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# The XMM-Newton View of GRS 1915+105<sup>★</sup>

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**Abstract.** Two XMM-Newton observations of the black-hole binary GRS 1915+105 were triggered in 2004 (April 17 and 21), during a long "plateau" state of the source. We analyzed the data collected with EPIC-pn in Timing and Burst modes, respectively; Reflection Grating Spectrometers were used only on April 21st. The source 2–10 keV flux is  $\sim 0.6$  (unabsorbed:  $0.9 \div 1.1$ )  $\times 10^{-8}$  in cgs units. While the light curves show just small amplitude variations (a few percent) at timescales longer than a few seconds, a QPO is seen at about 0.6 Hz – as expected in  $\chi$  variability modes of GRS 1915+105, when the phenomenological correlation with the source flux is taken into account – possibly with a harmonic signal at 1.2 Hz. The pn spectrum is well fitted without invoking thermal disk emission, on the base of four main components: a primary one (either a simple power law or thermal Comptonization models), absorbed by cold matter with abundances different than those of standard ISM; reprocessing from an ionized disk; emission and absorption lines; and a soft X-ray excess around  $\sim 1$  keV. However, the latter is not confirmed by the RGS spectra, whose difference from the EPIC-pn ones actually lacks of a fully satisfactory explanation. If real, the soft X-ray excess may be due to reflection from an optically thin, photoionized disk wind; in this case it may yield a way to disentangle intrinsic from interstellar absorption.

**Key words.** Black hole physics – Line: formation – Accretion, accretion disks – X-rays: binaries – X-rays: individuals: GRS 1915+105

## 1. Introduction

GRS 1915+105 is a well-known black-hole (BH) binary, sometimes classified as a *superluminal microquasar* because of the impressive radio jets, whose matter seems to move with superluminal velocity due to a relativistic effect (Mirabel & Rodriguez, 1994; Rodriguez & Mirabel, 1999; Fender et al., 1999). For a recent review on this source, including a general discussion on its properties, see Fender & Belloni (2004).

Due to very large obscuration, the spectral type of GRS 1915+105's companion (a K-M III star) was discovered only recently, via infrared observations (Greiner et al., 2001). This yielded evidence that the system, initially discovered in the  $\gamma$  rays (Castro-Tirado et al. 1992), belongs to the class

of low-mass X-ray binaries (LMXBs). Greiner et al. (2001), on the basis of the commonly adopted value of the source's inclination ( $70^\circ$ ; but see Maccarone, 2002, for a discussion), used the same observations to determine the mass of the central compact object, which has been constrained to  $M_c = 14 \pm 4 M_\odot$ , i.e. well above the maximum neutron star mass limit. GRS 1915+105 is therefore believed to host a BH with a gravitational radius  $r_g = \frac{GM}{c^2} \sim 21$  km, and thus classified as a BH binary.

Reig et al. (2003) noted that, although the phenomenology of GRS 1915+105 has been usually<sup>1</sup> described in terms of three spectral states, named A, B and C (the latter being low/hard and jet-dominated), in fact several source properties always correspond to the canonical *intermediate* (sometimes called *very high*) state of galactic BH sources (see Homan & Belloni, 2005, and Belloni et al., 2005, for a definition). This is confirmed by Done et al. (2004), who also note that the source's unique limit-cycle variability appears when the source radi-

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<sup>1</sup> On the base of the hardness ratio and position in the colour-colour diagram.

ates at super-Eddington luminosities. Many of GRS 1915+105 properties can be related to the binary’s evolutionary state: the giant companion star and the long orbital period (33.4 days, by far the longest of any LMXB (Greiner et al., 2001), indicate that the Roche lobe overflow occurs in a very wide binary. The cause of all the unique long as well as short-term variability of GRS 1915+105 is then linked to the evolution of the huge disk structure, which can contain enough material to maintain nearly-Eddington accretion rates over timescales of several years.

A strong and broad iron  $K\alpha$  line was present in *BeppoSAX* (Martocchia et al. 2002, 2004) and *ASCA* GIS (Miller et al. 2004) spectra of GRS 1915+105. *Rossini-XTE* PCA data are generally well-fitted including a broad iron line component (e.g. Done et al. 2004), but the energy resolution of this instrument does not allow to discriminate among different line models. Fe  $K\alpha$  emission was compatible also with the spectral fits of *Chandra* data, taken during a low/hard state of the source, as reported by Lee et al. (2002).

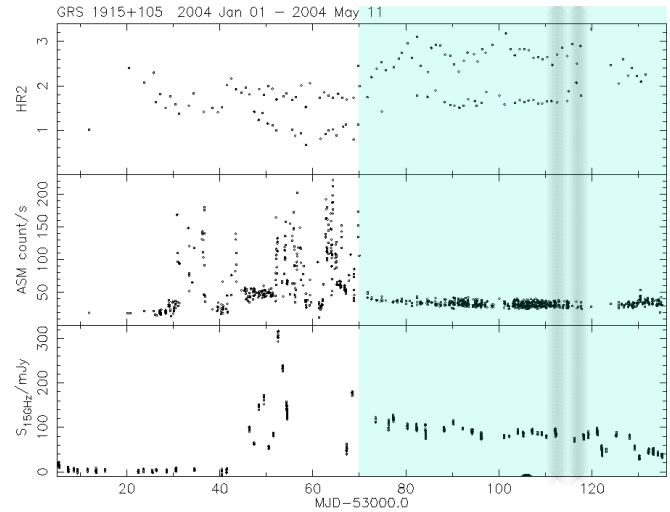
A XMM-Newton ToO observation of GRS 1915+105, which would allow to study the iron line profile with enhanced spectral resolution and sensitivity, was proposed in AO2. The observation was intended to be triggered by the occurrence of a “plateau” state of the source (Foster et al. 1996), similar to that observed during the *BeppoSAX* 1998 observation; this was necessary also in order to have the source in a less dramatic variability state, and at a lower flux level to minimize technical problems due to instrumental pile-up and telemetry. For the latter reason we also planned to observe in Timing Mode.

The observation was triggered in April 2004, divided into two parts with  $\sim 4$  days separation one from the other. In this paper, while presenting the results, we offer for the first time a XMM-Newton view of GRS 1915+105. First of all we describe (in Section 2) the observations and data extraction methods, and details of the state of the source (based on its variability and general considerations). In Section 3 the best-fit spectral models and related physical parameters are given. Finally, in the Conclusions, we summarize these XMM-Newton measurements and put our results in the context of the known source properties and open issues.

## 2. Observations and data reduction

To trigger the ToO observation, we had set a condition on variability (*rms* single-dwell points of *Rossini-XTE* ASM to stay  $< 9\%$  within one day) and a condition on flux (*Rossini-XTE* ASM daily rate below a threshold of 30 cts/s).<sup>2</sup> Both triggering conditions were met during the long “plateau”, a lower-flux state also characterized by a more intense radio emission (Foster et al. 1996), which lasted several weeks during the Spring 2004 (Fig. 1). The observation could be therefore started, in EPIC Timing Mode, on April 17, 2004 (hereafter: *OBS1*).

<sup>2</sup> ASM is the *All Sky Monitor* onboard *Rossini-XTE*. We used the public ASM “Weather Map” mask, accessible at: [http://heasarc.gsfc.nasa.gov/xte\\_weather/](http://heasarc.gsfc.nasa.gov/xte_weather/).



**Fig. 1.** Radio and X-ray light curves of GRS 1915+105 before and during the Spring 2004 “plateau” (starting at  $\sim 2453070$  JD), from: <http://www.mrao.cam.ac.uk/~guy/>. Radio data at 15 GHz were observed with the Ryle Telescope, Cambridge (see Pooley & Fender, 1997, for a description of the techniques used); X-ray data (ASM flux and hardness ratio) are taken from the *Rossini-XTE* quick-look web pages: [http://xte.mit.edu/ASM\\_lc.html](http://xte.mit.edu/ASM_lc.html). The approximate dates of *OBS1* and *OBS2* are shaded. Note the anti-correlation between X and radio.

