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Retrieval of vertical constituents and temperature profiles from passive submillimeter wave limb observations of the Martian atmosphere: a feasibility study

Joachim Urban, Karin Dassas, François Forget, and Philippe Ricaud

The investigation of the Martian atmosphere is of key importance for an understanding of the planets present and past. Passive limb observations of thermal radiation at submillimeter wavelengths in the 320–350-GHz range by use of a state-of-the-art satellite receiver on a low Mars orbit allow important parameters such as the mixing ratios of H_2O , HDO, ^{12}CO , ^{13}CO , O_3 , and H_2O_2 as well as the thermal profile to be retrieved with high precision and unprecedented vertical range and resolution, providing valuable information for better understanding of the planet’s water cycle, atmospheric dynamics, and photochemistry. The feasibility of these kinds of measurement is demonstrated by means of model simulations based on realistic atmospheric, spectroscopic, and instrumental parameters. Temperature can be retrieved to ~ 90 km with half-scale height vertical resolution from single-scan measurements of emission lines of the long-lived species ^{12}CO and ^{13}CO . The global water-vapor distribution can be measured even under dry or wet conditions with good vertical resolution from the surface to ~ 45 km, and simultaneous observations of HDO allow useful information on the D/H ratio up to an altitude of ~ 30 km to be derived. The sensitivity of the limb-sounding technique also permits information on the photo-chemically important minor species O_3 and H_2O_2 to be obtained. It is shown that spectral averaging may improve precision, altitude range, and resolution of the retrieved profiles. Other frequency bands are explored, and the 435–465-GHz range is suggested as a possible alternative to the 320–350-GHz range.

1. Introduction

A. Motivation

The motivations for an investigation of the Martian atmosphere are numerous. Playing a central role in the evolution of the planet, its study is crucial for a better understanding of the present and past of

planet Mars. The current Martian climate system is a complex combination of atmospheric dynamics coupled with dust, CO_2 , and water cycles. Understanding how this system works is a key to the history of the planet and its evolution. Because of the similarities between Earth and Mars, a systematic comparative study of the Martian atmosphere’s meteorology, climatology, and aeronomy should furthermore tell us much about the evolution of our own planet.

Observational data are needed to constrain chemical and dynamical models, which are major tools nowadays available for the study of the Martian atmosphere and climate system.^{1–3} Global observations of temperature and wind fields are required for a complete view of the atmospheric circulation. Water vapor and related isotopic species are parameters to be studied for a better understanding of the seasonally asymmetric water cycle, characterized by seasonal and possibly diurnal exchange cycles with the surface. Finally, with photolysis extending to the ground owing to the planet’s low surface pressure, photochemical cycles that control the abundances of major (CO_2 , 95%, N_2 , 3%) and minor

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(CO, 0.08%, H₂O, 0.01%, O₂, 0.1%) constituents play a major role in the Martian atmosphere and need to be investigated by measurements of key species.^{4–7}

Several space missions have so far contributed to the study of key parameters for atmospheric dynamics. Temperature, determined by meteorological and radiative processes, is the major parameter that controls physical (e.g., condensation) and chemical processes. Most-recent measurements of infrared spectrometers such as the Thermal Emission Spectrometer (TES) of the Mars Global Surveyor⁸ (MGS) mission provided global profile information at moderate vertical resolution larger than one scale height.^{9–12} Winds have never been directly measured by spaceborne sensors but are indirectly derived from the thermal gradients.¹² Some sparse observational information on winds has also been obtained from Earth-based radio observations of Doppler shifts.¹³ The water cycle has so far been studied only by measurements of the water-vapor column, e.g., from TES nadir infrared measurements¹⁴ or Earth-based telescopes.^{15,16} Present missions such as Mars Express,¹⁷ which started its observations in January 2003, are designed to extend and improve on TES measurements to better characterize the dynamics and photochemical cycles of the Martian atmosphere. Vertically well-resolved temperature and ozone profiles are measured by UV solar and stellar occultation,¹⁸ with the disadvantage of the limited global and temporal coverage of this technique. Columns of ozone from UV nadir sounding during daytime as well as columns of water vapor from IR spectroscopy will also become available,¹⁹ and in addition some information on the column abundances of HDO and CO should be obtained. In the near future, NASA will launch the Mars Reconnaissance Orbiter,²⁰ carrying the Mars Climate Sounder, which is designed for limb observations of temperature, aerosol, and water-vapor profiles employing filter radiometry at thermal IR wavelengths.

B. Submillimeter-Wave Radiometry

In this study we investigate the potential of spaceborne heterodyne spectroscopy at submillimeter wavelengths for the exploration of the Martian atmosphere. Compared to other spectroscopic techniques such as UV, visible, and thermal IR, microwave temperature and constituent retrievals have certain advantages. Intense emission lines exist for major molecules of the Martian atmosphere, and accurate profiles can be obtained for extreme meteorological conditions and during day and night. Radiative transfer at millimeter and submillimeter wavelengths ($\lambda > 500 \mu\text{m}$) is relatively insensitive to aerosols,^{21,22} given the long wavelengths compared with the typical dimensions reported for dust and cloud particles in the Martian atmosphere ($r_{\text{eff}} < 4 \mu\text{m}$).^{23,24} Thus information can be retrieved under conditions with heavy aerosol loading, which means during periods of particular scientific interest (dust storms, cloud formation). Spectroscopy of rotational lines in the vibrational ground state is relatively simple and, in particular, the local thermal

equilibrium (LTE) assumption remains valid for rotational lines to above 120 km (non-LTE processes are a major problem for temperature retrievals in the IR above ~ 50 km).²⁵ Moreover, reliable space-proven (all solid state) superheterodyne receiver technology is nowadays available up to ~ 700 GHz, and latest digital spectrometer technology allows the narrow lines emitted at the very low pressures of the upper Martian atmosphere to be resolved.^{26–28}

H₂O and HDO are the key target species for the study of the diurnal and seasonal water cycles and the isotopic fractionation. D/H ratios that are five to six times that on Earth have been reported,^{29,30} which might be caused by mechanisms that favor the escape to space of H rather than D. Seasonal and diurnal condensation and fractionation cycles involving approximately half of the available atmospheric water vapor, with HDO being more concentrated in its icy phase because of its lower vapor pressure, may also play a controlling role for the atmospheric water vapor abundance and may lead to considerable seasonal variations.^{16,31,32} CO is a long-lived species whose variations³³ can be used to trace slow meridional transport processes, in addition to direct observations of winds from Doppler shifts. Temperature and pressure profiles, essential parameters for atmospheric dynamics and chemistry, can be retrieved simultaneously from observations of ¹²CO and ¹³CO lines. Moreover, ¹²CO and ¹³CO as well as minor constituents such as O₂, O₃, and H₂O₂ with strong transitions in the submillimeter wavelengths range are potential targets for the study of the photochemical state of the Martian atmosphere and surface. H₂O₂, never detected on Mars, is assumed to be a key reservoir for the chemically active HO_x radicals. With the three-body recombination of CO and O being slow, the equilibrium composition of the Martian CO₂ atmosphere without catalytic cycles would contain much lower mixing ratios of CO and O₂ than observed.³⁴ H₂O₂ probably also acts as a sterilizing oxidant near the surface, which could also diffuse in the subsurface.³⁵ Ozone is controlled by catalytic cycles (HO_x, NO_x, O_x) as in the Earth's middle atmosphere, and observations would allow the predicted variation due to significant water-vapor and temperature variations to be investigated.^{5,36} Measurements of other species such as radicals involved in the photochemical cycles would also be desirable but are more difficult owing to the species' low abundances or their weak spectroscopic line strengths.

In what follows, we first investigate in detail the feasibility of profile measurements of the aforementioned key target molecules H₂O, HDO, ¹²CO, ¹³CO, O₃, and H₂O₂ by simulating a state-of-the-art receiver on a low-altitude Mars orbit performing passive observations of thermal emission lines in the 320–350-GHz spectral range at the atmospheric limb and at nadir. This spectral range provides a compromise with respect to the line strengths of the major target lines (increasing with frequency), the sensitivity of the receiver (decreasing with frequency), and

the field of view for a reasonably sized antenna (decreasing with frequency). The study was conducted while the state-of-the-art microwave sounder Mars Atmosphere Microwave Brightness Observer (MAMBO) was being designed.⁶ The instrument's general design, performance, and scientific objectives (which also include Doppler wind measurements and surface science) will be discussed elsewhere. Here we focus on the retrieval of vertical constituents and temperature profiles.

In Section 2 we describe the MOLIERE-5 model (Microwave Observation Line Estimation and Retrieval, version 5), used in this study for the simulation of spectroscopy, radiative transfer, instrument characteristics, and profile retrieval. Section 3 provides an overview of the instrument concept that we have adopted. The assumptions made for the retrieval simulations are described in detail in Section 4, and the resultant theoretical capabilities of submillimeter wave limb sounding of the Martian atmosphere are summarized in Section 5. Alternative target line selections are discussed in Section 6, and conclusions are drawn in Section 7.

2. Simulating Radiative Transfer and Inversion: the MOLIERE-5 Model

Measurements of a state-of-the-art submillimeter wave limb sounder as well as constituents and temperature retrievals were simulated by use of the general forward-and-inversion model for the millimeter and submillimeter wavelengths range, MOLIERE-5. This model, validated in several intercomparison activities,^{37–39} was explained in detail by Urban *et al.*³⁸

The MOLIERE-5 forward model comprises modules for spectroscopy, radiative transfer, and instrument characteristics. Important properties of the absorption coefficient module with respect to the Martian atmosphere are its capability of line-by-line calculation as well as its inclusion of H₂O and CO₂ continuum models.^{40,41} The radiative transfer module allows for calculations in different geometries such as limb and nadir sounding from orbiting platforms. A spherically stratified (one-dimensional) emitting and absorbing atmosphere is assumed. Weighting functions, required for inversions, are calculated by differentiation of the radiative transfer equation analytically after discretization. Several independent modules allow instrument characteristics such as the antenna's field of view, the sideband response of a heterodyne receiver, and the spectrometer's bandwidth and resolution to be taken into account.

The MOLIERE-5 inversion model consists of a module that calculates linear optimal estimation, a least-squares retrieval method that uses statistical *a priori* knowledge of the retrieved parameters for the regularization of ill-posed inversion problems.⁴³ Both forward and inversion modules are included within a loop allowing nonlinear inversions to be calculated iteratively by use of a Levenberg–Marquardt scheme. The model computes diagnostics such as the mea-

surement and smoothing error covariance matrices along with contribution and averaging kernel functions. Additional features of the inversion code are the possibility of retrieving typically unknown additional parameters such as offsets in the nominal pointing angle of a limb scan or in the assumed temperature of the ground.

