sup>(<sup>4</sup>S), H<sup>+</sup>, He<sup>+</sup> ions and the processes of transfer of thermal energy of electrons and ions were neglected. At the lower boundary (120 km) the viscosity terms in the momentum equations, the diffusion processes of O<sub>2</sub><sup>+</sup> and NO<sup>+</sup> ions, O(<sup>1</sup>D), vibrationally excited molecules and vibrational quanta are neglected in the northern and southern hemispheres. The upper boundary conditions are assumed as  $V_i = 0$  for O<sub>2</sub><sup>+</sup> and NO<sup>+</sup> ions (see Eq. A2),  $\frac{\partial}{\partial z} \alpha = 0$ ,  $F_j(N_2^*) = 0$ ,  $\frac{\partial}{\partial r} U_\theta = 0$ , and  $\frac{\partial}{\partial r} U_\phi = 0$  at 700-km altitude and  $F(O(<sup>1</sup>D))) = 0$  at 1500-km altitude in the both hemispheres.

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