However, the source counts were such to saturate the telemetry, even after switching off the MOS and RGS. Thus a second part of the observation was performed on April 21, 2004 (hereafter: *OBS2*), this time in (EPIC-pn) Burst Mode.

Given the typical “plateau” behaviour, with strong jet activity and high radio level, we conclude that the source was caught in the conventional “hard” (C) spectral state and  $\chi$  variability class, as defined by Belloni et al. (1997a,b, 2000), in both observations. This is similar to the condition in which the source was seen by simultaneous *INTEGRAL* (rev.57) and *Rossini-XTE* observations performed in the Spring 2003 (Hannikainen et al. 2003, 2004, 2005; see Fuchs et al. 2003 for a multi-wavelength study of one of these observations).<sup>3</sup>

Count rate thresholds for pileup in timing modes are reported in the XMM User’s Handbook:<sup>4</sup> the MOS count rate registered in both observations ( $\sim 130$  and  $160$  cts/s respectively) is well above the quoted limit (100), therefore EPIC-MOS data are unusable<sup>5</sup> and we restricted the analysis to EPIC-pn and RGS data (the latter are available only for *OBS2*).

Since, in this state, the source is stable, we can use the whole integration time to obtain very high statistics and signal/noise. Through adopting the customary method of extracting high-energy ( $E > 10$  keV) background light curves, we

<sup>3</sup> This radio-loud state had been described already by Munro et al. 2001, and by Trudolyubov 2001 (as “type II”).

<sup>4</sup> See Table 3 in: Ehle et al. (2004) – [http://xmm.vilspa.esa.es/external/xmm\\_user\\_support/documentation/uhb/node28.html#2828](http://xmm.vilspa.esa.es/external/xmm_user_support/documentation/uhb/node28.html#2828)

<sup>5</sup> Excising the core of the PSF is enough to cope with this problem in the case of imaging modes; this is not applicable for timing modes.

checked for intervals of low contamination in both observations. While during OBS1 no background flaring is apparent, initially we did select a 9 ks low-background interval (198951000 ÷ 198960000 s) in OBS2. We performed preliminary OBS2 data analysis with this selected dataset; afterwards, we verified that the results are not modified when considering the whole observing interval. Therefore, in the following, we report results of the data analysis which has been performed using *all* observing time, e.g. with the highest possible statistics. This allowed also to refine, as much as possible, the spectro-variability analysis (e.g. the analysis of root mean squared variability vs. energy: Ponti et al. 2004) which we performed to possibly disentangle the various contributions to the observed spectra, and thus to get an independent validation of the adopted emission/absorption model. The high number of counts allows a high-resolution analysis of *rms* vs. *E* over the whole EPIC-pn energy band (about 35 energy bins for OBS2, and several dozens of energy bins for OBS1; the minimum timescales of variation we could consider were, respectively,  $\sim 1000$  and  $\sim 100$  s). We used background-subtracted light curves in order to avoid contamination at the lowest and highest energies ( $E \lesssim 2$  and  $E \gtrsim 7$  keV), and found a mean *rms* value of the order of 5% ( $0.02 \div 0.07$ ) in both observations, compatible with a null *rms* when Poisson noise is removed, with no visible features in the “*rms* spectra”. We further verified this by studying  $F_{\text{var}}$  vs.  $E$ .<sup>6</sup>

The data have been reduced with the standard SAS v6.1 and FTOOLS software packages. For EPIC-pn, single and double events were used in the timing modes (PATTERN  $\leq 4$ ); the fit was restricted to energies  $> 0.5$  ( $> 0.4$ ) keV in Timing (Burst) mode in order to avoid the increased noise. As recommended by the XMM-Newton Helpdesk, no selection on the RAWY coordinate was applied (since this is related to a fine-time and the selection would therefore exclude certain time periods). However, RAWY  $< 160$  was imposed as a condition in Burst mode to avoid direct illumination by the source. Response and auxiliary matrices were created with the SAS tools *rmfgen*, *argen*, *rgsrmfgen*.

For the fit procedure we assumed a source distance of 12.5 kpc and inclination  $i \sim 70^\circ$  (Rodriguez & Mirabel, 1999; Fender et al., 1999).

### 2.1. First observation

The April 17 observation (OBS1: XMM-Newton rev.798, start time: 14:18:56 [JD2453113.096], stop time: 20:03:37 [JD2453113.336]) was performed in (both pn and MOS1) Timing Mode, as requested in the AO2 proposal mainly with the aim to avoid pile-up problems. Moreover, the MOS2 cam-

<sup>6</sup> The nominal time resolution is 29 and 7 microseconds for EPIC-pn Timing and Burst modes respectively. However, all results which are derived from the light curves must be taken with caution: in the first observation, even after removing all telemetry drops, the collected photons “saturated” the instrument’s bandwidth; therefore, energy-dependent oscillations in the count rate may have been “cut off”. On the other hand, photons collected in the second observation, i.e. in Burst Mode, are selected on the base of an *ad hoc* duty cycle in the reading process (events are not registered for 97% of the time).

era and both RGS had to be switched off, to avoid overloading the telemetry. Nevertheless, the source flux was still so high that telemetry drops were frequent even in pn data, resulting in short gaps in the light curves, lasting a few to several seconds, for a total of about 75% of the observing time. They cause the search for (quasi-) periodic oscillations to be senseless on  $\sim 1$  Hz timescales. Source light curves do not show any appreciable variations at timescales  $\sim a\ few$  to  $\sim 100$  s; however, as explained above (see Footnote 3), due to the full scientific buffer and telemetry drops, this result should be taken with caution.

Even though affected by telemetry drops for up to about 75% of the observing time, EPIC-pn data are well below the pile-up count rate threshold<sup>7</sup> and can be therefore safely used for spectral analysis.

The data were selected from a four-column strip centered on the source, all along the Timing-mode (one-dimensional) “image”; a similar strip was selected, far away from the source, as a background sample.

### 2.2. Second observation

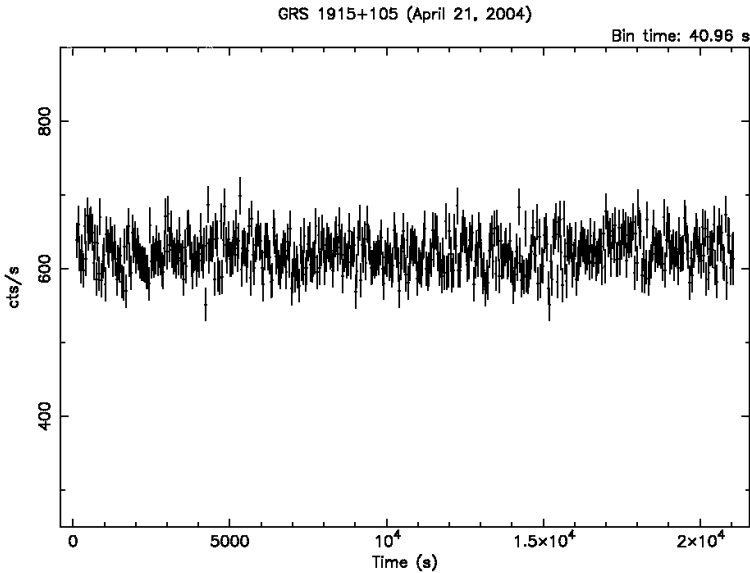
While the source essentially persisted in the same state, a second observation was started on April 21st (OBS2, rev. 800, start time: 13:49:23 [JD2453117.076], stop time: 20:56:55 [JD2453117.373]). To cope with the telemetry problem, this time we used EPIC-pn in Burst mode, MOS in Timing mode, and RGS in High Count Rate (HCR) mode. The other cameras were off.

In this observation, a background flare is apparent at the very beginning both in EPIC and in RGS data; however, due to the very high S/N and shortness of the flare, it does not really affect the spectra.