3. Instrumental Concept

To simulate a realistic performance of a state-of-the-art submillimeter-wave heterodyne receiver, we adopted the instrumental characteristics of the Mars Atmosphere Microwave Brightness Observer (MAMBO), the concept of a passive microwave sounder that was extensively studied by a consortium of research institutes and companies from France, the United States, Germany, and Sweden.⁶ This receiver, designed to characterize the dynamics and the composition of the Martian atmosphere by measurements at the atmospheric limb and at nadir, was optimized for observations of selected lines in the 320–350-GHz range of the key species H₂O (325.153 GHz), HDO (335.396 GHz), ¹²CO (345.796 GHz), ¹³CO (330.588 GHz), O₃ (326.901 GHz), and H₂O₂ (326.982 GHz). Key elements of the front-end design are a 230-mm-diameter main reflector with a 3-dB beam width of 0.28° at 335 GHz, corresponding to ~7.6 km in terms of tangent altitude for a 350-km orbit, and two planar Schottky-diode mixers that are subharmonically pumped by phase-locked Gunn local oscillators fixed tuned at 329.198 and 341.811 GHz. The downconverted signals of the double-sideband heterodyne receiver are spectrally analyzed by a combination of two state-of-the-art chirp-transform ($\Delta\nu = 100$ kHz; maximum bandwidth, 200 MHz) and eight autocorrelator spectrometers (each with 128 channels; maximum bandwidth, 1024 MHz). It should be noted that the overall receiver concept was constrained to meet the stringent requirements on mass, size, power consumption, and data rate that are typical for such a mission.⁶ The instrumental characteristics derived represent therefore realistic assumptions for the simulations of this study.

4. Setup for the Simulation Study

A. Simulation of Measurements

Measurements were simulated for the individual bands by use of the forward model part of MOLIERE-v5.32.

Spectroscopic line parameters for the line-by-line calculation were taken from the Verdandi database,^{43,44} which merges frequencies, line intensities, and lower-state energies from the Jet Propulsion Laboratory catalog⁴⁵ with pressure-broadening parameters from the HITRAN compilation.⁴⁶ Moreover, broadening parameters for target lines were adapted to Martian conditions by use of measured or calculated values for CO₂ broadening or, if measurements were not available, air-broadening parameters were scaled by a factor of 1.65 as estimated from the ratio of collisional broadening coefficients in air and CO₂

Table 1. Target Lines and Back-End Spectrometer Configuration for the 320–350-GHz Range

Band	Species	ν_0 (GHz)	Chirp Transform Spectrometer		Autocorrelator Spectrometer	
			Bandwidth (MHz)	$\Delta\nu$ (MHz)	Bandwidth (MHz)	$\Delta\nu$ (MHz)
A	^{12}CO	345.796	200	0.1	2×1024	8
	^{13}CO	330.588	80 ^a	0.1	1×512	4
B	H_2O	325.153	120 ^a	0.1	2×1024	8
C	HDO	335.396			1×256	2
D	O_3	326.901			1×64	0.5
	H_2O_2	326.982			1×256	2

^aShared spectrometer.

for molecules and lines already studied.^{47,48} Whereas this approximation was adequate for our simulations, all relevant line-broadening parameters and their temperature dependence should be accurately measured in the laboratory to allow for the analysis of actual atmospheric measurements. CO_2 and H_2O continuum emissions were modeled with appropriate parameterizations.^{40,41}

The atmospheric model was adopted from a general circulation model of the Martian atmosphere³ (see plots of the profiles in Section 5). Because of the expected high variability or large uncertainty of the concentrations of target species, extreme cases (wet and dry atmospheres, upper limit conditions) were investigated additionally.

Refraction was not considered in the radiative transfer calculations because, owing to the low ground pressure (approximately 6–7 hPa), its effect is negligible. Cosmic background radiation was taken into account. When the line of sight hit the ground, the temperature of the surface was used for the background term.

The sensor model assumed a platform altitude of 350 km, as proposed as baseline for the Mars orbiter, and a main antenna with a 230-mm diameter leading to a 3-dB field of view of approximately 7.6 km at the tangent point (Mars radius, ~ 3397 km) for a frequency of 335 GHz. A continuous limb scan from 0 to 130 km with a constant velocity of 4 km/s was modeled, according to the instrument specifications, leading to a widening of the effective antenna field of view. A spectrometer readout in 1-s intervals, corresponding to the acquisition of a limb measurement every 4 km in terms of tangent altitudes, was simulated. For nadir sounding, an integration time of 5 s was assumed. Local-oscillator frequencies of 329.198 GHz (for the bands of H_2O , HDO , O_3 , and H_2O_2) and 338.820 GHz (for ^{12}CO and ^{13}CO) and a perfect sideband ratio of 0.5 were chosen.

Simulations were performed assuming the back-end design proposed for MAMBO. Each target line can be observed with a combination of two kinds of spectrometer: in total, two 200-MHz-wide chirp transform spectrometers with high spectral resolution of 100 kHz to resolve the line centers as well as eight flexible wideband autocorrelator spectrometers, each having 128 channels and covering the line

wings, were used. The back-end design adopted for the individual target bands is summarized in Table 1. Figure 1 shows the results of radiative transfer calculations for standard atmospheric conditions.

B. Retrieval Simulations

To estimate the nominal capabilities of submillimeter-wave limb sounding for the study of critical parameters of the Martian atmosphere we performed nonlinear retrieval simulations, using the MOLIERE-v5.32 inversion model.

Mixing ratio profiles of all species with lines residing in a target band were retrieved simultaneously. H_2O and HDO profiles were retrieved from individual bands, whereas O_3 was retrieved together with H_2O_2 . Temperature was retrieved along with the mixing ratios of ^{12}CO and ^{13}CO . Pressure was recalculated after each iteration from the retrieved temperature profile by use of the hydrostatic equation and a reference pressure of 7 hPa. For each band a profile of CO_2 was retrieved to account for continuum emissions at low tangent altitudes. Mixing ratio profiles, temperatures, and pressures were in general calculated on the grid defined by the tangent altitudes of the measurement in 4-km altitude steps. For comparison, retrieval simulations that employed a reduced altitude grid with 10-km resolution were also performed. To ensure realistic retrieval simulations, additional parameters such as a bias in the pointing angle of the limb scan, a typically critical but uncertain parameter of a limb sounder, and an offset in the temperature of the ground, that accounts for emissions of the Mars surface captured by the antenna at low tangent altitudes and at nadir, were retrieved simultaneously.

Nonlinear retrievals were performed, starting from a first guess (*a priori*), which differed from the (so-called *true*) scenario used to simulate the measurement. To ensure smooth convergence a Marquardt parameter of 0.1 with a change factor of 5 was used, according to the implemented algorithm³⁸ that used the *a priori* variances as a diagonal scaling matrix.

Measurement errors were assumed to be uncorrelated for each spectrometer channel and were estimated from the standard radiometer formula with the nominal integration time and assuming a double-sideband receiver noise temperature of ~ 1600 K.

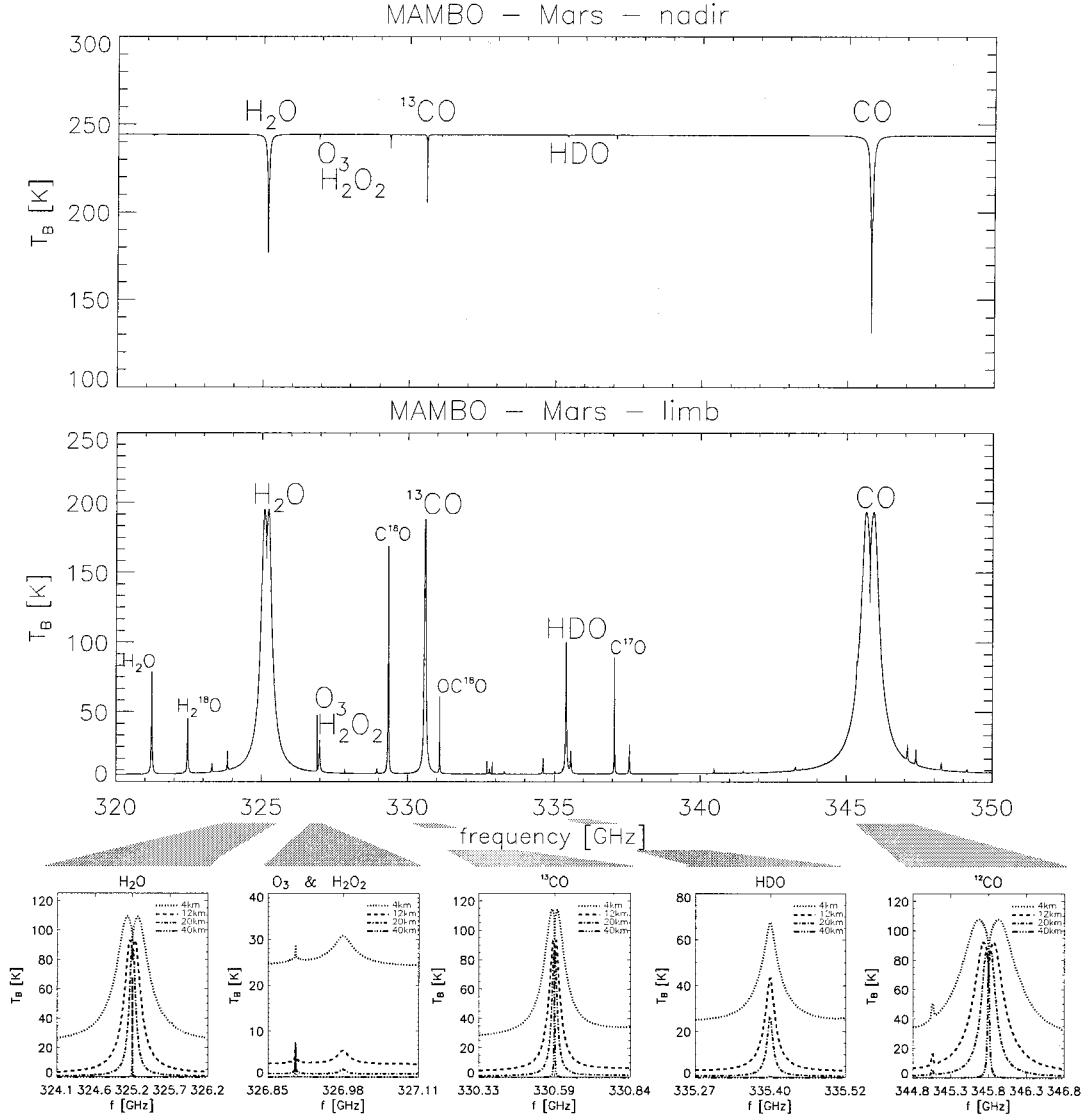


Fig. 1. MOLIÈRE-5 model calculations in the 320–350-GHz range for standard atmospheric conditions. In the nadir observation geometry the transitions appear as absorption lines viewed against the blackbody radiation of the warmer surface (top). By pointing at the limb, an observation system would measure the atmospheric line emission against the weak cosmic background radiation (middle). The tangent altitude of the limb view shown in this example is 6 km. Bottom: model calculations for the individual double-sideband receiver bands of H_2O (325.153 GHz), O_3 and H_2O_2 (326.901 and 326.982 GHz), ^{13}CO (330.588 GHz), HDO (335.396 GHz), and ^{12}CO (345.796 GHz) at selected tangent altitudes (4, 12, 20, and 40 km).

A priori errors in the target species volume mixing ratios (VMR) of 3×10^{-4} for H_2O , 5×10^{-7} for HDO , and 3×10^{-7} for O_3 and H_2O_2 were chosen, which means that uncertainties were constant over altitude and that parameters were assumed to be uncorrelated. A somewhat smaller uncertainty of 20% of the maximum of the corresponding *a priori* profile was used for the less-variable species ^{12}CO ($\sim 1.6 \times 10^{-4}$), ^{13}CO ($\sim 1.2 \times 10^{-6}$) and for continuum CO_2 (~ 0.18). An *a priori* error of 20 K was chosen for temperature, which should largely cover the expected uncertainty and variability about a climatological *a priori* value. Abrupt global temperature increases of 5–15 K, for example, have been observed during intense dust storm events.^{11,12} The *a priori* error of the scan bias, representing the remaining uncertainty after satellite atti-

tude reconstruction, was assumed to be 0.01° (~ 500 m), whereas the ground temperature was assumed to be uncertain to 5 K after being retrieved from nadir continuum observations. Note that the values of the *a priori* errors stated here correspond to the assumed $1 - \sigma$ uncertainties of the retrieval parameters.

5. Results

A. Retrievals of Constituents

Results of the nonlinear retrieval simulations for the target species H_2O , HDO , H_2O_2 , and O_3 are presented in Figs. 2–4. The left-hand sides of the figures show the assumed (true) and the retrieved volume mixing

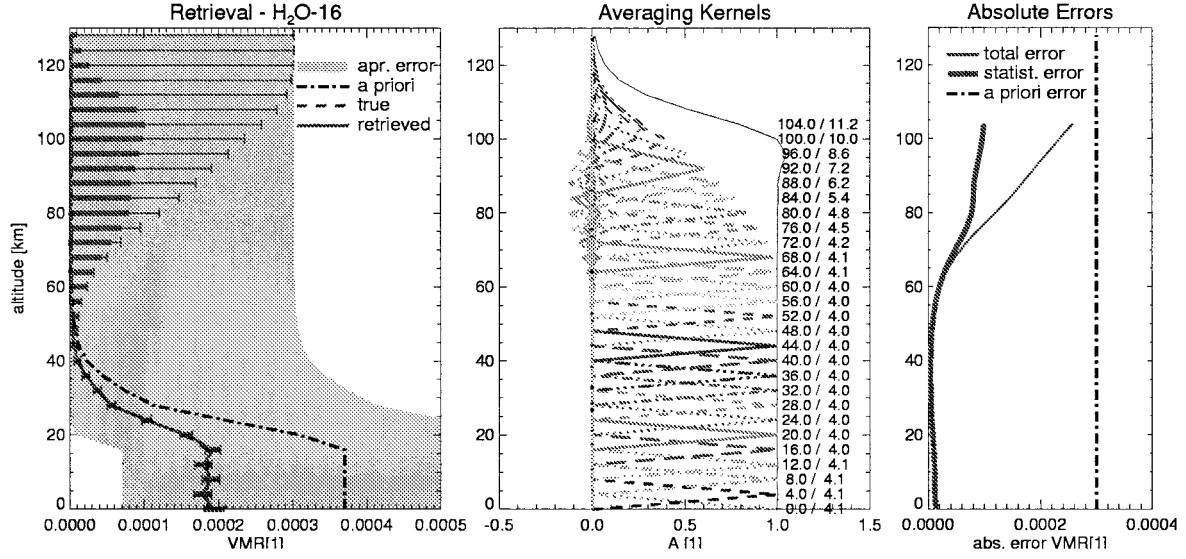


Fig. 2. Results of the nonlinear retrieval simulations for H_2O at 325.153 GHz. Left, retrieved, *a priori*, and true profiles (see legend). Thick error bars indicate the retrieval error that is due to intrinsic receiver noise. The total retrieval error, also including the smoothing error that is due to the limited altitude resolution, is given by the thin error bars. Shaded area: assumed *a priori* error. Middle: averaging kernel functions for each retrieval level. The altitude resolution (FWHM) of the measurement is indicated for each retrieval level. The envelope is the measurement response. Values close to 1 indicate the altitude range where the information from the measurement dominates the *a priori* information. Right: retrieval errors.

ratio profiles. Error bars indicate the statistical errors that are due to intrinsic receiver noise as well as the total error, which also includes contributions from the smoothing error due to the limited altitude resolution of the simulated measurement. Additionally, the *a priori* (or first-guess) profile is provided, along with its assumed $1 - \sigma$ uncertainty. For clarity, the absolute retrieval errors are plotted on the right-hand side of each figure. The middle graph provides the averaging kernel functions for each retrieval altitude level. The full width at half-maximum (FWHM) of these functions are reasonable estimates of the altitude resolution of a given observation sys-

tem and are also shown. Moreover, the sum over all averaging kernel functions at a given altitude level provides the total measurement response: If the value is close to 1, the retrieved parameter is determined by the measurement; values close to 0 indicate that the information comes entirely from the *a priori* parameter.

The standard model assumes a water-vapor profile that corresponds to a total column of 17 precipitable micrometers (pr- μm), as is typical for midlatitude summer conditions.^{3,14} The results of the retrieval simulation for this case are shown in Fig. 2. Information on H_2O from the measurement is theoretically

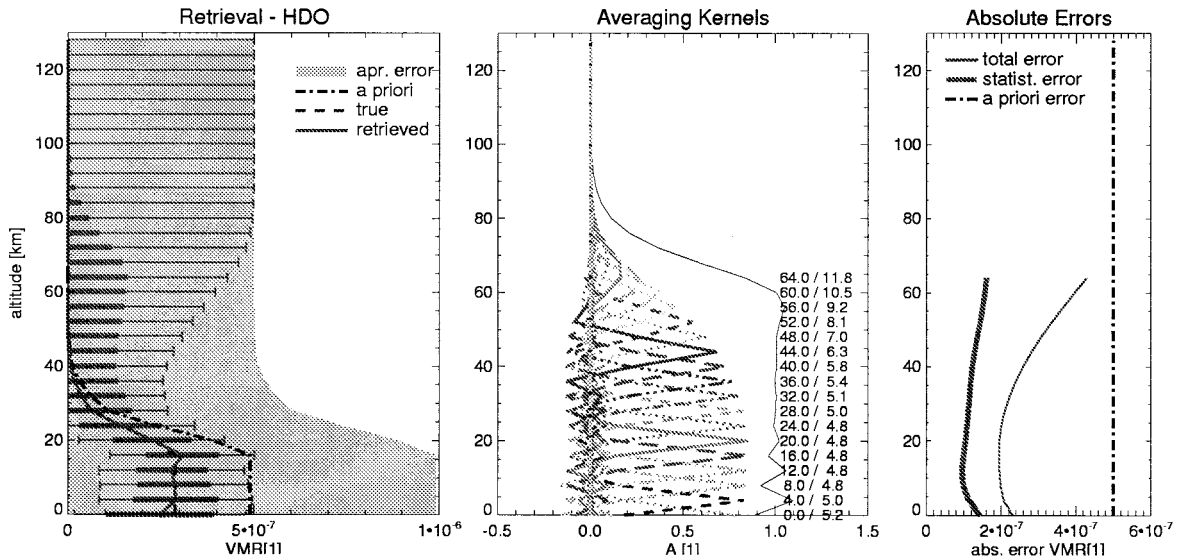


Fig. 3. Results of the nonlinear retrieval simulations for HDO at 335.396 GHz. See the caption of Fig. 2 for an explanation.