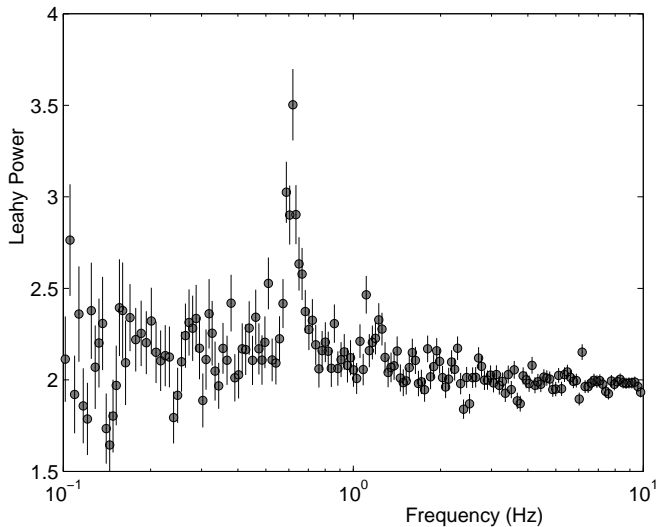
Source data were selected from a six-column strip (i.e. two columns more than in Timing mode, in order to increase the statistics), all along the Burst-mode (one-dimensional) “image”; a similar strip was selected, far away from the source, as a background sample.

EPIC-pn data are neither affected by telemetry drops nor by pile-up, in this observation, and can be therefore used for both spectral and timing analyses. However, as far as timing/variability studies are concerned, some caution is necessary since Burst-mode data suffer from the intrinsic photon loss related to the 3% duty cycle (cp. Footnote 6). While the light curves show only tiny variations (up to a few percent, but compatible with zero: see Fig. 2) at all timescales bigger than a few seconds, a Quasi Periodic Oscillation (QPO) at about 0.6 Hz, with a possibly harmonic signal at 1.2 Hz is seen in the EPIC-pn data collected in Burst mode (Fig. 3). This frequency is expected in  $\chi$  variability modes of GRS 1915+105, when the phenomenological correlation with the source flux is taken into account (c.p. Fender & Belloni 2004). A forest of narrow peaks at  $\nu > 100$  Hz is apparent, due to spurious signals of instrumental origin.

<sup>7</sup> For EPIC-pn Timing mode the pile-up threshold for a point source is 1500 cts/s, the telemetry threshold is 300 cts/s, while we registered  $\sim 650$  cts/s.



**Fig. 2.** The source light curve in the 0.4 – 10.0 keV band during OBS2 (EPIC-pn, Burst mode, *not* background-subtracted). The whole 20946 s observation has a mean 618 (i.e. 18.54/0.03) cts/s; the background has just about 18 (i.e. 0.5320/0.03) cts/s.

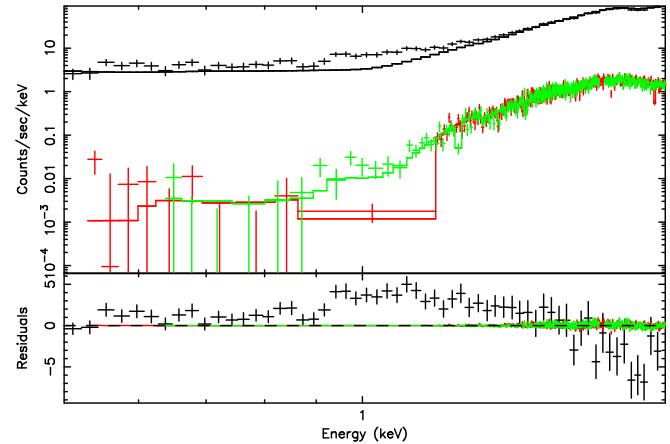


**Fig. 3.** Power spectrum in the 0.4 – 13.0 keV band of the OBS2 (April 21) EPIC-pn data.

### 2.2.1. The RGS data

Thanks to their excellent energy resolution, RGS cameras may in principle reveal subtle details in the 0.3–2.0 keV band. However, the source is so strongly absorbed below  $1 \div 2$  keV that these data do not actually add much information. We used them mainly to check the EPIC-pn spectrum: in Fig. 4 we show the 0.5–2 keV spectra of both instruments, fitted with a simple absorbed power law model, with  $\Gamma$  fixed to 2 (see Section 3.2).<sup>8</sup>

<sup>8</sup> Multiplicative constants to the RGS have been also added, to allow for intercalibration problems (the best fit values for the constants



**Fig. 4.** The RGS and pn (Burst mode) spectra fitted in the 0.5–2 keV range with a simple absorbed power law. The RGS and pn spectra clearly disagree each other.

The RGS and EPIC-pn spectra are clearly different, especially with regard to the broad excess around 1 keV in the latter spectrum. The RGS spectrum shows a fast decline (more than two orders of magnitude, with very few counts at the lower energies), without any apparent features. In Section 3.1.5 we will go through the alternative hypotheses which can be invoked to explain the RGS–pn discrepancy.

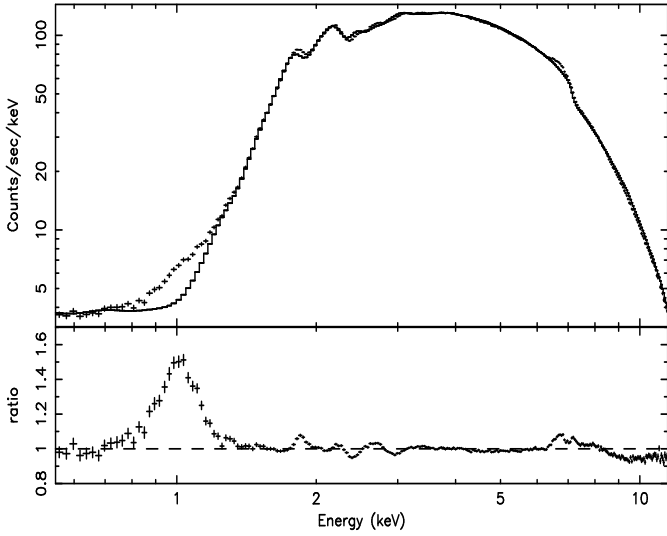
RGS normally have 15 ms timing resolution in HTR mode; we produced standard light curves, but did not get any useful results due to the low statistics on most of the band.

## 3. Spectral analysis

At the flux level of GRS 1915+105 ( $\sim 10^{-8}$  erg cm<sup>-2</sup> s<sup>-1</sup>) it is very likely that systematic errors cannot be always neglected with respect to the statistical ones. However, introducing an energy-independent systematic uncertainty would prevent a meaningful use of the  $\chi^2$  statistics. Features at the 1–2% level must be therefore taken with caution, being possibly of instrumental origin (cp. Kirsch 2005).

Even though the background issue is not relevant on most of the energy band, due to the extremely high S/N, nevertheless it may play a role at the highest energies ( $E \gtrsim 10$  keV) as well as at the energies where the source is strongly absorbed ( $E \lesssim 1.5$  keV). We therefore performed background subtraction; all reported results refer to background-subtracted analysis, unless stated otherwise. Of course, if a dust halo is present (a hypothesis we cannot check with our data, given the lack of imaging capabilities of timing modes; but see Greiner et al., 1998, on the issue) background subtraction may cause a spectral distortion in the soft band, depending on the distance from the source at which the background spectrum is taken (see below, Section 3.1.5).

are 1.28 and 1.22 for the RGS1 and RGS2, respectively). RGS data, collected in the standard CCD read-out scheme, are marginally affected by pile-up, which amounts to a few percent and mainly affects the energies where the spectrum peaks.



**Fig. 5.** The OBS1 spectrum and data/model ratio when fitted with a simple power law and cold absorption with variable abundances.

Spectra were rebinned in order to have about three bins per energy resolution element, and at least 20 counts per bin at the same time. Spectral fits were performed with the XSPEC v.11.3.0 software package. In the following, all errors refer to 90% confidence level for one interesting parameter ( $\Delta\chi^2=2.71$ ).

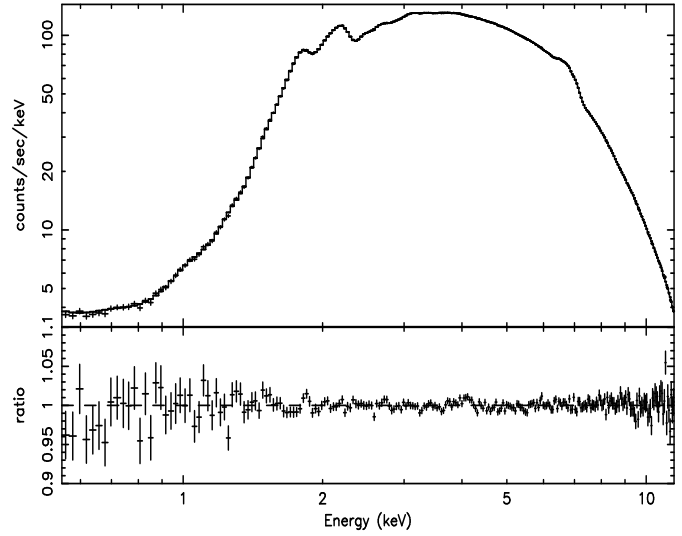
### 3.1. The Timing-mode observation

As discussed in the previous Section, we are confident that the telemetry problems do not affect significantly the spectrum. Given the much larger total number of counts in OBS1, we decided to use it first for the spectral analysis.

To illustrate the spectral complexity, in Fig. 5 we show the best fit spectrum and data/model ratio after fitting with a simple power law absorbed by cold matter. Element abundances were left free to vary with respect to the solar value ratios. Several features are apparent: an excess around 1 keV, several wiggles between  $\sim 1.5$  and 3 keV, an excess at the energies of the iron line complex, and a deficit above 8 keV.

After several attempts, we found a – both statistically ( $\chi^2/\text{d.o.f.}=317.5/227$ ) and physically – reasonable description of the spectrum in terms of (see Table 1, and Fig. 6): a power law with  $\Gamma \sim 1.7$ ; several absorption and emission features; an ionized, optically thick reflection component (and/or a warm absorber); and an optically thin reflection component to model the 1 keV excess. Given the fact that all remaining deviations of the data from the model are no larger than about 1-2%, i.e. of the order of the systematic errors in the calibration (which, for most of the energy range under analysis, is larger than the statistical errors), we did not try to further improve the fit.

We will now discuss the various spectral components separately for the sake of clarity.



**Fig. 6.** The OBS1 spectrum and data/model ratio for the best fit model (see text and Table 1).

#### 3.1.1. The continuum

First of all, we note that a simple power law is basically sufficient to describe the continuum, apart from an additional component required to fit the  $>8$  keV deficit, and an apparent, broad excess around  $\sim 1$  keV. Substituting the power law with simple Comptonization models (like [Comptonization](#) or [Comptonization](#)) does not lead to any significant improvement in the fit. No thermal emission from an accretion disk is apparent (the inclusion of a multi-color disk component is not required by the data). Of course, this strongly suggests that the innermost disk regions were absent during the observation. As far as the apparent 1 keV excess is concerned, it is far too narrow to be fitted by even a single temperature blackbody (i.e. the narrower among physically motivated continuum models) or to be residual from “distorted” disk thermal emission. The nature of this feature will be discussed separately in Section 3.1.5.

#### 3.1.2. The cold absorber

The photoelectric absorption has been fitted with the XSPEC model [tbabs](#), which assumes neutral matter and allows to vary the elemental abundancies separately. In practice, for the sake of simplicity, we grouped the elements, so that within each group the elements are linked one another. The groups are: H, He, C, N, O; Ne, Na; Mg, Al; Si; S; Cl, Ar, Ca, Cr; Fe, Co, Ni. The rationale behind this choice is both physical and practical. In fact, we grouped together elements whose origin is likely common, but also those which are less abundant (and therefore cannot be easily measured independently one from the other) with very abundant ones: e.g. Co and Ni with Fe. We also grouped H and He with C, N and O because their effect is small in our energy range.

In Table 1 the best fit values are reported. They are given in units of Hydrogen equivalent column densities for solar abundances as in Anders & Grevesse (1989). A significant overabundance of the heavier elements with respect to the lighter

ones is apparent. These results are only partly consistent with those derived from the *Chandra* HETG observation (Lee et al. 2002). In particular, Fe and S abundances are roughly the same, while our values for Mg are about twice, and 2/3 for Si. Fixing the abundances to the values of Lee et al. (2002), the fit is unacceptable ( $\chi^2/\text{d.o.f.}=659.9/231$ ). We note that our observation has a lower energy resolution, but many more photons.

These overabundancies with respect to the solar values (the differences are even greater assuming the ISM values, see Wilms et al. 2000) suggest that a significant fraction of the absorber is local to the source. The value for the low Z elements (i.e.  $1.6 \times 10^{22} \text{ cm}^{-2}$ ) can then be assumed as the upper limit to the ISM absorber column density. It is interesting to note that this value is similar to that estimated by fitting the  $\sim 1$  keV excess (see below).

### 3.1.3. Reflection or warm absorption?

The deficit above 8 keV can be fitted equally well either by a simple smeared edge model ( ), or by an ionized and relativistically blurred reflection component ( ). As the smearing edge is not physically meaningful *per se*, but is usually assumed to represent the reflection component, we will not further discuss the fits with this component.

The model adopted for the disk reflection component includes the relativistic effects in Schwarzschild metric. The inner and outer radii are linked to those of the H-like iron line (see below). The disk inclination angle has been fixed to  $70^\circ$ , and the matter temperature to  $10^6$  K. The very large value of the inner radius (about  $320r_g$ ) justifies *a posteriori* the use of the Schwarzschild metric instead of the Kerr one (differences between the two metrics, even for a maximally rotating BH, are insignificant above  $\sim 10r_g$ ), and agrees with the lack of any detectable thermal disk emission. The iron abundance is five times the solar value, roughly consistent with the overabundance found in the cold absorber (see above). The ionization parameter is pretty large, in agreement with the presence of He- and H-like iron lines. The significant solid angle subtended by the disk to the illuminating source, i.e.  $0.3 \times 2\pi$ , implies that the emission occurs at distances comparable to the inner disk radius. The adopted model is necessarily a simplification of the real situation, as it adopts a unique ionization for the whole disk, while the ionization likely decreases with increasing radius. The derived parameters, therefore, must be taken with caution and considered only as indicative.

The presence of a H-like iron absorption line (see below) suggests warm absorption. Substituting the reflection component with a warm absorber (model ), a significantly worse, but not completely unacceptable fit is found ( $\chi^2/\text{d.o.f.}=384.5/226$ ), with a comparable iron overabundance ( $A_{\text{Fe}}=6.6$ ) and power law photon index ( $\Gamma \sim 1.69$ ). However, the column density ( $\sim 7.7 \times 10^{21} \text{ cm}^{-2}$ ) is larger than that implied by the value of the EW of the absorption line (see Table 1), given the large iron overabundance (e.g. Bianchi et al. 2005). Moreover, the presence of ionized disk lines (see below) *requires* the presence of the ionized reflection component. We thus fitted the spectrum with both the reflection and the warm

absorber. No significant improvement is found with respect to the fit with the reflection only, but the upper limit to the column density,  $\sim 2 \times 10^{21} \text{ cm}^{-2}$ , is now consistent with the absorption line EW.