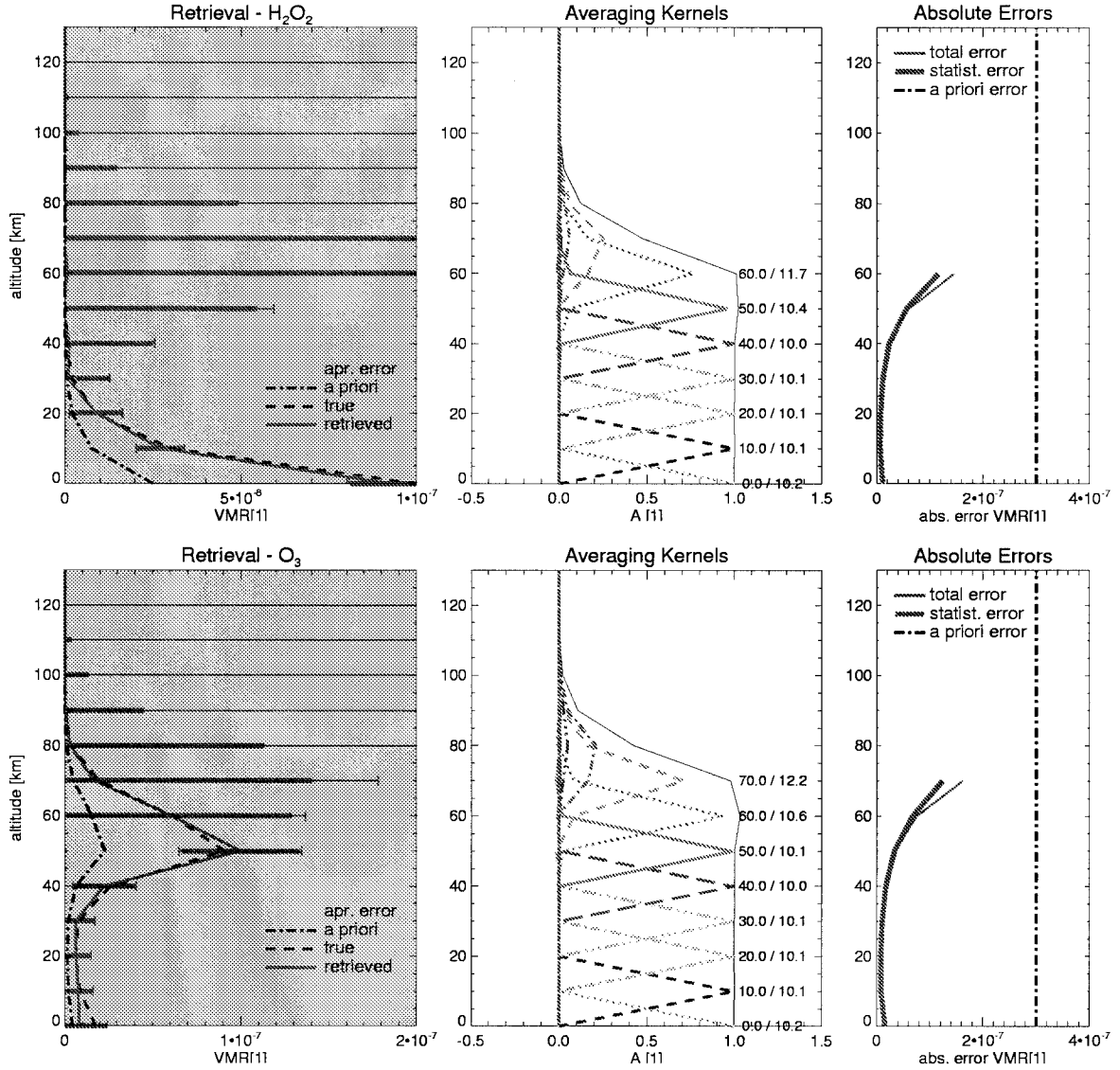


Fig. 4. Results of the nonlinear retrieval simulations for the species H_2O_2 at 326.982 GHz (top) and O_3 at 326.901 GHz (bottom), retrieved simultaneously on an inversion altitude grid of 10-km resolution. See the caption of Fig. 2 for an explanation. The *a priori* error for H_2O_2 and O_3 is 3×10^{-7} .

available up to 100 km with a vertical resolution of the order of the scan readout interval of ~ 4 km below ~ 70 km. The corresponding measurement precision in the mixing ratio depends strongly on altitude: Whereas below 50 km the error that is due to measurement noise is smaller than $\sim 1 \times 10^{-5}$, it increases to values of the order of $5\text{--}10 \times 10^{-5}$ above. Best values of the measurement precision of $\sim 5 \times 10^{-6}$ are obtained at 40 km. The relative error is of the order of 5–10% below 30 km and increases above because of the decreasing H_2O mixing ratio. Compared to the expected mixing ratios of the standard scenario, useful information can practically be retrieved on a single-scan basis to roughly 40–45 km.

For comparison, in nadir geometry one also may retrieve water vapor from 4 to 60 km with a precision of the order of 5×10^{-5} from a single measurement

with an integration time of 5 s but with a considerably poorer altitude resolution ranging from 10 km in the lowest layers to 20 km at the upper limit of the exploitable altitude range. The results of this simulation are shown in Fig. 5. The estimated error that results from statistical receiver noise exceeds the mixing ratios of our theoretical profile above ~ 30 km.

Because water vapor is known to show a high seasonal variability,¹⁴ we also investigated extreme cases (Fig. 5). For extremely wet conditions (column ~ 120 pr- μm), corresponding to an upper limit for the water column at high latitudes in the summer Northern Hemisphere, we obtained results similar to those for standard conditions above approximately 15–20 km, whereas at lower altitudes the signal-to-noise ratio is affected by self-absorption of the strong water-vapor line. The relative error below 30 km is still of the order

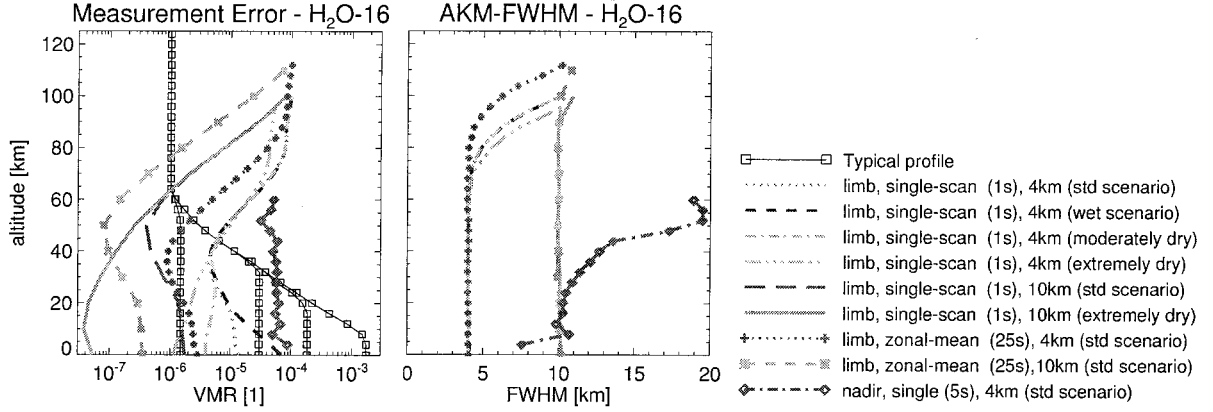


Fig. 5. Comparison of estimated measurement precision (left) and altitude resolution (right) achievable from single-scan (1-s), zonal-mean (25-s), and nadir (5-s) retrievals of H₂O at 325.153 GHz. Also shown are retrieval errors obtained with the assumption of a reduced inversion altitude grid of 10-km resolution. Typical and extreme water vapor profiles (extremely wet, 120 pr- μ m; standard, 17 pr- μ m; moderately dry, 3 pr- μ m; extremely dry, 0.15 pr- μ m), are indicated for comparison.

of 5–10%. Even in a moderately dry atmosphere (column ~ 3 pr- μ m), to be expected at middle latitudes of the winter Southern Hemisphere, water vapor can still be measured up to approximately 40–45 km with a vertical resolution of the order of 4 km and a precision in the range $2\text{--}6 \times 10^{-6}$. These values for measurement precision, which are not affected by self-absorption, represent the noise-equivalent limit for the retrieval of mixing ratios from the spectral measurements of a single limb scan and correspond roughly to a total water column of 0.15–0.2 pr- μ m. A profile representing extremely dry winter pole conditions with a H₂O column of ~ 0.15 pr- μ m is shown in Fig. 5 for comparison.

Averaging of spectra at the price of considerably degraded spatial resolution would allow us to reduce further the statistical error caused by intrinsic receiver noise and thus to obtain profile information at higher altitudes or even under extremely dry atmospheric conditions. Assuming, for example, zonal averages with an effective integration time of 25 s, useful information on H₂O could be obtained with 4-km resolution up to ~ 50 km if the corresponding retrieval error were compared with the mixing ratios of wet, standard, or moderately dry conditions (Fig. 5). This information would still be useful to constrain chemical-dynamical models of the atmosphere.

It should also be noted that for limb observations there is a trade-off between the resolution of the retrieval altitude grid and the value of the statistical retrieval error. By increasing the retrieval grid step, one obtains a smaller retrieval error and vice versa, according to the information from measurements at given tangent altitudes that contribute to the retrieved mixing ratio at a certain level. The previously derived quantitative estimates of the measurement capabilities for submillimeter-wave limb sounding have therefore to be considered in the context of a retrieval grid with 4-km resolution. For comparison, retrieval simulations were also performed with a retrieval grid with 10-km resolution. The single-profile retrieval on

the 10-km grid for standard conditions indicates error measurement that is due to measurement noise smaller than 2×10^{-6} below 60 km with a best value of $\sim 5 \times 10^{-7}$ at 50 km. Compared with the retrieval on the 4-km grid for the same standard atmospheric conditions, we thus obtain an improvement of roughly 1 order of magnitude and an extension of the useful measurement range to ~ 60 km. Zonal mean (25-s) retrievals on the 10-km grid would provide information even up to 70–75 km if we compared them with the expected mixing ratios in this altitude range. Moreover, single-scan retrievals under extremely dry atmospheric conditions do not suffer from self-absorption of the signal at low altitudes and would provide useful results up to ~ 60 km on the 10-km retrieval grid (see also Fig. 5).

The HDO retrieval simulation for standard atmospheric conditions shown in Fig. 3, resulted in a measurement response close to 1 up to ~ 60 km. The altitude resolution is of the order of ~ 5 km below 35 km and degrading above, with increasing altitude. The retrieval precision is approximately $1\text{--}1.5 \times 10^{-7}$ ($\sim 20\text{--}30\%$ up to 20–25 km) for mixing ratios retrieved on the 4-km grid. For extremely wet conditions, similar results were obtained (see Fig. 6).

As for H₂O, the precision of retrieval and thus the useful altitude range can be improved by averaging of measurements, which is useful in particular for the study of isotope fractionation above the condensation level, as discussed, for example, by Fouchet and Lelouch.³¹ Our standard HDO profile was derived from the standard H₂O profile and corresponds below 20 km to a D/H ratio of five times the isotopic ratio in standard mean ocean water (SMOW) of 1.5576×10^{-4} , a depletion of 80% above ~ 45 km, and a linear transition from 20 to 45 km. A profile for extremely wet conditions was derived in a similar way from the wet H₂O profile, but the assumed depletion above 45 km was only 20%. The error in the D/H ratio derived from the single-scan precision of H₂O

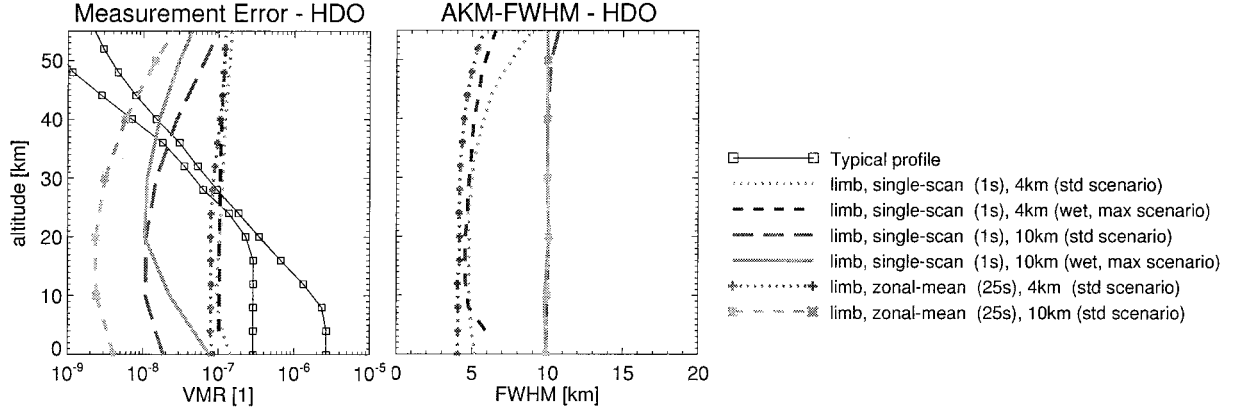


Fig. 6. Comparison of estimated measurement precision (left) and altitude resolution (right) achievable from single-scan (1-s) and zonal-mean (25-s) retrievals of HDO at 335.396 GHz. Also shown are retrieval errors obtained with the assumption of a reduced inversion altitude grid of 10-km resolution. Profiles for standard and wet (upper limit) atmospheric conditions are also shown.

and HDO when it is retrieved on the 4-km retrieval grid increases rapidly above ~ 20 km, which is insufficient if the isotopic fractionation above the condensation level is to be studied (see Fig. 7). However, retrieval simulations on a grid with reduced resolution of 10 km yield meaningful information on the D/H ratio up to an altitude of ~ 30 km for single-scan retrievals and up to ~ 35 km for zonal mean (25-s) retrievals. Obtaining information with better altitude resolution will require more extensive averaging.

The results of the retrieval simulations for H_2O_2 and O_3 are presented in Figs. 4 and 8. The estimated measurement precision for H_2O_2 , if it was retrieved on the standard 4-km retrieval grid given by the tangent altitudes of the measurement, was found to be approximately $5\text{--}10 \times 10^{-8}$ in the 0–60-km altitude range with a corresponding vertical resolution better than 5 km below 30 km. In the same band, an ozone profile can be simultaneously retrieved up to 70 km with retrieval precision and altitude resolution of the same order of magnitude as for H_2O_2 .

Given the expected mixing ratios of these species, useful information can thus be obtained only if data are averaged. Figure 4 shows therefore the results of single-scan retrievals on a grid with a reduced altitude resolution of 10 km. For H_2O_2 the resultant measurement error is smaller than 1×10^{-8} below 25 km, allowing useful information to be obtained up to an altitude of 20–30 km. Ozone may be retrieved up to ~ 50 km with a precision better than 3×10^{-8} . Figure 8 summarizes the theoretical measurement capabilities for single-profile and zonally averaged data products of H_2O_2 and O_3 . Note that for HDO, H_2O_2 , and O_3 the sensitivity of nadir sounding with an integration time of 5 s does not allow profiles to be retrieved.

B. Temperature Retrieval

Temperature information is contained in the limb measurements of optically thick lines such as ^{12}CO at 345.796 GHz and ^{13}CO at 330.588 GHz because the measured temperature is close to the atmospheric

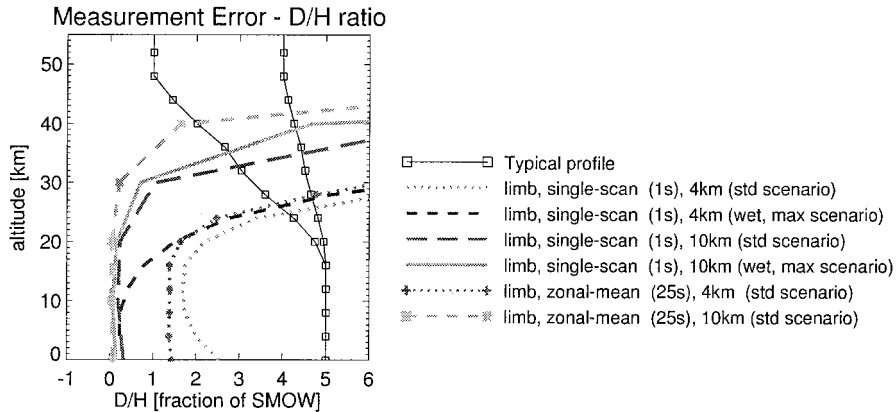


Fig. 7. Comparison of measurement precision for the D/H ratio derived from single-scan (1-s) and zonal-mean (25-s) retrievals of HDO at 335.1 GHz and H_2O at 325.1 GHz. Also shown are retrieval errors obtained with the assumption of a reduced inversion altitude grid of 10-km resolution. Profiles with depletion of 80% and 20%, derived from the standard and wet scenario for H_2O and HDO, are indicated for comparison.

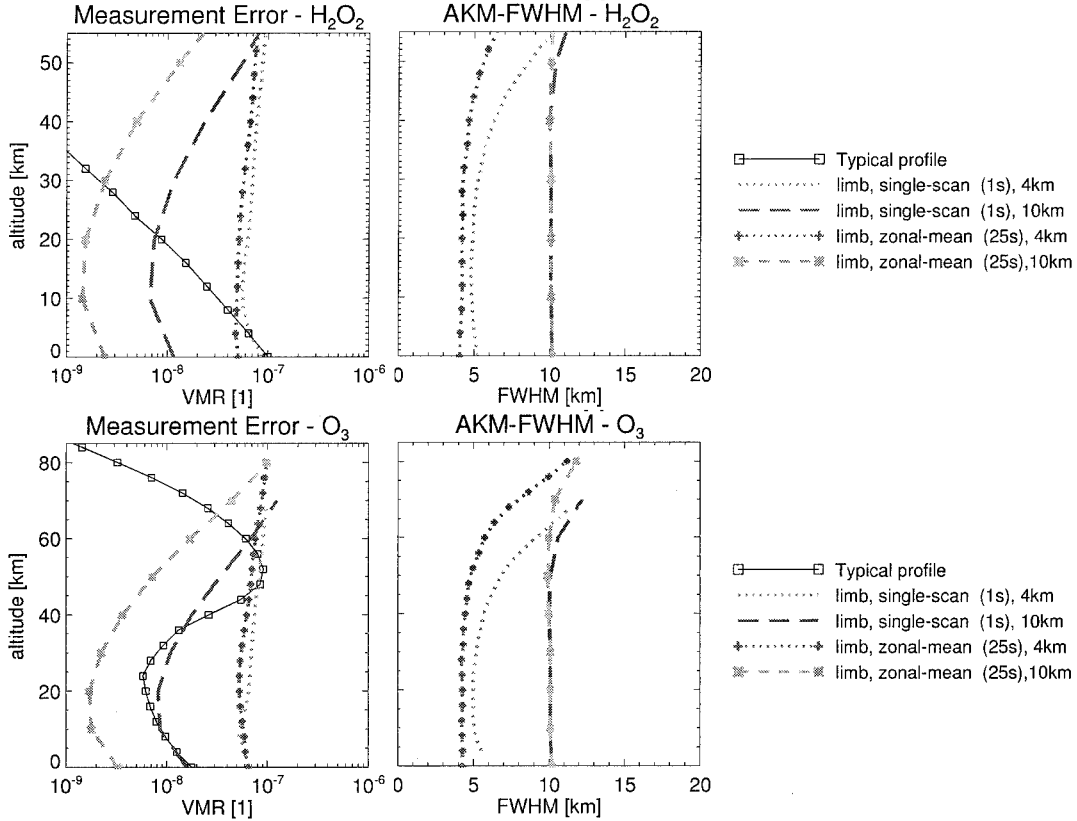


Fig. 8. Comparison of estimated measurement precision (left) and altitude resolution (right) achievable from single-scan (1-s), zonal-mean (25-s), and nadir (5-s) retrievals of H_2O_2 at 326.982 GHz (top) and O_3 at 326.901 GHz (bottom). Retrieval errors obtained from inversion simulations with the assumption of a reduced inversion altitude grid of 10-km resolution are also presented. Typical profiles are indicated for comparison.

temperature at the altitude where saturation occurs. This characteristic behavior is shown in Fig. 9, where brightness temperatures are plotted for both lines under consideration as a function of tangent altitudes. The temperature profile, converted to brightness temperatures, is also shown. It is evident that the maximum values of the spectra follow the atmospheric temperature profile up to altitudes of ~ 60 km for ^{13}CO and ~ 90 km for ^{12}CO . The combination of the information from both lines should consequently permit temperature to be retrieved up to ~ 90 km, as shown in the combined double-sideband spectrum of ^{12}CO and ^{13}CO at the right in Fig. 9.

Retrieval of temperature from these two long-lived molecules has the further advantage that the mixing ratio is relatively well known and much less variable compared with other species with strong lines such as water vapor. For the simulations we assume an *a priori* error of 20% (of the profile maximum) as a conservative worst-case estimate for the expected uncertainty in the ^{12}CO and ^{13}CO mixing ratios.

Nonlinear retrieval simulations with these assumptions resulted indeed in a significant measurement response range for temperature up to ~ 85 km, as shown in Fig. 10. For this example the altitude resolution is of the order of 4–5 km up to 55 km and in the range 5–10 km above. The precision is in the

range 3–6 K for temperature retrievals on the 4-km retrieval grid. Simulations that employed various different temperature profiles yielded similar results, and the upper altitude limit of the temperature retrieval was typically in the range 85–95 km.

Simultaneously, mixing ratio profiles for ^{12}CO and ^{13}CO were estimated. ^{12}CO can be retrieved up to 100 km with an error smaller than 1×10^{-4} (or 5–10%). The altitude resolution was better than 5 km below 40 km. Note that a minimum in the measurement response was observed near 80 km, which is obviously correlated with the temperature minimum. For this extreme example, the signal-to-noise ratio of the measurement in this altitude range is thus not sufficient to retrieve ^{12}CO along with temperature. For ^{13}CO , the approximate range with significant information from the measurement is 0–75 km, and the measurement precision is estimated to be better than 1×10^{-6} (5–10%). The FWHM value of the averaging kernel functions is smaller than 7 km for retrieval levels below 40 km.