### 3.1.4. The iron lines

In Fig. 5, a clear broad excess at the energy of the iron line complex is apparent. The excess is at energies larger than 6.4 keV, strongly suggesting ionized iron. We fitted this excess with two unresolved emission lines ( $\sigma=10$  eV) peaking at 6.7 keV and 6.96 keV, which correspond to He- and H-like iron, respectively, and (after Lee et al. 2002) we also included an unresolved absorption line with energy free to vary around the value of the H-like iron. The fit is unacceptable ( $\chi^2/\text{d.o.f.}=617.2/225$ ); not surprisingly, the absorption line is not found. A much better, but still hardly acceptable, fit ( $\chi^2/\text{d.o.f.}=425.6/223$ ) is found when leaving the widths of the lines free to vary. We found  $\sigma=129$  eV and 174 eV for the He- and H-like lines, respectively; an absorption line centered at  $\sim 6.95$  keV is also found. The large values of  $\sigma$ , along with the likely presence of the disk reflection component, strongly suggest that the lines are relativistically broadened: we therefore substituted the broad gaussians with models, with outer radius linked to that of the reflection component, as well as the inner radius of the H-like line (the inner radius of the He-like left free to vary, in the assumption that the matter is more ionized in the innermost regions). The results are reported in Table 1. The fit is significantly better than with the gaussians and, as expected on both physical grounds and the fact that the width of the He-like line is smaller, the inner radius for this line is larger than that of the H-like line (the values of the radius must of course be taken merely as indicative, see discussion above). The energy of the narrow absorption line,  $6.95^{+0.01}_{-0.03}$  keV, is just smaller than the rest frame value.<sup>9</sup> The line was assumed to be narrow ( $\sigma \equiv 1$  eV); leaving  $\sigma$  free to vary, no improvement in the fit is found.

We also tried to fit the iron emission with just one disk line model, with the energy free to vary. The fit is of comparable statistical quality, but the line energy,  $6.86 \pm 0.01$  keV, does not correspond to any rest frame atomic energy of significance. Moreover, the good fit is obtained at the expense of a very large value for the iron abundance of the reflection component ( $\sim 18$  times the solar value). Therefore, the two disk lines model has to be preferred.

The EW of both lines (28 eV) seems rather small when compared to the amount of reflection component (Matt et al. 1996). However, one must note again that the ionization likely changes along the radius in the disk. While reflection occurs (even if with changing spectra) all along the disk, line emission occurs only if the corresponding ion is present, which likely happens only in a fraction of the disk.

<sup>9</sup> Actually, this is a blend of two lines, with centroid energy at 6.966 keV.



### 3.1.5. The $\sim 1$ keV excess

In Section 2.2.1 we have shown that, during the Burst-mode observation, the RGS and pn disagreed, with the former not showing the 1 keV excess. (Unfortunately, during the Timing-mode observation the RGS were switched off). There are several possible explanations for both the excess and the RGS-pn discrepancy. As we will see, none of them is fully satisfactory, and we consider the issue still open, waiting for improvements in calibrations.

A pn instrumental origin for the broad excess seems unlikely, for various reasons. The amplitude (about 40% of the underlying continuum, see the data vs. model ratio in Fig. 5) is far too large, given the known calibration uncertainties (cp. Kirsch 2005). Moreover, the excess is present in the Timing as well as in the Burst mode, and was observed also with *BeppoSAX* (cp. Fig. 3 in: Feroci et al. 1998). A similar feature has been seen also in other binaries (see e.g. the case of the dipping LMXB 4U 1323-62: Boirin et al. 2005), but not in all; in particular, it is not present in the Burst mode observation of the Crab Nebula (Kirsch et al. 2005).

However, given the unusual brightness of our source, and the fact that the excess arises in the part of the spectrum where the count rate drops precipitously due to absorption, a calibration problem, possibly related to the redistribution matrix, cannot be completely excluded at the moment.

Let us then assume that this excess is real. We could not fit this feature with any continuum model put behind the whole absorber. As we have seen in Section 3.1.1, it cannot be fitted with blackbody models (the only model able to fit it is a broad gaussian with an EW of several tens of keV, which has no physical explanation). Good fits are instead found if the emission is screened by a thinner absorber. There are three possibilities in this respect: (i) a confusing source; (ii) a dust scattering halo; (iii) a component related to the system, but lying outside the intrinsic absorber.

The confusing source hypothesis seems unlikely. In timing modes, source data are extracted from a narrow strip (here about  $16''$ ); the required, “polluting” source would have to be bright (0.5-2 keV observed flux of about  $4 \times 10^{-12}$  erg cm $^{-2}$  s $^{-1}$ ): but no source is present in the *ROSAT* All Sky Survey within a radius of  $0.3^\circ$  around GRS 1915+105 (we checked with the HEASARC tool PIMMS that such a source would be well above the typical RASS detection limit; of course, the possibility of a variable source cannot be ruled out).

Since the dust scattering is energy dependent, ignoring a possible halo component leads to a distortion of the spectrum in a small energy band, because of an incorrect background subtraction (depending on the distance from the source at which the background spectrum is taken); however, less soft photons would be “observed”. A possibility is that we are including the halo in the extraction region. Since there is no obvious dust scattering model to fit the data, imaging may be the only way to test the halo hypothesis: given the particular modes adopted, this sort of analysis is very difficult (and altogether impossible in the Burst mode). We only note that the dust scattering halo observed in Cyg X-2 (Costantini et al. 2005) has quite a broad

spectrum (when normalized to the source spectrum), broader than the feature under discussion. Another problem with the scattering halo hypothesis is that the relative contribution of the excess is similar in the Timing and Burst mode observations, which is possible only if most of the halo is within  $16''$ , i.e. the width of the strip from which the Timing mode spectrum was extracted.

Even if unsatisfactory in many respects, the abovementioned two hypotheses have the merit to invoke extended or off-axis emission, which can in principle help explaining the presence of the excess in the pn spectra but not in the RGS. However, as the RGS spectrum is extracted over a region of about 2.5 arcminutes, the halo hypothesis seems to be excluded for the reasons above discussed. A confusing source, provided it lies at a greater distance, is instead still tenable. We note that all *Chandra* observations of GRS 1915+105 have been performed in the CC (i.e. non imaging) mode, therefore the abovementioned hypotheses could not be tested yet.

In the assumption that it is the RGS spectrum which suffers from instrumental/calibration problems, the 1 keV component could be related to the system, and a possible candidate would be the emission by a disk wind. Both a collisionally ionized plasma model and a photoionized plasma model do indeed fit to the EPIC-pn data well.

In the first case, the temperature of the plasma is about 1.3 keV (obtained using the `tbabs` model), with metal abundance 1.2 times the solar value; the ISM absorber has a  $N_{\text{H}}$  of  $9 \times 10^{21}$  cm $^{-2}$ . The emission integral is  $7.2 \times 10^{57} d_{10}^2$  cm $^{-3}$  (where  $d_{10}$  is the distance to the source in units of 10 kpc), which, assuming that the density of the matter is constant along the emitting region, translates into a mass of the wind  $M_w \approx 6 M_{\odot} d_{10}^2 n^{-1}$ , where  $n$  is the number density. Assuming that the size of the disk wind is at most of the order of the binary separation (about  $5 \times 10^{10}$  cm according the binary system parameters given by Harlaftis & Greiner 2004), we obtain a lower limit to  $n$  of about  $8 \times 10^{12} d_{10}$  cm $^{-3}$ . With this size, the Thomson optical depth would already be substantial ( $\sim 0.25$ ), and further increasing with decreasing size. Thus, the matter would not be really optically thin, and this solution is therefore not self-consistent.

Alternatively, the matter may be in photoionization equilibrium, the emission being due to reflection of the primary X-ray radiation. To model this case, we adopted a power law with the same photon index as the primary one (to model Thomson scattered radiation) plus a gaussian line. In this case, the ISM  $N_{\text{H}}$  is larger,  $\sim 1.5 \times 10^{22}$  cm $^{-2}$  (this value may be assumed as an estimate of the ISM column density)<sup>10</sup>. From the observed flux, the optical depth of the wind is about  $0.001/f$ , where  $f$  is the covering factor of the wind to the primary radiation. The centroid energy of the gaussian line ( $\sim 0.97$  keV), its width (90 eV), and its EW against the reflected continuum (5.6 keV), suggest a blend of Ne K and Fe L lines.

<sup>10</sup> For the sake of simplicity we fitted the absorption with a pure gas photoelectric absorption (model `tbabs` in XSPEC). Using a more appropriate ISM model, which includes the effects of dust (TB `tbabs` in XSPEC, Wilms et al. 2000), does not change the results significantly. The fit returns basically the same  $\chi^2$  and the same column density.