Figure 11 compares the results obtained for temperature to nadir and zonal-mean retrievals. Zonal-mean temperature profiles (integration time 25 s) typically extend to approximately 90–100 km with an altitude resolution of 4–5 km below 55 km (increas-

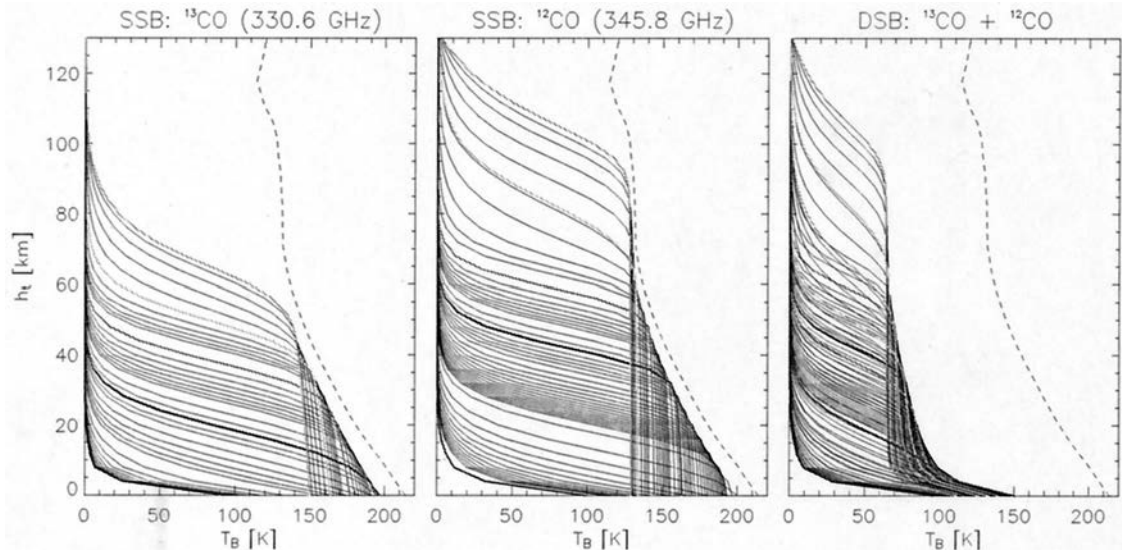


Fig. 9. MOLIÈRE-5 model calculations for selected channels of the ^{13}CO (left) and ^{12}CO (middle) bands, assuming standard atmospheric conditions. The curves represent brightness temperatures for individual channels as a function of tangent altitude (SSB: single-sideband spectrum). The resultant double-sideband spectrum (DSB) is shown at the right. The temperature profile, converted to brightness temperatures, is indicated as well (dashed curves).

ing to 10–15 km above) and a precision of the order of 1–2 K. Zonal-mean (25-s) retrievals on the 10-km grid would even allow information to be obtained up to ~ 120 km. Nadir observations (integration time, 5 s) would provide temperature information up to an altitude of ~ 70 km with a precision of 5 K but with considerably poorer vertical resolution that increases with altitude from ~ 6 to 16 km, which may cause considerable difficulties in reproducing the true profile above ~ 40 km. Information on the mixing ratios of the CO isotopes are not contained in a 5-s nadir measurement, which means in turn that the nadir temperature retrieval is relatively insensitive to variations in the CO mixing ratio.

An overview of the measurement capabilities of a state-of-the-art submillimeter-wave limb sounder operating in the 320–350-GHz range is provided in Table 2.

6. Alternatives for Line Selection

The search for a combination of target bands that contain suitable spectral lines of the most important target species H_2O and HDO as well as ^{12}CO and ^{13}CO for temperature and wind retrieval in one simple, compact, lightweight, and low-power-consuming instrument is the main driver for the selection of the frequency range of a potential submillimeter-wave limb sounder of the Martian atmosphere. The choice of the 320–350-GHz range is, as demonstrated above, one possible solution for the design of such an instrument. Alternatives for the selection of the frequency bands are now discussed.

^{12}CO as a linear molecule has a simple rotational spectrum with transitions at 115.3, 230.5, 345.8, 461.0, 576.3, and 691.5 GHz for the main isotope. Lines of the isotope ^{13}CO are always found in the vicinity of those of the main isotope at 110.2, 220.4,

330.6, 440.8, 550.9, and 661.1 GHz. We showed in Subsection 5.B that the combination of neighboring CO isotope lines with very different line strengths is highly beneficial to the retrieval of temperature. Strong water-vapor lines close to the CO isotope line pairs are available only at 183.3, 325.1, 439.2, 448.0, 474.7, 489.5, and 556.9 GHz, further restricting the choice. HDO transitions can be found in all the remaining windows: at 225.8, 335.4, 464.9, 479.9, and 481.7 GHz. Even stronger HDO lines could be measured at 490.6 or 599.9 GHz. Secondary-target species such as O_3 and H_2O_2 have more-complex spectra with many potential target lines in the submillimeter-wave range and are therefore not critical for the band selection. The millimeter- and submillimeter-wave spectra of O_2 , with only a few strong transitions at 118.8, 368.5, 424.8, and 487.3 GHz, would, however, influence the choice of an instrument's spectral range if scientific priority were given to measurements of this minor species.

Robust subharmonically pumped mixers, required for a simple design of the instrument optics, are nowadays available for all these spectral ranges. However, the receiver's sensitivity decreases considerably by roughly a factor of 2 from 300 to 600 GHz. The antenna's field of view, on the other hand, improves linearly with frequency, and a higher frequency range could therefore be beneficial for the altitude resolution.

In Figs. 12–14 we compare the theoretically achievable measurement precision and altitude resolution for the key parameters temperature, H_2O , and HDO when they are retrieved from single-scan limb observations of target lines in the 150–650-GHz range. To ensure the comparability of the spectral bands we based the retrieval simulations on the same assumptions about the instrumental characteristics that were previously adopted for the retrievals from the corre-

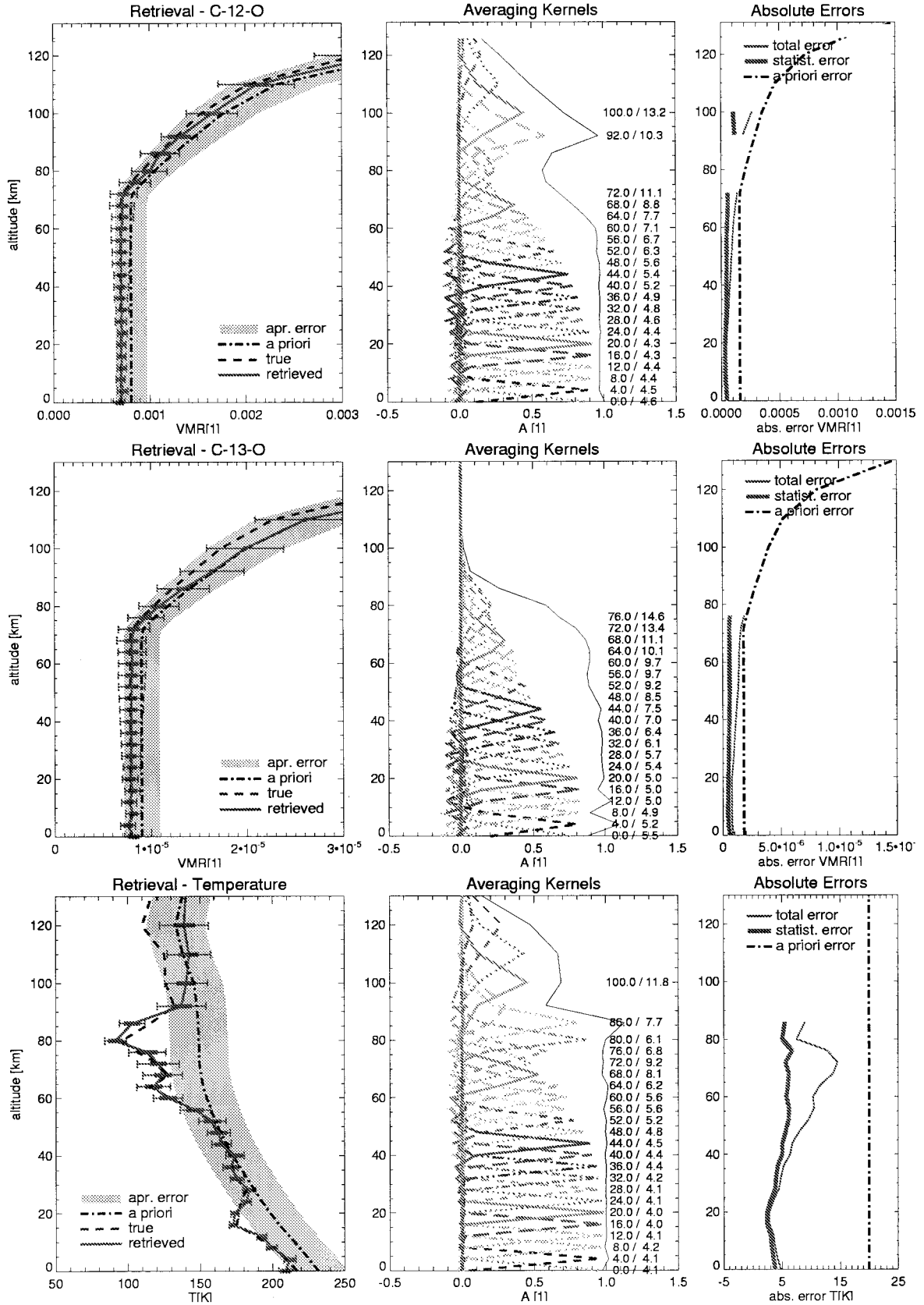


Fig. 10. Results of the nonlinear retrieval simulations for ^{12}CO at 345.796 GHz (top), ^{13}CO at 330.588 GHz (middle), and the simultaneously retrieved temperature. See the caption of Fig. 2 for an explanation.