### 3.1.6. Other emission/absorption lines

Other emission and absorption lines are required to fit the spectrum. They are: two unresolved emission lines at about 1.85 keV and 2.25 keV, with EW=19 eV and 8 eV, respectively; two broad absorption lines at about 2.5 and 2.93 keV, with  $\sigma$  of 0.11 and 0.14 keV and EWs of  $-17$  and  $-30$  eV, respectively.

Some of these features may be instrumental artifacts, since this energy range is notoriously troublesome due to the changes in the effective area.<sup>11</sup> However, some of them appear too prominent to be entirely due to calibration problems. The first emission line can be readily identified with the Si recombination line. If arising from the same photoionized material responsible for the  $\sim 1$  keV excess (see above), its total EW would be about 16 keV, a very large but not fully unrealistic value, given the large overabundance we found in the local absorber and provided that the reflecting matter is still optically thin to resonant scattering (Matt et al. 1996), which is possible if the matter is largely turbulent. The same is true for the second emission line, which can be identified with Si and S, even if the best fit value of the energy is slightly lower.

More difficult to explain are the two broad absorption lines. The energies correspond to the  $K\alpha$  and  $K\beta$  recombination lines from S. In this case, the absorbing material cannot be the same as that responsible for the Fe absorption line (see Section 3.1.4), since the latter is more ionized. It is also worth recalling that the modeling of the absorption edges is oversimplified, because it does not take into account X-ray Absorption Fine Structures (XAFS) which are generated if the material is in grains (e.g. Lee et al. 2002); this may introduce spurious features in the residuals, even if it cannot fully explain them. Moreover, we stress again that it is not inconceivable that a significant part of these features be of instrumental origin. Data with higher energy resolution and optimized calibration are necessary to solve this issue.

### 3.2. The Burst-mode observation

We fitted the Burst-mode spectrum with the same model used for the Timing-mode spectrum, fixing those parameters which are not expected to vary. In particular, we fixed the interstellar absorption column density, the ratio between the elemental abundances of the local absorber, and the iron abundance of the reflection component at the best fit values obtained for the Timing-mode observation. A good fit ( $\chi^2/\text{d.o.f.}=248.2/219$ ) is found. The best fit parameters are summarized in Table 2.

In comparison with the Timing-mode observation (see Fig. 7), the spectrum is steeper ( $\Gamma \sim 2$  instead of  $\sim 1.7$ ), the reflection component much stronger ( $R \sim 1.7$  instead of 0.3) and more ionized ( $\xi \sim 3300$  instead of  $940 \text{ erg cm s}^{-1}$ ). The inner and outer radii are much smaller ( $r_i < 20$ ,  $r_o \sim 100$ ). The local absorber is slightly larger, while the ionized iron disk lines are not detected (upper limits to the EW of about 10 eV for both the He- and H-like lines. The Fe H-like absorption line

is still present, with a centroid energy of  $6.98 \pm 0.02$ , i.e. consistent with both the rest frame value and the value obtained in the Timing-mode observation. The EW is larger, about  $-50$  eV, due to the fact that the line is now best fitted with a resolved width  $\sigma$  of  $0.11 \pm 0.03$  keV. The 1.85 and 2.24 keV emission lines are not required, with upper limits to the EWs of 5 and 2 eV, respectively. A 6.4 keV iron emission line is instead marginally found, with an EW of  $6 \pm 4$  eV.

The above-mentioned results are consistent with a picture in which the disk is now more extended downwards and ionized (in fact, so much ionized that no significant relativistic iron line is observed). The steeper spectrum can be due to the more efficient cooling due to the increase of soft disk photons. The fact that no thermal disk emission is required, however, implies that the values for the inner and outer radii must be taken with caution. It must be noted that they are now determined only via the reflection component, with all the associated uncertainties (see discussion in the previous Section), as no disk line is present.

We must however remark that the abovementioned results (especially those regarding the reflection components) can be at least partly an artifact due to calibration problems. A deficit of photons above  $6 \div 7$  keV, i.e. similar to what we find (see Fig. 7) is indeed present also in the Burst mode observation of the Crab Nebula (Kirsch et al. 2005), a source where of course no significant reflection component is expected.

Regarding the 1 keV excess, which still persists, it can be again well fitted by a power law plus an emission line, the latter with a flux slightly larger than that obtained in the Timing-mode observation. In the reflection scenario discussed in the previous section, the optical depth is about a factor 10 larger, which implies a much lower EW of the line. These findings can be explained, at least qualitatively, assuming that the matter is now thicker to the lines, with consequent self-absorption effects (see e.g. Matt et al. 1996).

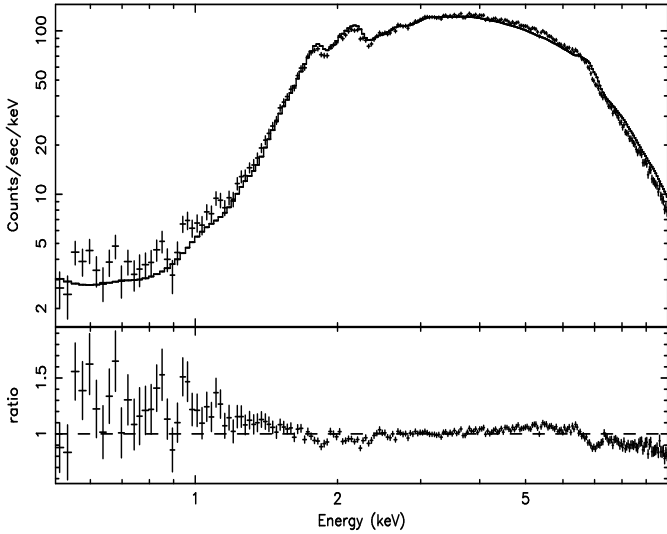
## 4. Conclusions

In this paper for the first time we have presented XMM-Newton observations of GRS 1915+105. Several problems make it difficult to observe the source fruitfully with this satellite, first of all due to the extremely high flux – in our case:  $F_{2-10 \text{ keV}} \sim 6.0 \div 6.6$  (unabsorbed:  $\sim 8.7 \div 10.7$ )  $\times 10^{-9} \text{ erg cm}^{-2} \text{ s}^{-1}$  – and to the notorious, erratic source variability, which does not help the triggering of observations. We succeeded at both a) observing the source in a well-defined, stable physical/spectral state and b) collecting EPIC-pn useful data, only marginally corrupted by telemetry problems.

In both (April 17th and 21st, 2004) observations the source has been caught in a “plateau” state, also characterized by stronger jet activity and higher radio level, which we identify with the conventional “C” state /  $\chi$  variability mode as defined by Belloni et al. (1997a,b, 2000). In this state the source shows a QPO at  $\sim 0.6$  Hz, i.e. what is expected in “plateau” intervals when the frequency vs. spectral hardness correlation is taken into account.

In order to try disentangle the different spectral components, and thus to better discriminate among the viable spectral/physical models, we used the *rms* vs. *E* method by Ponti et

<sup>11</sup> See Fig.30 and 31 in the XMM-Newton User’s Handbook: [http://xmm.vilspa.esa.es/external/xmm\\_user\\_support/documentation/uhb/node32.html](http://xmm.vilspa.esa.es/external/xmm_user_support/documentation/uhb/node32.html).



**Fig. 7.** The spectrum and data/model ratio of the Burst-mode observation with the best fit model for the Timing-mode observation. Clearly, iron line emission is over-subtracted, and the deficit at higher energies results to be more important than in OBS1; note again the strong excess at  $\sim 1$  keV and the wiggles at  $\sim 2$ – $3$  keV (cp. with Fig. 5).

al. (2004). The resulting *rms* is lower than 0.1 all over the energy band. We analyzed further the variable  $F_{\text{var}}$ , which takes the Poisson noise into account, and concluded that this variable is compatible with being null on the whole energy band, i.e. all spectral components are compatible with being constant during the two observations.