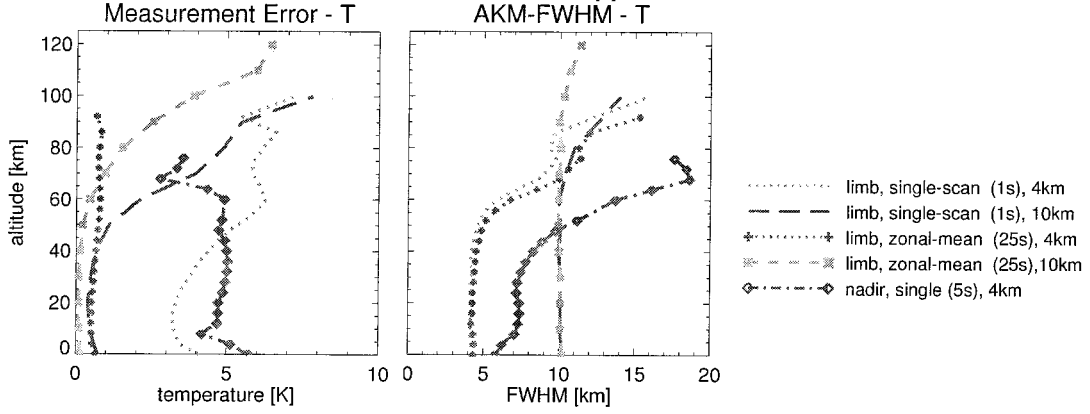


Fig. 11. Comparison of estimated measurement precision (left) and altitude resolution (right) achievable from single-scan (1-s), zonal-mean (25-s), and nadir (5-s) temperature retrievals from spectral lines of ^{12}CO at 345.796 GHz and ^{13}CO at 330.588 GHz. Also shown are retrieval errors obtained with the assumption of a reduced inversion altitude grid of 10-km resolution.

sponding target lines in the 320–350-GHz range. In particular, calculations were always performed for a double-sideband receiver and for the back-end spectrometer bandwidths and spectral resolutions listed in Table 1. For temperature retrieval from the CO line pairs the intermediate frequency was adjusted to allow for measurements of the ^{12}CO and ^{13}CO lines in the upper and lower sidebands in the same parts of the spectrometer as for the 330–345-GHz case. A linear dependence on frequency was assumed for the deterioration of the receiver noise as well as for the narrowing of the antenna’s field of view.

Figure 12 shows the results for the single-scan temperature retrieval from the four CO isotope line pairs that reside in the 150–650-GHz range. The measurement precision below ~ 50 km and the altitude resolution above ~ 60 km improve considerably for the higher-frequency bands; this improvement mainly reflects the strongly increasing spectroscopic line strengths with frequency. The narrowing of the antenna’s main lobe with increasing frequency might additionally play a role in the improvement of the altitude resolution above 60 km.

A comparison of retrieval results derived from

Table 2. Theoretical Capabilities of a Submillimeter-Wave Limb Sounder in the 320–350-GHz Range for Observation of Key Parameters in the Martian Atmosphere^a

Band	Species (Integration Time, Retrieval Grid Step)	ν_0 (GHz)	Altitude Resolution (km)	Altitude Range (km)	1- σ Precision (ppv or K)
A	^{12}CO (1 s)	345.796	≤ 5 (<40 km)	0–100	$4\text{--}10 \times 10^{-5}$ [5–10%]
	^{13}CO (1 s)	330.588	≤ 7 (<40 km)	0–75	$5\text{--}7 \times 10^{-7}$ [5–10%]
	T (1s)		4–5 (<55 km)	0–85	3–6 K
	T (25 s)		4–5 (<55 km)	0–110	1–2 K
	T (1s, 10 km)		10 (<80 km)	0–100	≤ 1 K (<50 km)
	T (25 s, 10 km)		10 (<100 km)	0–120	≤ 1 K (<70 km)
	T (5 s, nadir)		6–16	0–70	~ 5 K
B	H_2O	325.153	~ 4 (<70 km)	0–100	$\leq 1 \times 10^{-5}$ (<50 km)
	H_2O (25 s)		~ 4 (<85 km)	0–110	$\leq 3 \times 10^{-6}$ (<50 km)
	H_2O (1 s, 10 km)		10	0–100	$\leq 2 \times 10^{-6}$ (<60 km)
	H_2O (25 s, 10 km)		10	0–110	$\leq 1 \times 10^{-6}$ (<75 km)
	H_2O (5 s, nadir)		10–20	4–60	$4\text{--}6 \times 10^{-5}$
C	HDO (1 s)	335.396	~ 5 (<35 km)	0–60	$10\text{--}15 \times 10^{-8}$ (<40 km)
	HDO (25 s)		~ 4 (<35 km)	0–75	$8\text{--}15 \times 10^{-8}$ (<40 km)
	HDO (1s, 10 km)		10	0–60	$2\text{--}3 \times 10^{-8}$ (<40 km)
	HDO (25 s, 10 km)		10	0–70	$3\text{--}6 \times 10^{-9}$ (<40 km)
D	O_3 (1 s, 10 km)	326.901	10	0–70	$1\text{--}3 \times 10^{-8}$ (<50 km)
	O_3 (25 s, 10 km)		10	0–80	$2\text{--}6 \times 10^{-9}$ (<50 km)
	H_2O_2 (1 s, 10 km)	326.982	10	0–60	$\leq 1 \times 10^{-8}$ (<25 km)
	H_2O_2 (25 s, 10 km)		10	0–70	$\leq 2 \times 10^{-9}$ (<25 km)

^aShown are target line frequencies, altitude resolutions derived from the averaging kernel functions (FWHM), altitude ranges where the *a priori* contribution is negligible, and single-scan precision for 4- and 10-km step retrieval grids. Results obtained for 5-s nadir observations as well as for 25-s limb (zonal mean) retrievals are also provided for comparison.

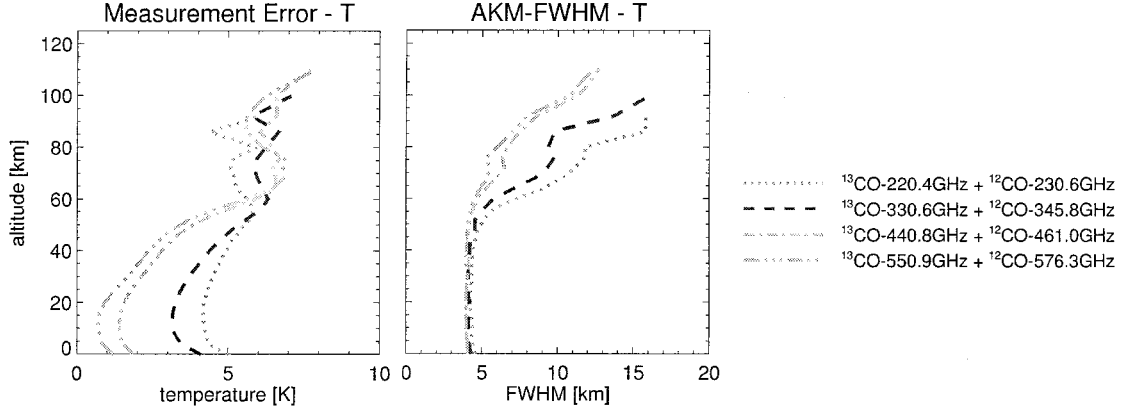


Fig. 12. Comparison of measurement precision (left) and altitude resolution (right) for single-scan profile retrievals of temperature from different CO isotope line pairs in the 150-650 GHz range.

H₂O target lines is shown in Fig. 13. Since H₂O in the Martian atmosphere is characterized by a significant seasonal variability, retrieval simulations were also calculated for extreme dry and wet conditions. The results show that, for the standard and wet conditions, all H₂O lines considered in the 150–650-GHz range are affected by self-absorption, leading to a considerable degradation of the retrieval precision at low tangent altitudes. In practice, self-absorption should limit the exploitable

altitude range, in particular for the strongest lines, because of the strong nonlinearity of the retrieval problem. Consequently, satisfactory results for the strongest water-vapor line at 556.9 GHz are obtained only at altitudes above 20 km or under extreme dry atmospheric conditions. Good sensitivity can be obtained by observation of the slightly weaker H₂O lines at 380.2, 448.0, or 620.7 GHz, still sufficient for the detection of the low mixing ratios of the dry profile below 20 km but, in the case of an

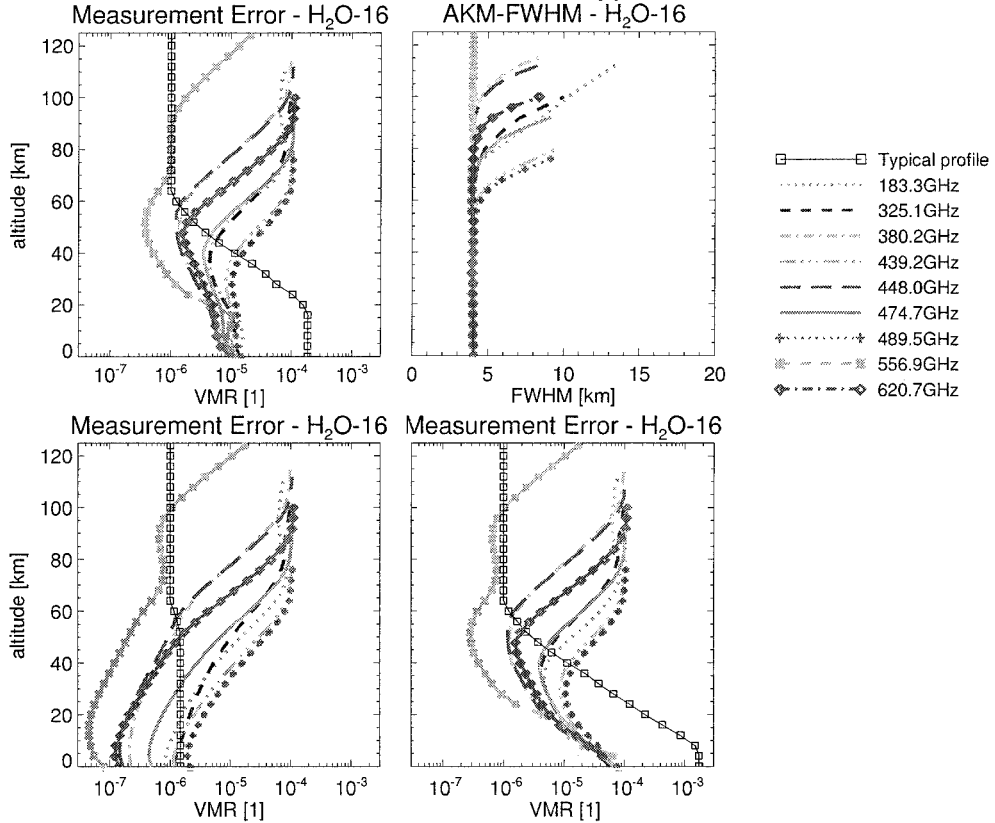


Fig. 13. Top, comparison of measurement precision (left) and altitude resolution (right) for single-scan profile retrievals from H₂O target lines in the 150–650-GHz range (standard conditions). Bottom, retrieval precision for extremely dry (left) and extremely wet (right) conditions. Typical profiles are given for comparison.