A priori, some of the features in the spectrum may be affected by dust halo scattering (although this cannot explain the 1 keV excess: see Section 3.1.5). However, testing this hypothesis would first of all require an accurate analysis based on proper source imaging (if available); further, spectral modelling of such effects is not easy (see e.g. Costantini et al. 2005). Being beyond the scope of this paper, this study should be addressed in future works.

We adopted a power law continuum model, which mimics emission by a hot corona. Such a continuum could be in principle due to Comptonized thermal emission e.g. from the jet basis (as proposed by Rodriguez et al., 2004); however, an optically thick reflector is required to account for the smeared edge at  $\sim 7$  keV (with a covering ratio of  $\sim 0.4 \div 1.7 \times 2\pi$  in the two observations respectively). This may be evidence of an accretion disk being present, or just optically thick, only at quite large distance from the central compact object, at least in the first observation ( $r_i/r_g > 300$  in OBS1,  $\sim 20$  in OBS2). The relatively large amount of the reflection components implies that the primary X-ray emitting region should have a size at least comparable to the inner disk radius. That the disk is truncated, i.e. not present in the innermost part, is suggested also by the non-detection of a thermal disk emission component.

Several line residuals are superimposed on the modeled continuum. Part of these may be due to calibration uncertainties, especially at the energies where changes in the EPIC effective area take place (e.g. 1–3 keV). However, we found clear

evidence of ionized iron emission around  $\sim 7$  keV: data are well fitted with two ionized Fe  $K\alpha$  lines, possibly affected by mild relativistic broadening (being produced far away from the BH event horizon), plus a narrow absorption feature at  $\sim 6.95$  keV.

The puzzling presence of an intense, broad excess around 1 keV can be explained (assuming, as a working hypothesis, that it is real, and that the discrepancies between the RGS and pn spectra discussed in sections 2.2.1 and 3.1.5 have to be attributed to RGS calibration problems, an admittedly unsubstantiated assumption at the moment) in terms of reflection by an optically thin wind. The excess is indeed well fitted with a power law plus a line, unobscured by material intrinsic to the system. The centroid energy of the gaussian line ( $\sim 0.97$  keV), its width (90 eV), and its EW against the reflected continuum (5.6 keV), point to a blend of Ne K and Fe L lines. The value of the equivalent H column density is interestingly similar to the value of the obscuration by low Z elements (H, He, C, N, O) at the source core –  $N_{\text{H}} \sim 1.6 \times 10^{22} \text{ cm}^{-2}$ : in the disk wind hypothesis, this may therefore be taken as an upper limit to the *interstellar* matter column density. This value matches well with the expected galactic absorption in that direction (Dickey & Lockman 1990).

On the other hand, a significant fraction of the absorber must be *local* to the source. We adopted a variable absorption model (in XSPEC), assuming neutral matter and grouping the elements on the base of both physical and practical considerations: elements which have probably a common origin, but also elements which are not very abundant (and therefore cannot be easily measured independently one from the other) with very abundant ones (e.g. Co and Ni with Fe): see the results in Table 1. A significant overabundance of the heavier elements with respect to the lighter ones is apparent, which suggests that a significant fraction of the absorber, traced by heavier species, is local to the source. These results are somewhat different from those derived from *Chandra* HETG data (Lee et al. 2002): in particular, Fe and S abundances are roughly the same, while our values for Mg are about twice, and 2/3 for Si. The overall scenario is, however, very similar.

Clearly, the intrinsic absorption may be subject to substantial changes on longer timescales, as already observed with *Rossi-XTE* in correspondence of similar “plateaux” (class  $\chi$ , state C): from a mean value of  $\sim 4.5 \times 10^{22} \text{ cm}^{-2}$ , the total  $N_{\text{H}}$  may occasionally rise up to  $\sim 6 \times 10^{22} \text{ cm}^{-2}$  (Belloni et al. 2000).

Assuming the customary value of  $\sim 12.5$  kpc for the source’s distance (Rodriguez & Mirabel, 1999) we obtain an estimate of the intrinsic luminosity  $L_{2-10 \text{ keV}}^{\text{unabs}} = 1.7 \div 2.0 \times 10^{38} \text{ erg s}^{-1}$  in the two observations respectively.<sup>12</sup>

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<sup>12</sup> Some authors claim that the distance could be substantially smaller (by up to a factor  $\sim 1/2$ : see Chapuis & Corbel 2004), in which case a corresponding correction of the luminosity (and BH mass) estimate should be considered.

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## References

- Anders, E., & Grevesse, N. 1989, *Geochim. et Cosm. Acta* 53, 197
- Belloni, T., Mendez, M., King, A.R., van der Klis, M., & van Paradijs, J. 1997a, *ApJ* 479, L145
- Belloni, T., Mendez, M., King, A.R., van der Klis, M., & van Paradijs, J. 1997b, *ApJ* 488, L109
- Belloni, T., Klein-Wolt, M., Mendez, M., van der Klis, M., & van Paradijs, J. 2000, *A&A* 355, 271
- Belloni, T., Homan, J., Casella, P., van der Klis, M., Nespoli, E., Lewin, W.H.G., Miller, J.M., Mendez, M. 2005, *A&A*, in press
- Bianchi, S., Matt G., Nicastro F., Porquet D., Dubau J., 2005, *MNRAS*, 357, 599
- Boirin, L., Mendez, M., Diaz Trigo, M., Parmar, A.N., Kaastra, J.S. 2005, *A&A* 436, 195
- Castro-Tirado, A.J., Brandt, S., & Lund, N. 1992, *IAU Circular* n.5590
- Chapuis, C., & Corbel, S. 2004, *A&A* 414, 659
- Costantini, E., Freyberg, M.J. & Predhel, P. 2005, *A&A* – in press: astro-ph/0508129
- Dickey, J.M., & Lockman, F.J. 1990, *ARA&A* 28, 215
- Done, C., Wardzinski, G., Gierlinski, M. 2004, *MNRAS* 349, 393
- Ehle, M., Breitfellner, M., González Riestra, R., Guainazzi, M., et al. 2004, "XMM-Newton Users' Handbook" Issue 2.2, online edition – [http://xmm.vilspa.esa.es/external/xmm\\_user\\_support/documentation/uhb/index.html](http://xmm.vilspa.esa.es/external/xmm_user_support/documentation/uhb/index.html)
- Fender, R.P., et al. 1999, *MNRAS* 304, 865
- Fender, R., & Belloni, T. 2004, *ARA&A* 42, 317
- Feroci, M., Costa, E., Matt, G., Belloni, T., Tavani, M. 1998, *Nucl. Phys. B (Proc. Suppl.)* 69/1-3, 336
- Foster, R.S., Waltman, E.B., Tavani, M., Harmon, B.A., Zhang, S.N., Paciasas, W.S., Ghigo, F.D. 1996, *ApJ* 467, L81
- Fuchs, Y., Rodriguez, J., Mirabel, I.F. et al. 2003, *A&A* 409, L35
- Greiner, J., Morgan, E.H., & Remillard, R.A. 1998, *New Astr. Rev.* 42, 597
- Greiner, J., et al. 2001, *A&A* 373, L37
- Greiner, J., Cuby, J.G., & McCaughrean, M.J., 2001, *Nature* 414, 522
- Hannikainen, D.C., Vilhu, O., Rodriguez, J., et al. 2003, *A&A* 411, L415
- Hannikainen, D.C., et al. 2004, in: Proc. of the 5th *INTEGRAL* Workshop, Munich 16-20 February 2004. Published by EDP – in press: astro-ph/0405349
- Hannikainen, D.C., Rodriguez, J., Vilhu, O., et al. 2005, *A&A* – in press: astro-ph/0502438
- Harlaftis, E.T., & Greiner, J. 2004, *A&A* 414, L13
- Homan, J., & Belloni, T. 2005, in: Proc. of the Workshop "From X-ray Binaries to Quasars: Black Hole Accretion on All Mass Scales", Eds. T.J. Maccarone, R.P. Fender, L.C. Ho (Dordrecht: Kluwer), in press – in press: astro-ph/0412597
- Kirsch, M. 2005, "EPIC Status of Calibration and Data Analysis", XMM-SOC-CAL-TN-0018 – [http://xmm.vilspa.esa.es/external/xmm\\_sw\\_cal/calib/documentation/index.shtml#EPIC](http://xmm.vilspa.esa.es/external/xmm_sw_cal/calib/documentation/index.shtml#EPIC)
- Kirsch, M., et al. 2005, "Crab: the standard X-ray candle with all (modern) X-ray satellites" – astro-ph/0508235
- Lee, J.C., et al. 2002, *ApJ* 567, 1102
- Maccarone, T.J., 2002, *MNRAS* 336, 1371
- Martocchia, A., Matt, G., Karas, V., Belloni, T., Feroci, M. 2002, *A&A* 387, 215
- Martocchia, A., Matt, G., Karas, V., Belloni, T., Feroci, M. 2004, in: Proc. II *BeppoSAX* Symposium, *The restless high-energy universe*, Amsterdam (The Netherlands), 5-8/5/2003. Eds. E.P.J. Van den Heuvel, J.J.M. In't Zand and R.A.M.J. Wijers, Elsevier, *Nucl. Physics B (Proc. Suppl.)* 132, 404
- Matt, G., Brandt, W.N., Fabian, A.C. 1996, *MNRAS*, 280, 823
- Miller, J.M., Fabian, A.C., Nowak, M.A., Lewin, W.H.G., 2004, in: Proc. 10th Annual Marcel Grossmann Meeting on General Relativity, in press – astro-ph/0402101
- Mirabel, I.F., & Rodriguez, L.F. 1994, *Nature* 371, 46
- Muno, M., Remillard, R., Morgan, E. et al. 2001, *ApJ* 556, 515
- Ponti, G., Cappi, M., Dadina, M., Malaguti, G. 2004, *A&A* 417, 451
- Pooley, G., & Fender, R., 1997, *MNRAS* 292, 925
- Reig, P., Belloni, T., van der Klis, M. (2003), *A&A* 412, 229
- Rodriguez, L.F., & Mirabel, I.F. 1999, *ApJ* 511, 398
- Rodriguez, J., Hannikainen, D.C., Shaw, S.E., Cabanac, C., Fuchs, Y. 2004, in: Proc. of the "Semaine de l'Astrophysique Francais", Scientifics Highlights 2004. Eds. F. Combes, D. Barret, T. Contini, F. Meynadier, L. Pagani – astro-ph/0409561
- Trudolyubov, S.P. 2001, *ApJ* 558, 276
- Wilms, J., Allen, A., McCray, R. 2000, *ApJ* 542, 914