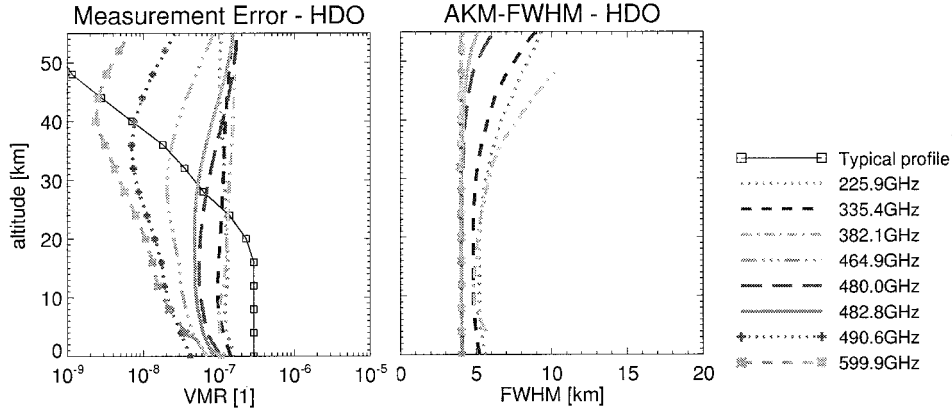


Fig. 14. Comparison of measurement precision (left) and altitude resolution (right) for single-scan profile retrievals from potential HDO target lines in the 150–650-GHz range. A typical profile is shown as well.

extremely wet atmosphere, affected by strong self-absorption at tangent altitudes below ~ 10 km. Spectral measurements of the 183.3-, 325.1-, and 474.7-GHz H_2O lines give similar results and allow single-scan profile information to be retrieved up to roughly 40–45 km for standard and wet conditions, but data have to be averaged to yield useful information on the extremely dry atmosphere.

The theoretically achievable precision and altitude resolution for the submillimeter wave transitions of HDO under investigation are compared in Fig. 14. Best results were found for the strongest lines at 490.6 and 599.9 GHz, allowing useful single-profile information to be retrieved up to 35–40 km, but self-absorption might here also be a limiting factor at the lowest tangent views below 10 km in an extremely wet atmosphere. In general, line intensities increase considerably with frequency, leading to an improved retrieval precision and exploitable altitude range for the higher-frequency lines of HDO. This result is also reflected in the measurement precision for the D/H ratio, which is shown in Fig. 15 for single-profile retrievals from different H_2O and HDO line pairs. The study of isotopic fractionation above the condensation level would cer-

tainly benefit if the strongest H_2O and HDO lines were measured.

To conclude, given the constraints on the design of a compact and low-mass instrument, the choice of the 435–465-GHz spectral range could represent a reasonable compromise and an alternative to the 320–350 GHz range. Main bands would cover lines of ^{13}CO and ^{12}CO at 440.8 and 461.0 GHz, respectively, as well as a strong H_2O transition at 448.0 GHz and a HDO line at 464.9 GHz. Strong transitions of the secondary target species O_3 and H_2O_2 are available in this range, and O_2 lines in the vicinity (424.8 and 487.3 GHz) could possibly be targeted as well. Simultaneous measurements of a slightly weaker H_2O line at 439.2 GHz could also help to improve the retrieval at the lowest tangent altitudes. The combination of different frequency ranges in one instrument would theoretically also be an option, for example, with the goal of observing stronger and weaker target lines of a species such as H_2O to maximize the measurement information at low and high altitudes but certainly at the cost of higher instrument design complexity, weight, and power consumption. Such technical constraints will have to be carefully investigated.

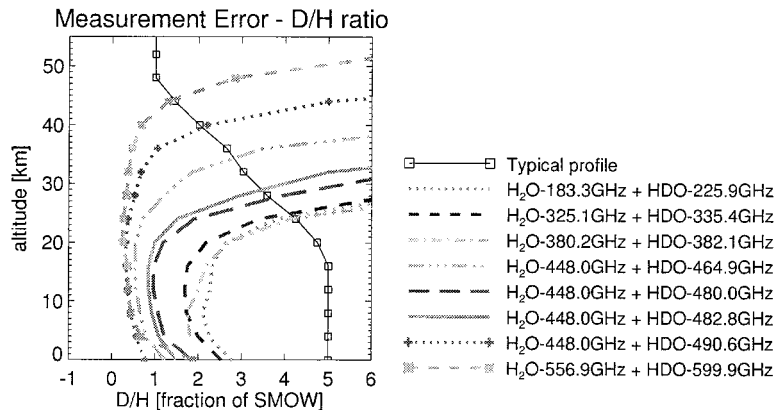


Fig. 15. Comparison of the measurement precision of the D/H ratio derived from single-scan profile retrievals from different H_2O and HDO line pairs in the 150–650-GHz range. The profile of the standard scenario (with 80% depletion above 45 km) is shown for comparison.

7. Discussion and Conclusions

Our investigation of submillimeter-wave heterodyne spectroscopy of the Martian atmosphere, based on a well-defined concept for the instrumentation, clearly demonstrates the feasibility of such measurements and thus confirms and extends conclusions drawn in earlier studies.^{31,49,50} In particular, we have provided quantitative estimates of achievable measurement precisions, altitude ranges, and vertical resolutions for the important parameters H_2O , HDO , ^{12}CO , ^{13}CO , temperature, O_3 , and H_2O_2 measured in the 320–350-GHz range with a state-of-the-art satellite receiver on a low Mars orbit.

Temperature is the main parameter that controls physical and chemical processes. Limb observations of lines of ^{12}CO (345.8 GHz) and ^{13}CO (330.6 GHz) would provide vertically well-resolved temperature information on a global scale during all seasons, during day and night, and even under extreme conditions such as dust storms. The altitude range from 0 to ~90 km covers the range where most of the atmospheric structures persist, and the vertical resolution of the order of one-half scale height allows wave structures such as gravity waves to be resolved. Zonal-mean temperature fields with a reduced altitude resolution of the order of 10–12 km would reach even altitudes of ~120 km.

Moreover, direct measurements of the slowly varying long-lived species ^{12}CO and its isotope ^{13}CO would become available, permitting, for example, slow meridional transport processes to be traced, in addition to direct observations of winds from the Doppler shifts of the simultaneously measured ^{12}CO and ^{13}CO lines. Note that a dedicated study that used a classic least-squares approach has shown that a rms measurement precision of better than 10 m/s from 20 to 100 km can be obtained from the spectral measurements of a single limb scan,⁶ to be compared with wind velocities that can be as high as 150–200 m/s in the middle atmosphere of Mars.¹³ Global observations based on a submillimeter-wave limb-sounder concept with dedicated bands in the 320–350-GHz range would thus allow a full description of the dynamics of the Martian atmosphere.

Single-scan limb observations of the 325.2-GHz H_2O line would allow global water-vapor fields to be measured from 0 to 40–45 km with unprecedented sensitivity, with half-scale height vertical resolution, and even under extremely wet (H_2O column, ~120 pr- μm) and moderately dry (~3 pr- μm) conditions, valuable information for the study of the diurnal, seasonal, and interannual exchange cycles with surface reservoirs such as the polar ice caps and the regolith. Averaging of data would provide information at even higher altitudes or under the extremely dry atmospheric conditions at the winter pole (~0.15 pr- μm), information that would still be useful to constrain chemical-dynamical models of the Martian atmosphere. Simultaneous HDO (335.4 GHz) limb observations would allow information on the D/H ratio in the lower Martian

atmosphere to be derived up to roughly 30 km, which would be particularly useful for the study of a variety of dynamic processes involving, for example, condensation.

It has been demonstrated that limb sounding provides much better altitude ranges, resolutions, and measurement sensitivities than nadir sounding. Moreover, global fields of photochemically important minor constituents such as O_3 (326.90 GHz) and the reservoir for odd hydrogen radicals H_2O_2 (326.98 GHz) can be obtained only by use of the much longer emission path of the limb observation geometry.

On the other hand, nadir observations with good horizontal resolution of better than 30 km along the satellite track, to be compared with the length of the limb emission path of ~330 km for a 4-km-thick layer at the tangent point, provide valuable complementary information, for example, on the detection of meteorological phenomena with strong horizontal gradients in temperature and humidity (e.g., weather fronts). Moreover, nadir observations of surface brightness temperatures in two polarizations, performed between limb scans, would permit mapping of surface characteristics on a global scale. Quantities such as the effective emission of the ground and the temperature profile in the lower atmosphere are in turn important for the initialization of nonlinear retrieval of atmospheric parameters from the limb observations.

Compared with the typical observation capabilities of present IR nadir sounders such as the Thermal Emission Spectrometer⁹ (TES) of the Mars Global Surveyor (MGS) mission or the Planetary Fourier-Transform Spectrometer¹⁹ (PFS) on Mars Express, a submillimeter-wave limb-sounder approach would provide a considerably better altitude coverage for temperature, profile measurements instead of total columns for H_2O and HDO , as well as profile measurements of ^{12}CO , ^{13}CO , O_3 , and H_2O_2 . Compared with solar and stellar limb occultation techniques such as those employed by SPICAM (Spectroscopy for the Investigation of the Characteristics of the Atmosphere of Mars¹⁸) on Mars Express, capable of providing, for example, accurate profile measurements of H_2O and O_3 , a major advantage of the concept investigated here is the much better temporal and global coverage of thermal emission measurements. The relative insensitivity to aerosols compared with UV-visible and IR techniques, as well as the validity of the LTE assumption even at high altitudes (in contrast to the thermal IR range), should also be recalled as general advantages of submillimeter-wave emission sounding.

Besides conducting a detailed investigation of the observation capabilities of a state-of-the-art limb sounder in the 320–350-GHz range, we also explored alternative frequency band selections and estimated retrieval precisions for the most important target lines of H_2O , HDO , ^{12}CO , and ^{13}CO in the 150–650-GHz range. For the design of a compact het-

erodyne receiver the 435–465-GHz spectral range represents a possible alternative that would allow somewhat stronger transitions of the key target species ^{12}CO , ^{13}CO , H_2O , and HDO to be observed. However, the technical feasibility of such an instrument would need to be analyzed in a more dedicated study, given the typically stringent requirements on dimensions, mass, and power consumption for spaceborne instrumentation. Complementarity with other instruments on the same platform would also be an important criterion for the target line selection of a future submillimeter limb sounder for the investigation of the Martian atmosphere.

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