|                          |   |
|--------------------------|---|
| $\Gamma$                 | <b>Power law</b><br>$1.686^{+0.008}_{-0.012}$                                 |
|                          | <b>Cold absorption</b>  |
| $N_{\text{H,He,C,N,O}}$  | $1.60^{+0.17}_{-0.29} \times 10^{22} \text{ cm}^{-2}$                         |
| $N_{\text{Ne,Na}}$       | $7.46^{+0.31}_{-0.72} \times 10^{22} \text{ cm}^{-2}$                         |
| $N_{\text{Mg,Al}}$       | $7.57^{+0.54}_{-0.16} \times 10^{22} \text{ cm}^{-2}$                         |
| $N_{\text{Si}}$          | $5.70^{+0.07}_{-0.12} \times 10^{22} \text{ cm}^{-2}$                         |
| $N_{\text{S}}$           | $4.69^{+0.07}_{-0.69} \times 10^{22} \text{ cm}^{-2}$                         |
| $N_{\text{Cl,Ar,Ca,Cr}}$ | $11.0^{+1.3}_{-1.6} \times 10^{22} \text{ cm}^{-2}$                           |
| $N_{\text{Fe,Co,Ni}}$    | $11.7^{+0.2}_{-0.2} \times 10^{22} \text{ cm}^{-2}$                           |
|                          | <b>Emis. lines</b>  |
| $E_l$                    | $1.846^{+0.006}_{-0.005} \text{ keV}$   |
| $F_l$                    | $1.35^{+0.15}_{-0.07} \times 10^{-2} \text{ ph cm}^{-2} \text{ s}^{-1}$       |
| $EW$                     | 19 eV   |
| $E_l$                    | $2.244^{+0.007}_{-0.010} \text{ keV}$   |
| $F_l$                    | $3.62^{+0.46}_{-0.62} \times 10^{-3} \text{ ph cm}^{-2} \text{ s}^{-1}$       |
| $EW$                     | 8 eV  |
|                          | <b>Abs. line</b>  |
| $E_l$                    | $6.95^{+0.01}_{-0.03} \text{ keV}$  |
| $F_l$                    | $-7.9^{+0.2}_{-0.1} \times 10^{-4} \text{ ph cm}^{-2} \text{ s}^{-1}$         |
| $EW$                     | -9 eV   |
| $\sigma$                 | 1 eV (frozen)   |
|                          | <b>1st disk line (He-like Fe)</b>   |
| $E_c$                    | 6.7 keV (fixed)   |
| $r_i/r_g$                | $580^{+210}_{-120}$   |
| $F_l$                    | $2.37^{+0.18}_{-0.18} \times 10^{-3} \text{ ph cm}^{-2} \text{ s}^{-1}$       |
| $EW$                     | 28 eV   |
|                          | <b>2nd disk line (H-like Fe)</b>  |
| $E_c$                    | 6.96 keV (fixed)  |
| $r_i/r_g$                | $320^{+80}_{-60}$   |
| $F_l$                    | $2.20^{+0.19}_{-0.21} \times 10^{-3} \text{ ph cm}^{-2} \text{ s}^{-1}$       |
| $EW$                     | 28 eV   |
|                          | <b>Reflection</b>   |
| $R/2\pi$                 | $0.35^{+0.02}_{-0.02}$  |
| $A_{\text{Fe}}$          | $5.2^{+0.7}_{-1.9}$   |
| $\xi$                    | $940^{+190}_{-80} \text{ erg cm s}^{-1}$                                      |
| $r_i/r_g$                | $320^{+80}_{-60}$   |
| $F_{2-10 \text{ keV}}$   | $\sim 0.6$ (unabs: $\sim 0.87$ ) $10^{-8} \text{ erg cm}^{-2} \text{ s}^{-1}$ |
| $\chi^2/\text{d.o.f.}$   | 317.5/227   |

**Table 1.** The best fit model parameters for the Timing-mode observations. The column densities are the Hydrogen equivalent column for solar abundances as in Anders & Grevesse (1989). The outer radii of the reflection and disk line models have been linked together in the fit, yielding  $r_o/r_g = 900^{+190}_{-90}$ ;  $r_i$  in the 2nd disk line model has been linked together to  $r_i$  of the reflection component. See the text for further discussion.

|                         |   |
|-------------------------|---|
| $\Gamma$                | <b>Power law</b><br>$2.04^{+0.01}_{-0.02}$                                    |
|                         | <b>Cold absorption</b>  |
| $N_{\text{H,He,C,N,O}}$ | $1.98^{+0.02}_{-0.02} \times 10^{22} \text{ cm}^{-2}$                         |
|                         | <b>Reflection</b>   |
| $R/2\pi$                | $1.69^{+0.16}_{-0.04}$  |
| $r_i/r_g$               | <20   |
| $r_o/r_g$               | $93^{+11}_{-17}$  |
| $\xi$                   | $3300^{+600}_{-600} \text{ erg cm s}^{-1}$                                    |
| $F_{2-10 \text{ keV}}$  | $\sim 6.6$ (unabs: $\sim 10.7$ ) $10^{-9} \text{ erg cm}^{-2} \text{ s}^{-1}$ |
| $\chi^2/\text{d.o.f.}$  | 248.2/219   |

**Table 2.** The best fit model parameters for the Burst-mode observation (OBS2). See the text (Section 3.2) for details.